

General Disclaimer

One or more of the Following Statements may affect this Document

- This document has been reproduced from the best copy furnished by the organizational source. It is being released in the interest of making available as much information as possible.
- This document may contain data, which exceeds the sheet parameters. It was furnished in this condition by the organizational source and is the best copy available.
- This document may contain tone-on-tone or color graphs, charts and/or pictures, which have been reproduced in black and white.
- This document is paginated as submitted by the original source.
- Portions of this document are not fully legible due to the historical nature of some of the material. However, it is the best reproduction available from the original submission.

NASA CR-
150987

LUNAR SEISMIC PROFILING EXPERIMENT
NATURAL ACTIVITY STUDY

FINAL REPORT

Project: NAS-9-14405

14 Sept., 1976

Prepared by F.K. Duennebier*

THE MARINE SCIENCE INSTITUTE
700 THE STRAND
GALVESTON, TX. 77550

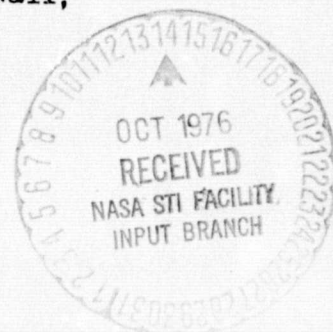
(NASA-CR-150987) LUNAR SEISMIC PROFILING
EXPERIMENT NATURAL ACTIVITY STUDY Final
Report (Marine Science Inst.) 46 p HC \$4.00
CSSL 03B

N76-33115

Unclas
05310

G3/91

*Now at the Hawaii Institute of Geophysics, University of Hawaii,
Honolulu, HI. 96822.



I. Introduction.

The Lunar Seismic Experiment Natural Activity Study has provided a unique opportunity to study the high frequency (4-20 Hz) portion of the seismic spectrum on the moon. The purpose of this project was to study the data obtained from the LSPE and determine as much as possible about 1) the origin and importance of the process that generates thermal moonquakes; 2) the characteristics of the seismic scattering zone at the lunar surface; and 3) the meteoroid flux for masses between 0.1 and about 10 grams.

The first problem has received the most effort and is the subject of the paper enclosed with this report (Appendix A). The detection of thermal moonquakes by the LSPE array made it possible to locate the sources of many events and determine that they are definitely not generated by astronaut activities but are the result of a natural process on the moon.

The second problem, that of the propagation of seismic waves in the near-surface layers, was studied in a qualitative manner. In the absence of an adequate theoretical model for the propagation of seismic waves in the moon, it is not possible to assign a depth for the scattering layer. The LSPE data does define several parameters which must be satisfied by any model developed in the future.

The third problem, that of meteoroid flux, was not studied. The vast numbers of events observed are almost certainly about 99% thermal moonquakes and, unlike the Apollo 14 PSE case (Duennebier and Sutton, 1974), there was no obvious way to discriminate between

ORIGINAL PAGE IS
OF POOR QUALITY

thermal moonquakes and meteoroid impacts. At station 14, thermal moonquakes always had a lower frequency spectrum than near-by impacts. At the LSPE array, thermal moonquakes are found with all ranges of spectra.

In addition to the above problems, a search was made for events detected by the PSE that were also detected by the LSPE. During the times when both instruments were operating, only one event was detected that was visible on both systems. This event, a fairly large meteoroid impact, was just barely visible on the LSPE records although easily visible at all PSE stations. The reason for the difference in amplitude is the different frequency responses of the two systems, the LSPE being less sensitive to frequencies below 4 Hz than the PSE. No high frequency teleseismic events (HFT) (Nakamura, et al., in prep) were observed during the times of LSPE operation. HFT events should be clearly recorded on the LSPE and the PSE. The LSPE is now being operated solely for the purpose of studying events that are recorded on both the LSPE and PSE network.

II. Data Collection.

During the short periods of time during which the LSPE was activated for the active portion of the experiment, several natural seismic events were noticed (Kovach, et al., 1973). In the hopes that these events were caused by thermal moonquakes of the type observed at the PSE stations (Duennebier and Sutton, 1974), the LSPE was activated for a 4-day period during the summer of 1973

ORIGINAL PAGE IS
OF POOR QUALITY.

to see if more of this activity could be observed. The results of this experiment showed that large numbers of thermal moonquakes were being detected by the LSPE. Because of the possibility of locating the source of many of these events using array techniques, it was decided to collect at least one full lunar month (lunation) of LSPE data. The data rates used by the LSPE require that all other experiments at the Apollo 17 site lose data when the LSPE is activated. To prevent long periods of data loss to other experiments, the LSPE was activated for 4-day periods at different parts of the lunation for a period of about one year until ten 4-day listening periods had been obtained. In this way data covering one full lunation (plus two periods of overlap) were obtained. In addition, a 4-hour period spanning the occurrence of a solar eclipse was recorded. For the more than 40 days of coverage, less than 5% of the data was lost due to poor transmission from the moon since the listening periods were planned by JSC personnel to coincide with times of optimal transmission.

Upon receipt of the data, about one full-reel tape per 4-hour period, compressed time scale playouts were produced (Figure 1) using the University of Texas PDP-15 computer. These records were scanned by eye and events above a given size (7 mm) were recorded. A listing of these events and copies of the compressed playouts were sent to R.L. Kovach at Stanford for his use as soon as they became available. Once events of interest were chosen, a set of event tapes was generated containing only periods where events occurred. These tapes were then used to generate expanded time scale playouts (Figures 2A and 2B of Appendix A) and power spectral analyses (Figure 2) of each event.

III. Analysis.

Thermal moonquakes.

The most time-consuming operation was the hand correlation of the more than one thousand events recorded. A property of thermal moonquakes is that their source mechanism is so repeatable that the same events are seen at about the same time of each lunation. Some types of events occur several times per lunation. When two events were found to have identical or nearly identical wave forms, they were assigned a number corresponding approximately to the longitude of the sun at the time of observation. Events generated by the LM were also identified. A complete listing of events observed is contained as Appendix B of this report. Of the 1730 events listed, 324 are caused by thermal noises in the LM, and 359 are thermal moonquakes with at least two events with nearly identical character observed. It is believed that almost all of the remaining events are thermal moonquakes for which no matching event was found either because of low signal-to-noise ratio or because the period of recording was not long enough to record a matching event. The reader is referred to Appendix A for further discussion as to the methods used in the study of thermal moonquakes at the LSPE site.

Seismic Propagation and Scattering.

The eight explosive package (EP) signals and the LM impact signal were analyzed to obtain information concerning the propagation of seismic waves and the effects of scattering. A problem encountered with the EP records is the occurrence of fairly high amplitude

ORIGINAL PAGE IS
OF POOR QUALITY

electronic noise at times when the detonation timing circuits are activated. This noise severely reduces the signal-to-noise ratio. Cooper^{and Kovach}, 1975, used prediction filtering to remove this noise. Close scrutiny of the signals reveals, however, that the noise repeats exactly with a periodicity of 21 data points and can thus be removed by simple subtraction of the noise waveform from the signal. Power spectra of the corrected EP data were then calculated and studies of the amplitude decay with frequency and range were made. The results were then compared with the predictions made by Nakamura, et al., *in prep.* using PSE data. His results predict that a two-dimensional scattering model should fit the data to ranges of at least 4 km. As shown by Cooper^{and Kovach} (1975) and in Figure 3, this is not the case. The two-dimensional scattering model is valid to ranges of only one km, beyond which the observed signals reach their peak amplitudes well before the model would predict. The early peak is apparently due to the increase in energy from seismic waves passing below the scattering zone, thus a change in diffusivity with depth that is not modelled in the two-dimensional case.

The variation of amplitude of the signal peak with range also does not follow the two-dimensional scattering theory (see Appendix A) and is strongly frequency-dependent in such a way as to imply that cylindrical spreading cannot account for the decay of amplitudes beyond a range that decreases with frequency, i.e. high frequencies attenuate faster than expected if values of Q are used that fit the data at short range.

When this information is coupled with that obtained from earlier studies, several qualitative observations concerning propagation of seismic waves in the moon can be made.

1) The high degree of scattering in the surface layers prevents the propagation of plane waves in any direction for more than about a wave length. The coherence of seismic waves on the three components of the PSE is very low and no coherence is observed in wave trains observed at the LSPE even though the separation between the closest sensors is about 50 meters. These observations imply that no azimuthal information as to the source of an event can be obtained from observing phase variations on orthogonal instruments or across an array.

2) The depth of the scattering zone is very shallow (probably less than about 500 m and most likely less than 100 m in most regions). This is verified by the observation that the two-dimensional scattering theory does not fit the data beyond distances that decrease with frequency.

3) The lunar surface is a very poor reflector of seismic energy. The process of scattering coupled with the strong increase in seismic velocity with depth cause almost all seismic energy that originates at or reaches the lunar surface from below to be trapped. Only seismic rays that leave the surface at angles within about one degree of vertical ever reach the lunar interior because the rapid increase in velocity with depth turns the energy back towards the surface. Seismic energy reaching the surface nearly vertically from below (such as from a deep moonquake) is severely scattered and very little energy is reflected at the near-vertical angles necessary to return to the interior. Once trapped at the surface,

ORIGINAL PAGE IS
OF POOR QUALITY

the seismic energy decays because of absorption and slow leakage into the interior. This property of trapping is observed in several ways. a) In general, seismic signal amplitude decreases monotonically with distance from the source at all times; even tens of seconds after an event, the source region will contain more seismic energy than regions farther away. b) Events originating at the lunar surface have long rise times while those originating below the scattering layer have sharp rise times. Events (such as impacts) with sources at the surface insonify a region with seismic energy that slowly spreads across the surface of the moon and, as the insonified surface gets larger, leaks energy into the interior. Events originating below the scattering layer generate seismic waves that insonify a large region of the lunar surface at once and have short rise times. c) For events from a surface source high frequencies are observed to be stronger earlier in the wave train than low frequencies. This is because the higher frequencies are more easily scattered out of the surface wave guide than low frequencies, thus arriving at the receiver earlier. This effect causes a statistical dispersion effect in lunar seismograms. d) Surface reflections are not observed. If the lunar surface were a good reflector of seismic waves, then seismic phases such as PP, PPP, SS, and SSS would be observed. The only reliable candidates for these cases are at frequencies of 0.2 Hz and lower, which may be low enough such that the scattering layer is no longer important. At higher frequencies, however, the lunar surface is far too good an absorber to allow such phases to be seen.

Three factors are necessary to explain the characteristics of all lunar events: 1) high Q in the surface regions, 2) a steep velocity gradient, and 3) scattering near the lunar surface. Without any one of these three factors the lunar seismograms would look vastly different. Unfortunately a model which incorporates all three of these factors has not been found. Models that include high Q and scattering do not include the change in scattering with depth or the increase of seismic velocity with depth. Models that include the increase in velocity with depth don't include scattering. Until an adequate model can be obtained it will be hard to quantify the above conclusions or accurately predict the characteristics of lunar seismograms and anelastic properties of the near-surface layers.

Meteoroid flux. As mentioned earlier, no new information has been obtained on the meteoroid flux because of the lack of features with which to discriminate impacts from moonquakes recorded on the LSPE. The seismic method itself (Duennebier, et al., 1975) is hampered by a lack of events with which to relate amplitude of a seismic event generated by a meteoroid impact to the energy of the impacting body. The SIVB impacts, used to obtain this parameter in the paper above, may not accurately reflect the same energy partitioning as a meteoroid impact since the density of an SIVB is about 0.01 g/cm^3 compared to a meteoroid density of greater than 1 g/cm^3 . The SIVB, being less dense, will probably transfer more of its energy into seismic waves and less into crater formation than a meteoroid impact with equivalent energy. The net result would be an underestimate of the meteoroid flux. Since no means to calibrate this effect are available, the degree of underestimation, if any, is not known.

ORIGINAL PAGE IS
OF POOR QUALITY

IV. CONCLUSIONS.

The LSPE Natural Activity Study has supplied important information concerning 1) the source mechanism and importance of thermal moonquakes in the degradation of the lunar surface and 2) the processes of scattering and propagation of seismic waves in the moon. It has been shown that thermal moonquakes are most likely caused by the movement of regolith in response to diurnal stresses at the lunar surface. These events would tend to move material from hot regions to cold regions and always in a down-slope direction, thus degrading lunar slopes. The rate of slope degradation is still uncertain but it could be fast enough to account for a significant amount of the observed smoothing of the lunar surface.

Data obtained from the LSPE explosive packages and LM impact supply valuable constraints on any forthcoming models for the propagation of seismic waves in the moon. The need for theoretical work in this area must be stressed.

There are several problems still to be solved in lunar seismology, and I have two main recommendations for future seismic experiments on the moon.

1) Low frequency array. Because of size and weight restrictions, the Apollo seismometers have limited response, and seismic background noise at frequencies below about 1 Hz is certainly below the threshold of the instruments. Because of the very low noise conditions on the moon, it may be possible to detect surface waves and normal modes of oscillation from very small events. These instruments should contain 3 orthogonal components and be spread as thickly and widely as possible on the lunar surface. The information obtained from such an array would yield considerably more information about the structure of the moon.

2) High frequency array. It is well known from the Apollo seismic experiments that a large portion of the seismic energy from natural lunar events is found in the frequency band from 0.4 to 10 Hz. However, even in this frequency band, the Apollo instruments could not detect the seismic noise level for long periods of time (especially during the lunar night). Thus more sensitive instruments are also needed in this frequency band.

At frequencies above 0.5 Hz there is very little coherence between orthogonal components, except possibly for the first arrivals. In fact, little information at all is obtained by having more than one component (other than redundancy). For this reason it would probably be more advisable to use the weight and data allocations on one high gain (vertical) instrument at each station. To obtain azimuth information it would be advisable to deploy sub-arrays with separations of from 2 to 10 km between components. In this way the phase velocity across the array could be used to obtain azimuths to teleseismic events and source locations of local events could be found. Efficient triggering systems could be employed to compress the amount of data returned.

A system of several (5 or 6) arrays containing one set of orthogonal high sensitivity long period seismometers and a sub-array of three or four higher frequency vertical instruments could vastly increase our knowledge of the deeper interior because of increased accuracy of source locations and lower threshold for event detection. In addition, much more can be learned about the meteoroid flux in the smaller mass range and about thermal moonquakes.

BIBLIOGRAPHY

Please see the reference list for Appendix A.

FIGURE CAPTIONS

1. Compressed scale payout of LSPE data. LSPE data tapes were played out on a Versatec matrix plotter using the Marine Science Institute's PDP-15 computer. The numbers on the left correspond to the geophone numbers. The data was expanded from its log-compressed state and then rectified. The year, day, hour, and minute are labeled every 10 minutes. The event labeled "MQ" is a known thermal moonquake. (The year should be "73" -- not "72" as labeled).
2. Spectrograms of LSPE events. Spectrograms of three common types of events are shown. The top section of each figure shows two band-pass envelope recordings of the signal, the dark squares at a frequency of 19.2 Hz and the light squares at 9.2 Hz. The bottom section shows the power level contoured in 10 db steps plotted with frequency as the vertical axis and time as the horizontal axis. The time in days, hour, and minute is given at the bottom. These spectrums are not corrected for the frequency response of the instrument. Note that, compared to the thermal moonquakes, the LM event (center) has a sharp rise time and heavily banded spectrum.
3. Comparison of explosive package (EP) signal envelopes to theoretical scattering model. The theory, given by Nakamura et al., *in prep.*, suggests that the signal amplitude should obey the function shown by the dashed line. While the fit is excellent at ranges (given in km by the number after the EP number) of less than 1 km, the theory obviously breaks down beyond 1 km. At lower frequencies, the theory fits to greater ranges.

PRECEDING PAGE BLANK NOT FILMED

FIGURE 1



72 195 1700

72 195 1650

72 195 1600

72 195 1630

72 195 1620

72 195 1610

①

②

③

④

ORIGINAL PAGE IS
OF POOR QUALITY

ORIGINAL PAGE IS
OF POOR QUALITY

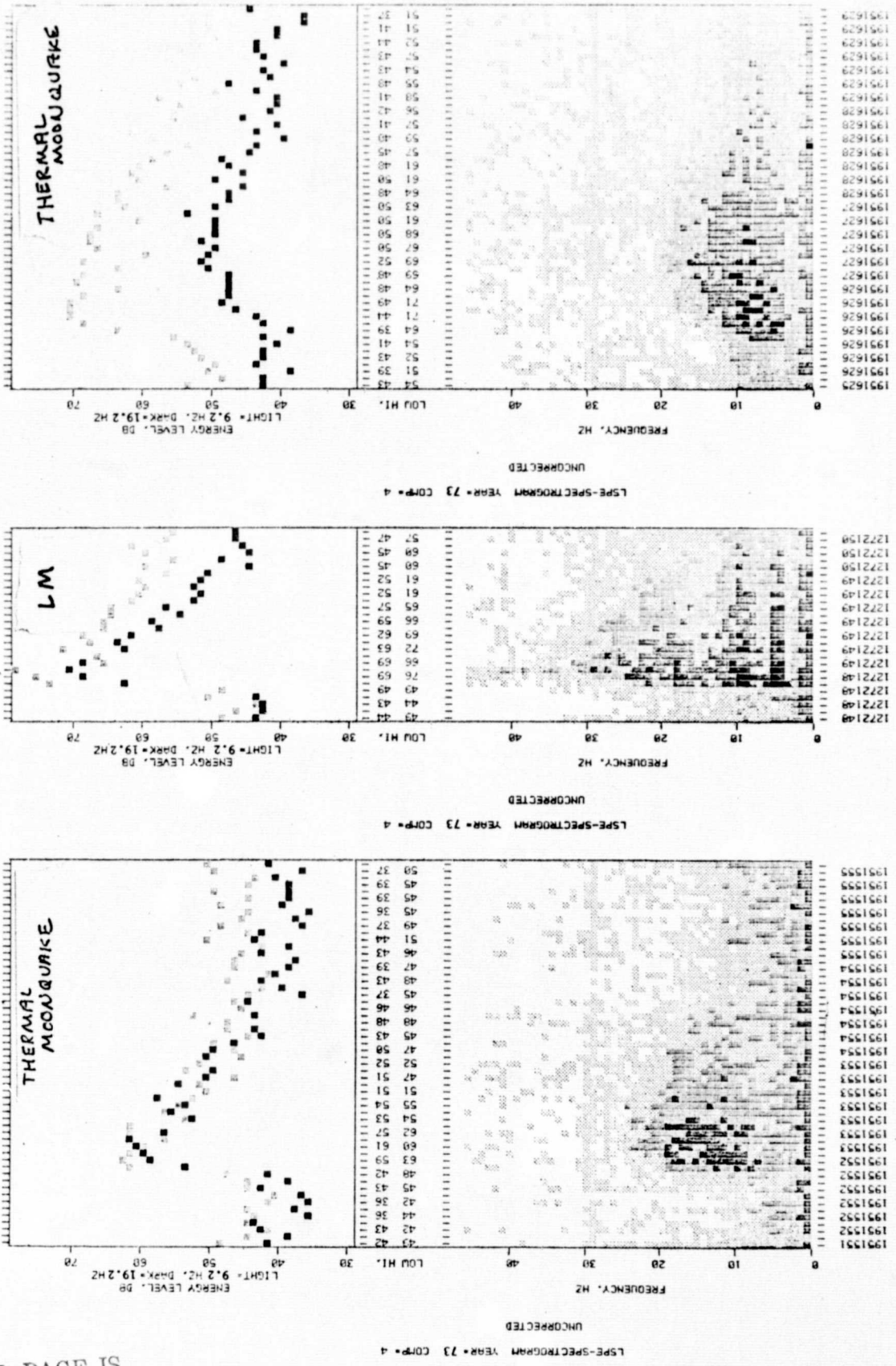


FIGURE 2

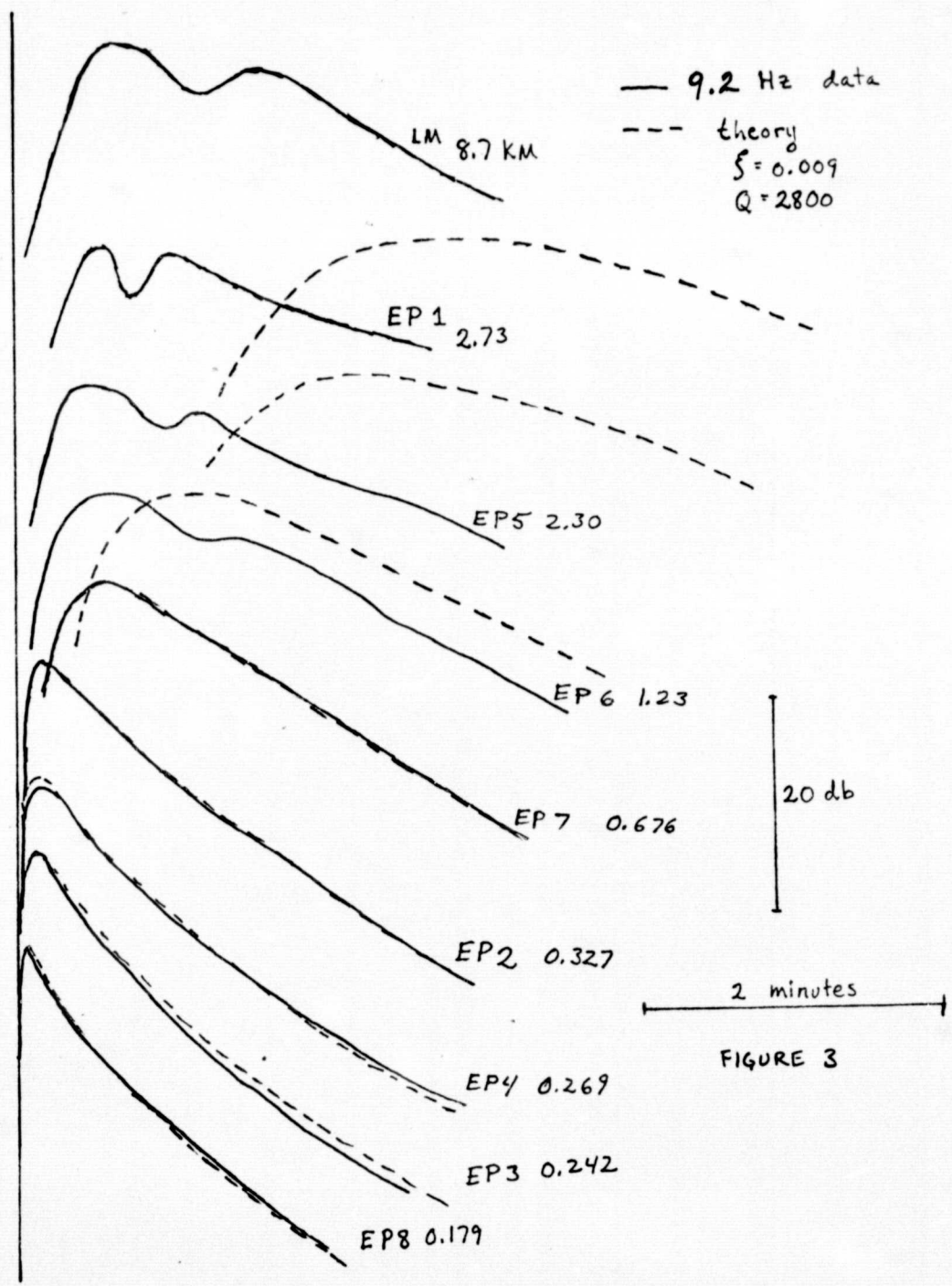


FIGURE 3

APPENDIX R

LSI (3) 329

Proc. Lunar Sci. Conf. 7th (1976), p. 000-000.
Printed in the United States of America

Thermal movement of the regolith

FREDERICK DUENNEBIER

The Marine Science Institute, University of Texas, Galveston, Texas 77550

Abstract—High frequency seismic events observed at the Apollo landing sites indicate that the regolith moves in response to diurnal variations of thermal stresses. While these events (thermal moonquakes) are observed by each of the Passive Seismic Experiment stations, the location of their sources was not possible until they were observed by the Lunar Surface Profiling Experiment (LSPE) array at the Apollo 17 site. Comparison of source locations with topographic features near the LSPE array indicate that thermal moonquakes are associated to some degree with large rocks and to a larger degree with craters. They are not associated with man-made disturbances of the lunar surface. It is suggested that thermal moonquakes are the seismic expression of a phenomenon that is actively degrading slopes on the lunar surface.

INTRODUCTION

THERMAL MOONQUAKES RECORDED by the Apollo Passive Seismic Experiments have been studied by Duennebier and Sutton (1974). Their paper provides a description of the important features of the observed seismic signals including periodicities, spectral content, and amplitudes. They suggest that thermal moonquakes are generated by fracturing of rocks along zones of weakness and by downslope movement of the regolith in response to thermal stresses as the surface temperature changes during each lunation. Because of their very small size, each thermal moonquake is recorded at only one station of the Passive Seismic Experiment, leaving no possibility of finding their source location. Estimates of the source distance were obtained by using knowledge of how the signal shape changes with distance, but this information yields only a rough guess at range, and no reliable correlations with surface topography were possible.

d/ With the deployment of the LSPE array it became feasible to locate nearby thermal moonquakes by standard triangulation techniques. The array consists of four geophones arranged in a triangle about 100 m on a side with one geophone in the middle. Although the prime experiment for the array was profiling of the surface layers using explosive sources and seismic refraction techniques (Kovach *et al.*, 1973), a four-day period of passive listening during the summer of 1973 showed that large numbers of thermal moonquakes were being detected by the array. To capitalize on this fact, a series of seven additional four-day listening periods were scheduled such that data from a complete lunation (29½ days) could be obtained over a period of about one year beginning in July, 1974. The listening periods were staggered to create a minimum of interference to other experiments at the Apollo 17 site. Two additional periods supplied redundant data for correlation of events occurring at times of maximum and minimum activity. A six-hour period bracketing a solar eclipse was also recorded. Events recorded

PRECEDING PAGE BLANK NOT FILMED

ORIGINAL PAGE IS
OF POOR QUALITY

during the extended listening periods numbered in the thousands, with 1730 events recorded that were sufficiently large to warrant detailed study. Except for 324 events generated within the lunar module (LM), and a few possible meteoroid impacts, all the events detected appear to be thermal moonquakes.

The origin of thermal moonquakes has been a mystery since they were first observed on the Apollo 11 records. At that time, the most likely source was thought to be the LM. Data obtained from the Apollo 14 mission showed that there had to be many sources for these events that tend to repeat themselves at the same time of each lunation in numbers correlating with the diurnal heating cycle (Fig. 1). Because of this repetition they could not be generated by meteoroid impacts and because there were multiple sources they could not all be generated in the LM, thus the term "thermal moonquakes" (Duennebie, 1973). Because of an apparent correlation of activity levels with the topography near each station it was suggested that thermal moonquakes could be generated by downslope movement of the regolith (Duennebie and Sutton, 1974). The calculated source

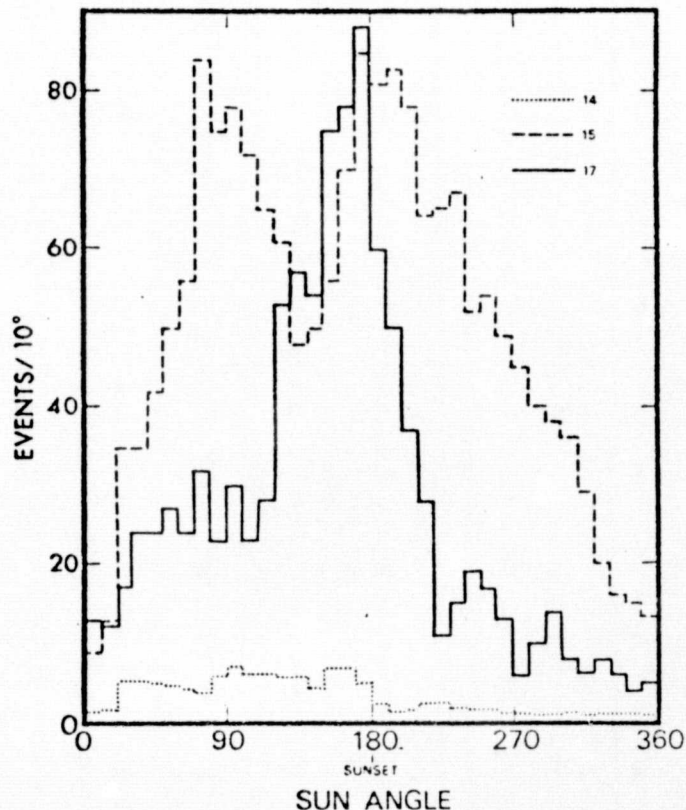


Fig. 1. Thermal moonquake activity at the Apollo 14, 15, and 17 sites. Ten degrees change in sun angle is about 20 hr. Activity increases abruptly about 2 days after sunrise and decreases rapidly after sunset at each station.

energies imply a change in slope of 1° per 100 m.y. in a 100 m deep crater (Cooper and Kovach, 1975). In this paper it will be shown that steep slopes are not a necessary condition for the occurrence of thermal moonquakes but that source locations do tend to correlate with crater locations.

ANALYSIS

Thermal moonquake data obtained from the Passive Seismic Experiment has been discussed by Duennebier and Sutton (1974). A short review of their findings is in order. Characteristics of thermal moonquake signals include emergent beginnings, durations of one to four minutes, frequencies between 2 and 10 Hz, and small amplitudes. Signals with nearly identical wave forms repeat each other regularly every lunation usually within one hour of the same sun angle. Some signals show gradual changes in wave forms from one event to the next and others show cyclic changes; some occur once per lunation, others occur many times per lunation. About 50 distinctly different types of events were observed at the Apollo 14 site (indicating 50 different sources) while more than 250 types are recognized at the Apollo 15 site. The short-period component does not operate at the Apollo 12 site, and thermal moonquakes have not been actively studied at the Apollo 16 site, although they apparently are visible on the long-period components as well as on the short-period component. Four years after the initial study was made at Apollo 14, an attempt was made to predict the occurrence of five types of events based on when they occurred four years earlier. Three of the five types were observed within an hour of when they were predicted. The wave forms had not changed noticeably in four years.

ive
Data from the Apollo 17 LSPE array has been described by Kovach *et al.* (1973) and Cooper and Kovach (1975). These data are an improvement over the PSE data in that each event is observed on four seismometers instead of one, thus making source location possible. An excellent set of calibration events at known ranges and with known energies are available. Characteristics of thermal moonquakes observed by the LSPE are nearly the same as those observed at the other seismic stations. Differences worthy of note are: (1) event activity has a single peak near sunset rather than multiple activity peaks earlier in the lunation as at the PSE stations (Fig. 1); and (2) while thermal moonquakes at the PSE stations are generally restricted to frequencies below 10 Hz, events with energy as high as 20 Hz are observed at the LSPE (Fig. 2).

SOURCE LOCATION

The standard technique for seismic source location is to use the relative times of phase arrivals at each seismometer and, knowing the velocity structure, triangulating to find the source. This method is not useful for LSPE data because there are no clear phases observed; even the largest signals have emergent beginnings. Signals observed at one end of the array show no obvious wave form correlations with signals from the same event observed 100 m away, or even 50 m

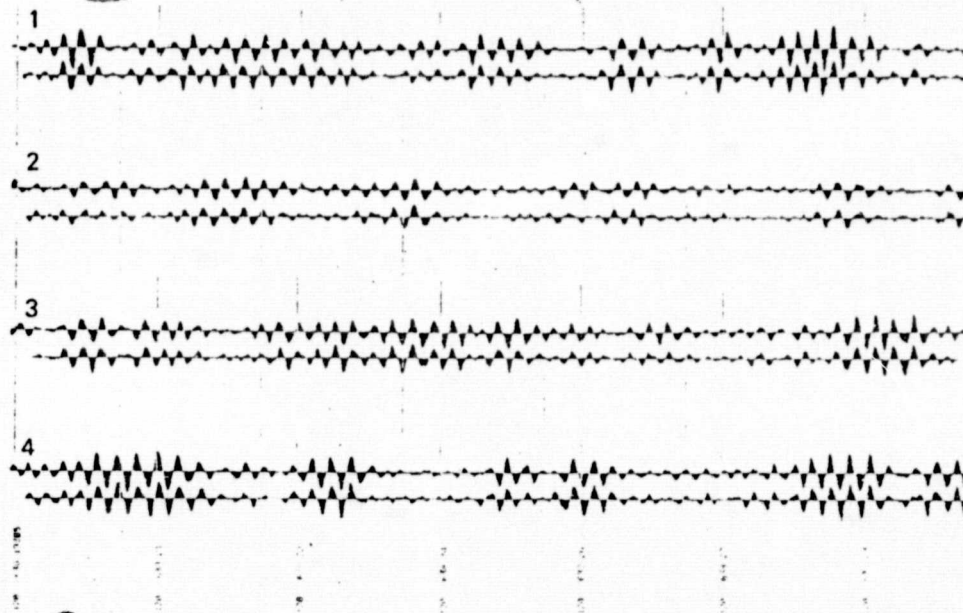
ORIGINAL PAGE IS
OF POOR QUALITY

invest A

4

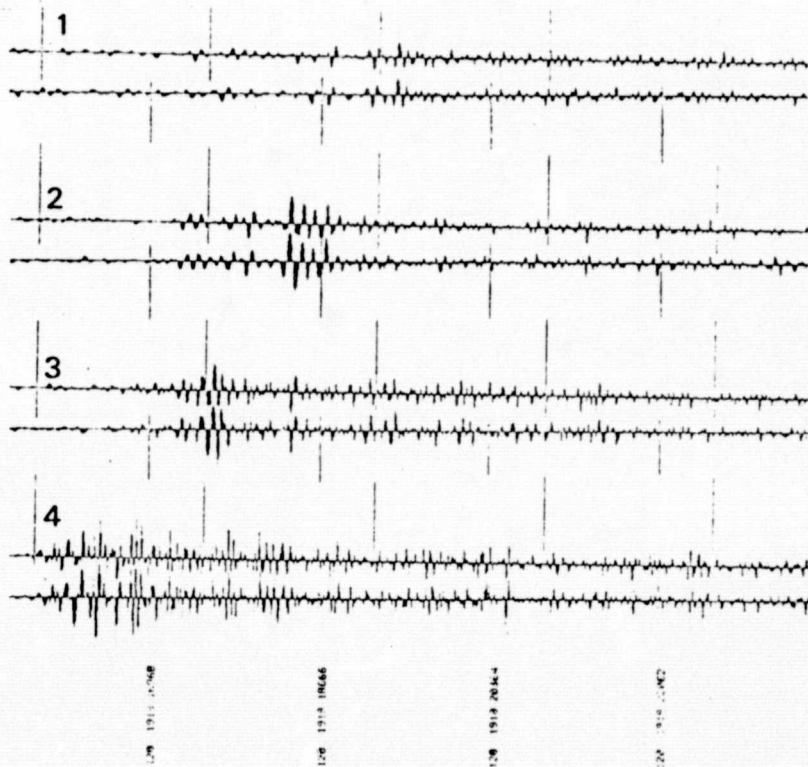
F. DUENNEBIER

A



invest B

B



away (see Fig. 2). However, the observed events can be divided into two types, those which show constant signal amplitude across the array and those that show significant variations in amplitude. Not including events from the LM, about 200 events fall into the group showing variations, or approximately 15% of the total number of natural events observed. If the rate of change in amplitude with range is known, then this amplitude variation can be used (in place of travel time differences) to locate events.

Records from the LSPE explosive packages (E.P.'s) detonated at ranges from 100 to 3000 m provide an excellent means of determining the change in amplitude with distance. Because high frequencies attenuate faster than lower frequencies, care was taken to filter the records sharply (in the frequency domain) at frequencies of interest to avoid contamination from energy at other frequencies. To eliminate the effects of changes in efficiency of the different E.P.'s in coupling into seismic waves, the source energy was not included as a known parameter and each E.P. was treated as independent of the others (Nakamura *et al.*). The equation

$$A = kr^{\alpha} \quad (1)$$

was used as a model where A is the amplitude at the peak of the signal and r is the distance from the source to the seismometer. The value of α was then found which best fit the E.P. data. The resulting analysis shows that the amplitude varies inversely as the distance at low frequencies near 5 Hz, ($\alpha = -1$), the exponent increasing to the square of the distance at 19 Hz ($\alpha = -2$). The higher frequency was chosen for determination of source locations since it is the highest frequency at which most nearby events have sufficient energy for measurement and because the large variation in amplitude with distance increases the accuracy of source locations. The calibration data are shown in Fig. 3.

The differences between measured amplitudes for natural events are used to triangulate using a least squares technique to find the source. Because of the small size of the array and a ± 1 db accuracy in amplitude determination, only those events within about 400 m of the center of the array could be located. As a check on the accuracy of the method, 30 events from the LM were also located. About 80% of these source locations fall within 50 m of the LM (located 187 m from the center of the array). Thus, while the method hardly gives pinpoint accuracy, it is capable of showing regions of high activity. Note that accuracy will decrease as distance from the array increases and that the average event size increases as

Fig. 2. (a) Thermal moonquakes recorded by the Apollo 17 LSPE. Recordings of two different but nearly identical events (177 days apart) are shown as recorded by the four LSPE geophones. Note the lack of coherence of wave forms across the array and the high coherence between the two events indicating that they both originate at the same source. Vertical lines are 1.7 sec apart. (b) Nearby thermal moonquakes. Note the change in amplitude and frequency content recorded at each geophone. These events occurred within 100 m of the array, closest to geophone 4. One event occurred 177 days before the other.

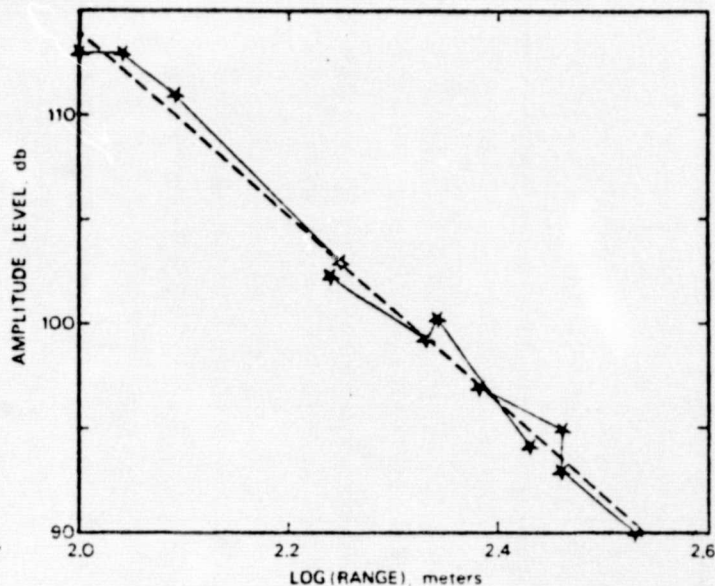


Fig. 3. Calibration data for amplitude-distance relationship. The data are from the LSPE explosive package recordings filtered at 20 Hz. The source energy for each event was not constrained to allow for changes in efficiency. The least-squares curve shown has a slope of -2 .

distance increases: i.e., small events recorded from sources inside the array may not be visible when their sources are outside the array.

The resulting source location map for natural events is shown in Fig. 4. The event activity is obviously not random but concentrated in two regions, one inside the array and one to the southwest. The concentration inside the array correlates well with the location of "geophone rock," a 2 m high rock to the southwest of the center geophone. However, none of the events from this region are large enough to be observed had they occurred more than 100 m from the array, indicating that rocks of this size are not the source of strong thermal moonquakes. Comparing Fig. 4 with Fig. 5, the concentration of activity to the southeast corresponds generally to the location of two subdued craters, approximately 100 and 200 m in diameter. The concentration to the northwest correlates with a crater (Rudolph) about 50 m in diameter. Just as important are the features that do not correlate with regions of thermal moonquake activity. The areas of astronaut activity in this region (concentrated mostly around the LM and the central ALSEP) station slightly to the north of the center geophone) show no signs of any unusual activity, thus thermal moonquakes do not appear to be generated by man-made activity. The relatively featureless region to the southwest is also an area of little activity. Note again that many events occurring near the edges and outside of the region shown in Fig. 4 were not located because the change in amplitude across the array is too small for adequate resolution. Had the array size been 1 km on a side rather than 100 m, many more events could have been located by this method.

SOURCE ENERGY

Previous estimates of thermal moonquake source energy were hampered by lack of adequate calibration events (Duennebier and Sutton, 1974). While this is still a problem, the LSPE has the advantage that the calibration events and thermal moonquakes are recorded on the same instruments, thus no conversion to actual ground motion needs to be made. The calibration events (eight small explosions and the Apollo 17 LM impact) are still not ideal in that the source mechanisms for explosion events (including the LM impact) and for moonquakes are vastly different and a much larger fraction of moonquake source energy will be converted into seismic waves than for explosives. In addition, no single model has yet been found that explains the characteristics of the signals. Cooper and Kovach (1975) noted that a two dimensional scattering equation failed to predict the observed signal rise times beyond a range of about 1 km. They also show that the value of k in Eq. (1) is frequency and range dependent (k decreasing as range and frequency increase). While Eq. (1) is an adequate model for near-range source location, a better model is needed for estimating the source energy of natural events. For this purpose the equation

$$A = kE_0^{1/2} r / \exp \left[\frac{-\pi f r}{QU} \right] \quad (2)$$

is used, where A is the amplitude at the peak of the signal envelope, E_0 is the source energy, f is the frequency, Q is the quality factor, and U is the average velocity of the signal peak. The values of A , E_0 , r , f , and U are measured and the values of k , α , and Q are found by regression. Only data within 1.5 km (6 events) are used since beyond this range the value of U becomes a complicated function of range and frequency. Within 1.5 km U averages about 25 m/sec. The regression to find the values of k , α , and Q was done for two frequency bands (4-13 Hz and 13-21 Hz) with 120 amplitude measurements used in each (5 frequencies per record, 4 records per event, and 6 events). The results yield $k = 141$ and 86 , $\alpha = -1.02$ and -0.97 , and $Q = 1400$ and 740 for the low and high frequency sets respectively.

To obtain source energy estimates for thermal moonquakes, the difference between explosives and moonquakes in the fraction of energy transformed into seismic waves must be accounted for. It is estimated that thermal moonquakes are more efficient at seismic wave generation such that the value of k should be increased by a factor of from 32 to 100 (Duennebier and Sutton, 1974). The resulting energies, based on the events located earlier and the higher frequency band parameters above, range from 10^4 to 5×10^5 ergs per event. Note, however, that only the events closest to the array are included, since it is only for these events that a reasonably reliable range is available. Events of similar amplitude at greater distances are observed implying that thermal moonquakes with source energies of 10^7 ergs and possibly larger do occur.

A reliable estimate of the total amount of energy released by thermal moonquakes within a given period does not appear possible, since most of the

ORIGINAL PAGE IS
OF POOR QUALITY

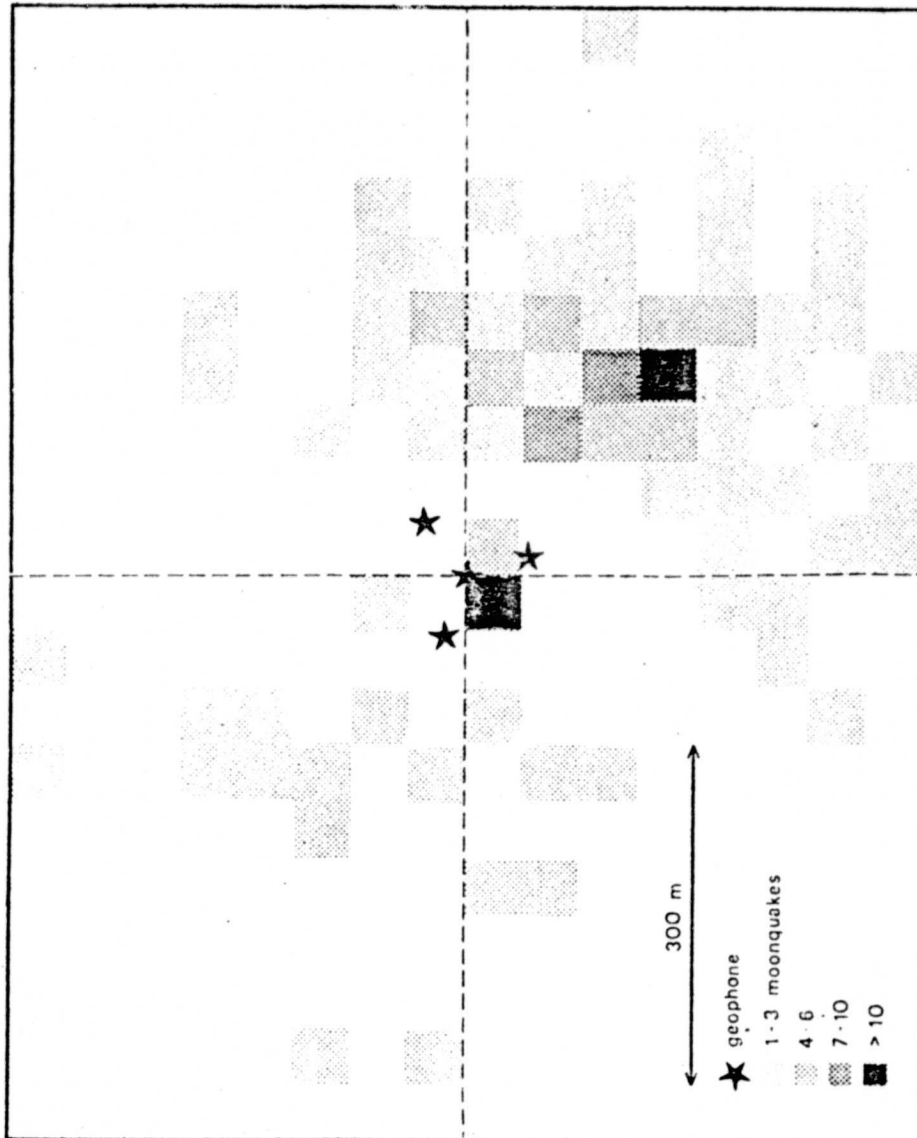


Fig. 4. Source location map. Density of thermal moonquake source locations in the region surrounding the LSPE array are shown. Comparison with Fig. 5 shows a correlation of activity with large craters and with geophone rock. This map does not include events that are known to originate at the L.M.

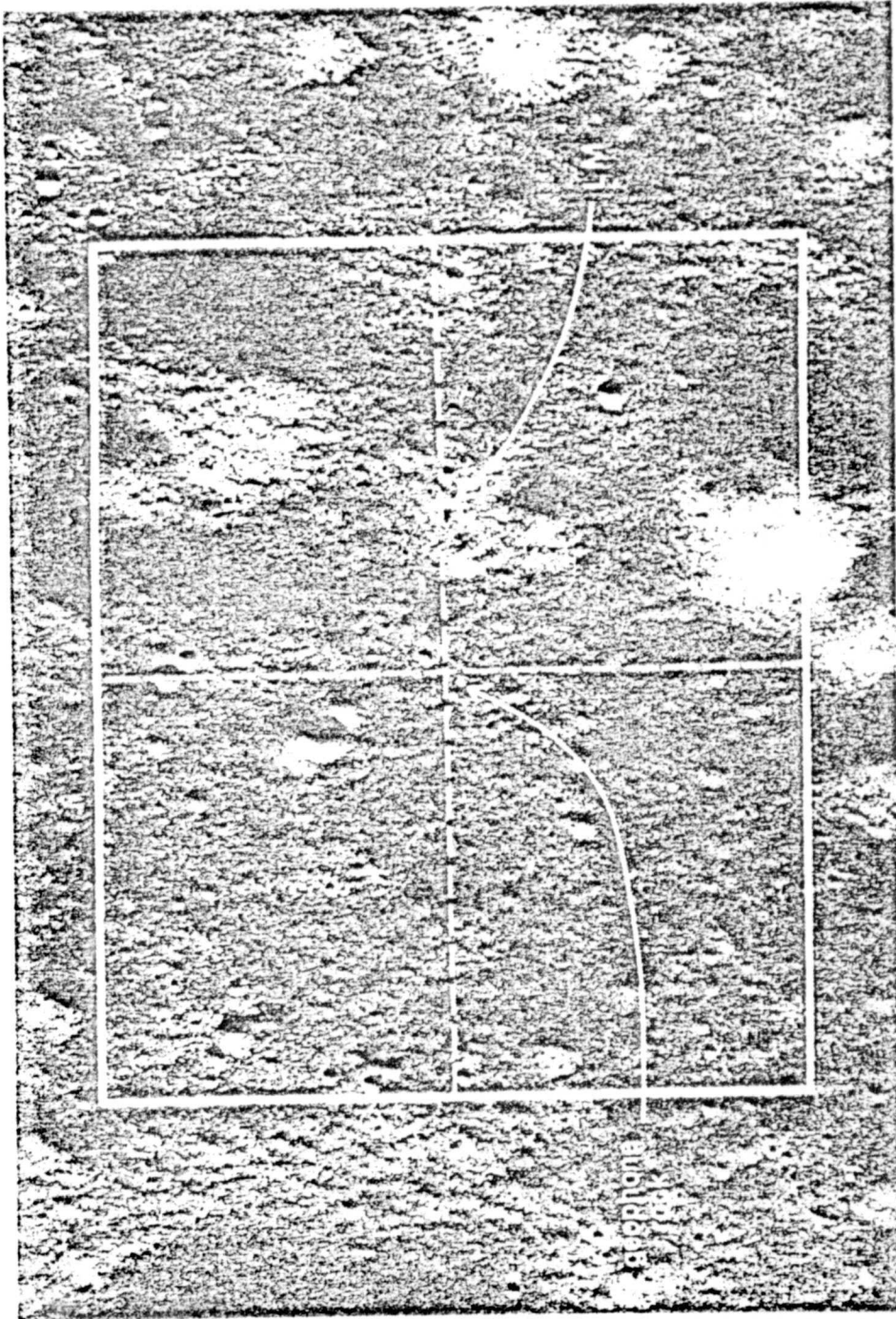


Fig. 5. Photograph of Apollo 17 landing site (AS 17-2309 Apollo 17 panoramic camera). The region within the lines is the same as that in Fig. 4. Note that the shadows of geophone rock and the LM are clearly seen.

ORIGINAL PAGE IS
OF POOR QUALITY

energy is apparently released by events smaller than the detection threshold of the system (Duennebie and Sutton, 1974; Cooper and Kovach, 1975). All plots showing cumulative number of events vs. (log) amplitude for thermal moonquakes have slopes steeper than -2 , implying that the large numbers of small events release more energy than the larger events. As a minimum estimate for total energy release from thermal moonquakes, consider the LSPE region where more than 100 events of 10^5 ergs occur each month within 1 km^2 . Extrapolating to the whole moon, 10^{15} ergs per year are released, as compared with 10^{13} ergs per year for the deep moonquakes (Lammlin *et al.*, 1974). Note that this estimate for thermal moonquakes is conservative in that the region around the LSPE site is relatively smooth whereas rougher regions are expected to have more activity.

For comparison with another energy source, a single 500 gm meteoroid impacting at 20 km/sec releases 10^{15} ergs of energy. An impact of this size or larger is expected about 2000 times per year on the lunar surface (Duennebie *et al.*, 1975). This difference in energies implies that the impacting process should be dominant in the sculpture of the lunar surface, which is no surprise since the surface is covered with craters. However, thermal movement observed as thermal moonquakes could be an important process in modification of existing features.

SOURCE MECHANISM

With the evidence presented in this paper and others (Duennebie and Sutton, 1974; Cooper and Kovach, 1975), the correlation of thermal moonquakes with the solar heating cycle appears inescapable. Additional confirmation was obtained during a solar eclipse for which the LSPE array was activated. The cooling period of the eclipse is characterized by relatively large thermal moonquakes occurring at about 25 events per hour (Fig. 6). During reheating after return of sunlight, activity is reduced to mostly small events, although activity is still much higher than normal. It is difficult to imagine a mechanism other than thermal stress variations which could cause such activity.

As mentioned earlier (and by Duennebie and Sutton, 1974), thermal cracking of rocks along zones of weakness can account for only a fraction of the events observed. While events are observed from the region around geophone rock, they have very small source energies compared to most thermal moonquakes, although this is one of the largest rocks in the vicinity. The change in wave form from one event to the next during each lunation for some types of thermal moonquakes implies movement of the source, thus excluding rocks as sources for these events (Duennebie and Sutton, 1974).

Most thermal moonquakes appear to be generated in the regolith itself. Duennebie and Sutton (1974) arrived at this conclusion and suggested that the dominant source mechanism is triggering of gravitational downslope regolith movement by thermal stress variations. While this conclusion still appears to be true, many of the source locations obtained in this paper do not correlate with steeply sloping surfaces. Apparently the thermal stresses alone supply enough

//

Thermal movement of the regolith

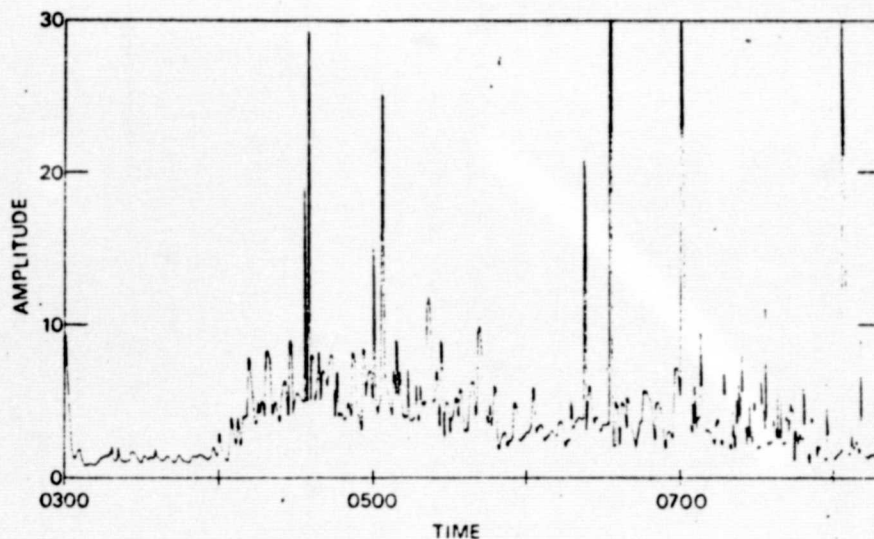


Fig. 6. Thermal activity at the LSPE during a solar eclipse. Trace amplitude is plotted as a function of time in hours.

energy for generation of thermal moonquakes and gravitational energy is not necessary.

Duenebier (1972) showed that, with reasonable assumptions as to the compressibility and strength of the lunar soil, the thermal stresses in the soil would be the same order of magnitude as the shear strengths, thus movement of the soil can be expected. Where and how these movements occur is a matter for debate. One might expect that the lunar soil would expand and contract elastically or "fracture" at close enough intervals such that no deformation or movement would be noticeable. Yet, if the new values of source energy are correct, using the method of Duenebier (1972) the source area will be about 6 cm in radius for a thermal moonquake of 10^6 ergs source energy, and maximum displacement in this region will be about 1 mm. This value is reasonable within an order of magnitude if one assumes a linear thermal expansion coefficient of 5×10^{-6} , in which case the diurnal surface temperature variation of 380°K requires 1 mm expansion for every 50 cm.

Regions in which one might expect the most movement are where there are strong contrasts in temperature or in physical properties; such contrasts exist across shadow lines caused by sloping surfaces or other surface features. The tendency will be for material on the hot (expanding) side to push into the cold (contracting) side. In craters, the shadowed portion is always downslope of the sunlight region such that material will always tend to move downslope (Fig. 7). Along sharp crater edges, material from outside the crater will tend to move into the crater. Generally, the process will be one of smoothing the lunar surface and eradicating discontinuities. Cooper and Kovach (1975) computed a slope degradation rate of 1° per 100 m.y. for a 100 m high slope using the method of Duenebier

ORIGINAL PAGE IS
OF POOR QUALITY

111 (a) 10) K

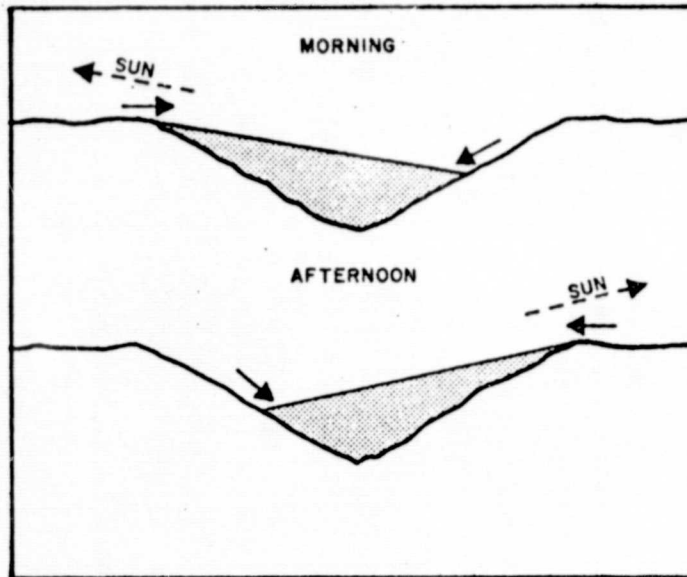


Fig. 7. Movement of the regolith caused by thermal expansion across shadows. The motion, shown by arrows, from hot regions to cold (shaded) regions, will always be downslope.

and Sutton (1974) which assumes that gravitational energy is the only energy released. Using this model, fewer events are necessary to change gentle slopes by one degree than steep ones in craters of the same radius since the change in gravitational energy per degree becomes less as the slope decreases. Intuitively, however, one expects slopes to change rapidly when they are steep and more slowly as the erosion continues. It seems likely that gravitational energy release is significant only for steeper slopes. For shallower slopes, thermal energy probably dominates. If we assume that a typical thermal moonquake displaces 200 cm^3 by 1 mm, then we can calculate how long it will take to degrade a crater by 1° by computing how much volume is moved in the transition. A 5 m deep crater with a 28 m radius will have about 130 m^3 of material transported an average of 14 m in a slope change from 10° to 9° . This process will require 9×10^7 average thermal moonquakes and will take about 9 m.y. if 1000 events occur per year. A crater with twice the radius (and four times the number of events occurring per year) will require four times as long for the same transition.

CONCLUSIONS

The movement of the regolith in response to diurnal stresses is observed as small high-frequency seismic events (thermal moonquakes). Most of this movement is apparently related to regolith materials moving downslope as sunlit areas expand and shadowed areas contract. Sources of thermal moonquakes located

near the Apollo 17 LSPE array suggest that slopes at angles much less than the angle of repose can be sources for these events indicating that thermal energy, rather than gravitational energy, is the main driving force. The effect of this activity over long periods of time will be to fill craters, lessen slopes, and generally eradicate sharp discontinuities.

The importance of this mechanism relative to others suggested for lunar erosion is hard to determine. Other mechanisms include impact cratering (Arnold, 1975), seismic shaking (Houston *et al.*, 1972; Schultz and Gault, 1975), and electrostatic transport (Gold, 1960; Criswell, 1972). Downslope movement is well documented by photographs and astronaut observations (Mattingly, 1973), and the fact that fines move downslope faster than large rocks can be seen from the observation that fillets occur preferentially on the uphill side of rocks (Muehlberger *et al.*, 1973). Crater life times are also significantly shortened on slopes (Basilevsky, 1976). In an earlier paper (Duennebieer and Sutton, 1974) it was suggested that more than enough energy was available in the thermal moonquake mechanism to account for significant erosion of lunar slopes. Although the source energy estimates have been lowered significantly since that paper, this conclusion still appears to be valid.

Acknowledgments—This research was sponsored by the National Aeronautics and Space Administration under contract NAS-9-14405. University of Texas, Marine Science Institute contribution # 090. Author is now at the Hawaii Institute of Geophysics, University of Hawaii, Honolulu, Hawaii 96822.

REFERENCES

- Arnold J. R. (1975) Monte Carlo simulation of turnover processes in the lunar regolith. *Proc. Lunar Sci. Conf. 6th*, p. 2375-2395.
- Basilevsky A. T. (1976) The rate of crater evolution (abstract). In *Lunar Science VII*, p. 33. The Lunar Science Institute, Houston.
- Cooper M. R. and Kovach R. L. (1975) Energy, frequency, and distance of moonquakes at the Apollo 17 site. *Proc. Lunar Sci. Conf. 6th*, p. 2863-2879.
- Criswell D. R. (1972) Lunar dust motion. *Proc. Lunar Sci. Conf. 3rd*, p. 2671-2680.
- Duennebieer F. K. (1972) "Moonquakes and meteoroids: Results from the Apollo passive seismic experiment's short-period data." Ph.D. Dissertation, Univ. of Hawaii.
- Duennebieer F. and Sutton G. H. (1974) Thermal moonquakes. *J. Geophys. Res.* 79, 4351-4363.
- Duennebieer F., Dorman J., Lammlein D., Latham G., and Nakamura Y. (1975) Meteoroid flux from passive seismic experiment data. *Proc. Lunar Sci. Conf. 6th*, p. 2417-2426.
- Gold T. (1960) Processes on the lunar surface. In *The Moon* (Z. Kopal, ed.), p. 433-439. Academic Press, London.
- Houston W. N., Moriwaki Y., and Chang C. S. (1972) Downslope movement of lunar soil and rock caused by meteoroid impact. *Proc. Lunar Sci. Conf. 3rd*, p. 2425-2435.
- Kovach R. L., Watkins J. S., and Talwani P. (1973) Lunar seismic profiling experiment. In *Apollo 17 Prelim. Sci. Rep.*, NASA publication SP-330, p. 10-1 to 10-12.
- Lammlein D. R., Latham G. V., Dorman J., Nakamura Y., and Ewing M. (1974) Lunar seismicity, structure, and tectonics. *Rev. Geophys. Space Phys.* 12, 1-21.
- Mattingly T. K. (1973) Impressions of the lunar highlands from the Apollo 16 command module (abstract). In *Lunar Science IV*, p. 513-514. The Lunar Science Institute, Houston.
- Muehlberger W. R., Baston M. R., Cernan E. A., Freeman V. L., Hait M. H., Holt H. E., Howard K. A., Jackson E. D., Larson K. B., Reed V. S., Rennison J. J., Schmitt H. H., Scott D. H., Sutton D. H.,

14

1/11/68

- Stuart-Alexander D., Swann G. A., Trask N. J., Ulrich G. E., Wilshire H. J., and Wolfe E. W. (1973) Preliminary geologic investigation of the Apollo 17 landing site. In *Apollo 17 Prelim. Sci. Rep.*, NASA publication SP-330, p. 6-1 to 6-91.
- Nakamura Y., Duennebier F., Latham G., Dorman J., and Lamlein D. Seismic scattering in the moon. In preparation.
- Schultz P. H. and Gault D. E. (1975) Seismically induced modification of lunar surface features. *Proc. Lunar Sci. Conf. 6th*, p. 2845-2862.

APPENDIX B

EVENT LIST

The following list identifies all events of amplitude 7 mm or greater on the compressed time scale playouts described earlier. The date, time, amplitude, selenographic longitude of the sun, and event type are given for each event. The event type identifies the event as either unknown (blank), thermal moonquake (number) or noise from the lunar lander (LM, LM1, or LM2). LM1 and LM2 events are noticeably different LM events. When a number is assigned to an event, this means that another event has been found identical to it. The number assigned is as close to the longitude of the sun at that time as possible. At periods of high activity a letter may replace the units digit. The last page of the list shows events detected during an eclipse.

YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1973	194	17	17	20	019.2	
1973	194	18	15	20	018.7	
1973	194	18	40	00	018.5	
1973	194	20	44	30	017.5	350
1973	194	21	56	00	017.4	017
1973	194	21	02	00	017.3	
1973	194	21	18	30	017.3	
1973	194	21	24	40	017.2	
1973	194	23	02	00	017.1	
1973	194	23	28	00	016.1	
1973	194	23	26	30	016.1	
1973	195	00	35	00	015.6	012
1973	195	00	46	00	015.5	
1973	195	01	12	00	015.4	
1973	195	01	19	00	015.2	
1973	195	02	33	10	015.1	
1973	195	02	33	30	014.5	
1973	195	04	39	00	014.0	
1973	195	04	55	30	013.5	
1973	195	04	55	00	013.3	012
1973	195	07	44	15	012.1	
1973	195	09	44	20	010.9	
1973	195	11	32	40	010.0	
1973	195	12	00	00	009.8	
1973	195	12	28	45	009.6	
1973	195	15	01	00	008.2	
1973	195	15	08	10	008.2	
1973	195	15	52	30	007.8	
1973	195	16	26	30	007.5	350
1973	195	16	38	50	007.4	
1973	195	16	50	40	007.3	
1973	195	17	01	40	007.2	
1973	195	17	03	40	007.2	
1973	195	18	10	00	006.6	
1973	195	18	15	30	006.6	
1973	195	18	32	00	006.4	
1973	195	18	44	30	006.3	
1973	195	19	37	30	005.9	
1973	195	19	58	00	005.7	
1973	195	20	21	30	005.5	
1973	195	21	53	20	004.7	
1973	195	22	22	10	004.5	
1973	195	22	28	00	004.5	
1973	195	22	40	10	004.3	
1973	195	23	09	10	004.1	
1973	196	00	49	00	003.3	
1973	196	05	53	00	000.7	
1973	196	06	06	40	000.6	
1973	196	06	13	00	000.5	
1973	196	06	37	00	000.3	

YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1973	349	07	03	50	000.1	
1973	LM2	07	26	50	359.9	
1973		07	27	40	359.9	
1973		07	30	50	359.8	
1973		08	01	30	359.6	
1973		08	23	10	359.4	353
1973		08	25	40	359.4	352
1973		08	27	50	359.4	
1973		08	38	20	359.3	
1973		11	13	40	358.0	
1973		11	27	00	357.9	354
1973		11	38	20	357.7	
1973		15	00	30	356.1	
1973		16	30	00	355.3	
1973		16	45	00	355.2	
1973		17	05	00	355.0	
1973		17	29	10	354.8	
1973		17	32	30	354.8	
1973		17	51	00	354.6	
1973		18	03	20	354.5	
1973		18	33	00	354.2	
1973		18	37	00	354.2	
1973		18	44	50	354.2	
1973		19	12	00	353.9	355
1973		19	19	40	353.7	
1973		19	46	00	353.6	
1973		19	52	00	353.6	
1973		19	54	00	353.5	
1973		20	06	00	353.5	
1973		20	47	10	353.1	
1973		21	00	20	353.0	
1973		21	05	00	353.0	
1973		21	25	20	352.8	
1973		21	50	40	352.6	354
1973		22	16	00	352.4	353
1973		22	19	30	352.4	
1973		22	30	00	352.3	
1973		22	38	10	352.2	347
1973		22	53	10	352.1	348
1973		23	02	00	352.0	
1973		23	06	30	351.9	
1973		23	14	00	351.9	
1973		23	21	00	351.8	
1973		23	57	10	351.5	
1973		00	06	30	351.5	
1973		00	18	00	351.3	352
1973		00	19	40	351.3	356
1973		00	28	10	351.3	
1973		00	35	10	351.2	

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1973	198	10	10	50	334.1		1974	064	05	26	03	042.4	043
1973	198	10	36	00	333.9		1974	064	06	28	03	041.9	
1973	198	10	59	10	333.7		1974	064	06	50	02	041.7	
1973	198	12	04	30	333.2		1974	064	07	29	03	041.3	
1973	198	12	06	30	333.1		1974	064	07	57	03	041.1	
1973	198	12	23	00	333.0		1974	064	07	51	03	041.1	
1973	198	12	28	00	332.9	333	1974	064	07	16	02	040.9	
1974	062	17	18	02	060.6		1974	064	08	23	05	040.9	041
1974	062	18	33	04	059.8		1974	064	08	31	03	040.8	
1974	062	19	58	04	059.5	057	1974	064	08	50	11	040.7	043
1974	062	21	35	04	058.5	054	1974	064	08	54	20	040.6	
1974	062	23	22	04	058.6		1974	064	08	58	40	040.6	
1974	062	23	03	09	057.8		1974	064	06	09	08	040.5	
1974	062	23	08	10	057.7	054	1974	064	09	19	04	040.4	
1974	063	01	20	03	056.6	042	1974	064	09	38	05	040.3	036
1974	063	01	42	10	056.4		1974	064	10	26	05	039.8	
1974	063	01	46	20	056.4		1974	064	11	33	04	039.3	042
1974	063	01	49	10	056.4		1974	064	11	49	04	039.1	
1974	063	02	27	08	056.0	057	1974	064	12	30	04	038.8	
1974	063	02	32	30	056.0	056	1974	064	12	43	04	038.7	
1974	063	02	46	04	055.9	054	1974	064	12	49	05	038.6	
1974	063	03	22	03	055.6		1974	064	14	01	05	038.0	038
1974	063	03	33	04	055.5		1974	064	14	09	04	038.0	041
1974	063	06	33	04	054.0	054	1974	064	14	38	04	037.7	
1974	063	07	01	03	053.7	042	1974	064	16	19	05	036.9	
1974	063	07	44	03	053.4	056	1974	064	16	54	06	036.6	036
1974	063	08	32	03	053.0		1974	064	17	29	04	036.3	042
1974	063	10	03	03	052.2	048	1974	064	19	14	04	035.4	
1974	063	10	21	03	052.0	050	1974	064	19	28	02	035.3	
1974	063	12	15	06	051.1	042	1974	064	20	47	09	034.6	
1974	063	12	43	02	050.8		1974	064	23	07	04	033.7	
1974	063	13	21	05	050.5	036	1974	065	00	35	03	032.6	
1974	063	14	22	04	050.0	050	1974	065	00	41	04	032.6	
1974	063	15	20	03	049.5		1974	065	00	47	04	032.6	
1974	063	16	53	03	048.7		1974	065	01	18	07	032.3	
1974	063	17	49	04	048.3	038	1974	065	02	59	07	031.5	041
1974	063	18	03	04	048.1	042	1974	065	03	04	04	030.5	
1974	063	17	56	03	048.1	048	1974	065	04	44	04	029.8	
1974	063	19	29	03	047.4	043	1974	065	06	12	09	029.8	
1974	063	20	50	04	046.7	044	1974	065	06	15	30	029.8	
1974	063	21	59	03	046.1	043	1974	065	06	35	07	029.7	
1974	063	23	22	03	045.4	042	1974	065	06	39	20	029.6	
1974	063	23	29	03	045.4		1974	065	07	07	05	029.4	
1974	064	01	24	03	044.4	048	1974	065	08	00	04	028.9	026
1974	064	01	28	03	044.4		1974	065	08	07	05	028.9	023
1974	064	02	17	03	044.0	041	1974	065	09	32	04	028.2	
1974	064	02	30	03	043.9	044	1974	065	10	08	04	027.9	
1974	064	05	18	06	042.4	042	1974	065	10	44	04	027.6	024
1974	064	05	21	50	042.4		1974	065	11	39	05	027.1	026

YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1974	065	12	27	04	026.7	
1974	065	13	26	04	026.2	
1974	065	13	58	04	025.9	
1974	065	14	04	05	025.9	
1974	065	14	43	03	025.5	
1974	065	15	01	06	025.4	
1974	065	15	05	10	025.4	
1974	065	15	21	04	025.2	
1974	065	15	48	05	024.9	
1974	065	16	08	03	024.6	
1974	065	16	33	06	024.6	
1974	065	16	50	06	024.5	
1974	065	18	58	04	023.4	
1974	065	20	22	04	022.7	
1974	065	20	39	05	022.5	
1974	065	22	47	06	021.4	
1974	065	22	26	04	021.3	
1974	224	10	54	10	246.7	217
1974	224	10	54	30	246.5	LM2
1974	224	11	46	40	246.0	
1974	224	12	51	50	245.5	
1974	224	14	06	40	244.8	
1974	224	14	49	20	244.5	
1974	224	14	54	00	244.4	
1974	224	17	36	30	243.1	
1974	224	18	59	10	242.4	
1974	224	19	39	10	242.0	
1974	224	20	00	20	241.9	
1974	224	21	56	40	240.9	
1974	224	22	13	40	240.7	
1974	224	22	33	10	240.6	
1974	225	00	25	40	239.6	
1974	225	01	01	15	239.3	
1974	225	01	04	20	239.3	
1974	225	01	13	50	239.2	
1974	225	02	12	00	238.7	
1974	225	02	20	00	238.6	
1974	225	02	21	40	238.6	
1974	225	05	10	15	237.2	
1974	225	06	20	00	236.6	
1974	225	07	04	00	236.2	
1974	225	07	48	00	235.8	
1974	225	09	04	30	235.2	
1974	225	11	03	30	234.2	
1974	225	11	28	40	234.0	
1974	225	12	11	00	233.6	
1974	225	12	47	00	233.3	
1974	225	13	11	30	233.1	
1974	225	13	41	00	232.9	
1974	225	13	59	20	232.7	
1974	225	14	46	30	232.3	
1974	225	14	58	20	232.3	
1974	225	14	01	00	232.2	
1974	225	14	05	00	232.2	
1974	225	16	18	10	231.5	
1974	225	17	23	10	231.0	
1974	225	19	25	00	229.9	
1974	225	19	31	00	229.9	
1974	225	20	21	10	229.4	
1974	225	20	49	10	229.2	
1974	225	22	26	00	228.4	
1974	225	22	26	00	228.4	
1974	225	23	25	00	227.9	
1974	226	00	18	00	227.4	
1974	226	00	18	40	227.4	
1974	226	02	03	00	226.5	
1974	226	04	27	00	225.3	
1974	226	06	06	00	224.5	
1974	226	07	25	50	223.8	
1974	226	08	34	00	223.2	
1974	226	09	37	00	223.2	
1974	226	09	50	00	223.0	
1974	226	12	12	30	222.6	
1974	226	12	59	30	221.4	
1974	226	13	40	30	221.0	
1974	226	13	50	30	220.6	
1974	226	13	50	30	220.5	
1974	226	21	22	00	217.2	
1974	226	23	17	00	216.7	
1974	226	23	41	30	215.7	
1974	227	00	30	40	215.5	
1974	227	00	42	00	215.1	
1974	227	01	19	00	215.0	
1974	227	01	28	10	214.7	
1974	227	04	39	00	214.6	
1974	227	04	43	40	213.0	
1974	227	06	36	00	212.9	
1974	227	07	15	00	211.6	
1974	227	07	57	20	211.3	
1974	227	09	25	30	210.5	
1974	227	10	02	00	209.7	
1974	227	10	59	30	209.7	
1974	227	17	04	30	206.6	
1974	227	17	28	20	206.4	
1974	227	17	35	30	206.4	
1974	227	17	16	00	204.5	
1974	227	21	41	30	204.5	

YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1974	228	00	35	30	202.8	LM1
1974	228	01	06	30	202.5	
1974	228	01	46	50	202.2	221
1974	228	08	14	10	198.9	195
1974	228	11	18	20	197.3	223
1974	228	11	53	40	197.0	LM2
1974	249	14	23	00	299.4	
1974	249	14	37	10	299.3	
1974	249	14	46	30	299.2	
1974	249	14	54	40	299.1	
1974	249	15	12	40	299.0	
1974	249	15	22	00	298.9	
1974	249	15	23	00	298.9	
1974	249	15	28	40	298.7	
1974	249	15	51	10	298.6	
1974	249	15	59	00	298.6	
1974	249	16	03	20	298.5	299
1974	249	16	13	00	298.5	297
1974	249	16	44	50	298.2	
1974	249	17	14	20	298.0	
1974	249	17	44	30	297.7	
1974	249	17	48	20	297.7	
1974	249	17	55	50	297.6	
1974	249	17	57	00	297.6	
1974	249	18	07	00	297.5	298
1974	249	18	50	50	297.1	293
1974	249	19	04	10	297.0	
1974	249	19	35	50	296.9	
1974	249	19	38	30	296.8	
1974	249	20	00	20	296.6	296
1974	249	20	06	40	296.5	295
1974	249	20	49	50	296.1	
1974	249	21	10	30	296.0	
1974	249	21	49	20	295.6	
1974	249	21	56	00	295.6	
1974	249	22	20	10	295.4	
1974	249	22	56	00	295.1	
1974	249	23	01	00	295.0	
1974	250	00	39	40	294.2	
1974	250	01	55	30	294.1	
1974	250	01	18	00	293.7	
1974	250	01	35	40	293.9	
1974	250	01	56	00	293.5	
1974	250	02	06	20	293.5	292
1974	250	02	43	50	293.1	LM1
1974	250	02	50	20	293.1	
1974	250	02	52	00	293.1	
1974	250	02	57	00	293.0	
1974	250	03	11	50	292.9	
1974	250	03	20	00	292.8	
1974	292.3					
1974	292.5					
1974	292.1					
1974	292.0					
1974	291.5					
1974	291.4					
1974	291.3					
1974	291.3					
1974	291.2					
1974	291.2					
1974	291.1					
1974	290.7					
1974	290.6					
1974	290.5					
1974	290.4					
1974	290.3					
1974	289.9					
1974	289.8					
1974	289.8					
1974	289.7					
1974	289.6					
1974	289.4					
1974	289.4					
1974	289.3					
1974	289.2					
1974	288.7					
1974	288.5					
1974	288.3					
1974	288.0					
1974	287.9					
1974	287.8					
1974	287.7					
1974	287.5					
1974	287.1					
1974	286.9					
1974	286.9					
1974	286.6					
1974	286.6					
1974	286.4					
1974	286.4					
1974	286.1					
1974	285.9					
1974	285.9					
1974	285.8					
1974	285.8					
1974	285.6					

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1974	250	18	03	00	285.4	283	1974	251	09	22	00	277.6	LM1
1974	250	18	10	30	285.3		1974	251	09	30	00	277.5	
1974	250	18	29	40	285.1	LM1	1974	251	10	14	20	277.9	
1974	250	18	34	00	285.1		1974	251	10	45	50	276.9	
1974	250	18	59	10	284.9		1974	251	10	52	20	276.8	
1974	250	19	34	40	284.6		1974	251	11	16	30	276.6	
1974	250	19	56	00	284.4		1974	251	11	49	00	276.3	
1974	250	20	01	00	284.4		1974	251	12	10	10	276.1	
1974	250	20	02	10	284.3	LM1	1974	251	12	30	30	276.0	
1974	250	20	12	00	284.3		1974	251	12	59	20	275.7	275
1974	250	21	09	10	283.8		1974	251	13	08	50	275.6	LM1
1974	250	21	30	10	283.6	LM1	1974	251	13	20	30	275.6	
1974	250	21	57	50	283.4		1974	251	13	28	40	275.5	
1974	250	22	15	00	283.2		1974	251	15	08	40	274.6	
1974	250	22	28	50	283.1		1974	251	15	14	40	274.6	LM2
1974	250	22	45	40	283.0		1974	251	15	26	00	274.5	
1974	250	22	52	40	282.9	LM1	1974	251	15	29	40	274.5	LM1
1974	250	22	59	40	282.8		1974	251	15	33	20	274.4	
1974	250	23	38	10	282.8		1974	251	15	37	10	274.4	283
1974	251	00	57	40	282.1	LM1	1974	251	16	08	50	274.1	
1974	251	00	57	20	281.8		1974	251	16	35	00	273.9	
1974	251	01	17	10	281.7		1974	251	17	41	40	273.3	LM1
1974	251	01	50	40	281.4		1974	251	18	17	20	273.0	
1974	251	02	19	50	281.2		1974	251	19	38	00	272.3	
1974	251	02	37	00	281.0	LM1	1974	251	19	39	50	272.3	
1974	251	02	51	50	280.9		1974	251	19	56	00	272.2	LM1
1974	251	02	56	10	280.8		1974	251	20	23	10	272.0	LM2
1974	251	03	00	00	280.8		1974	251	22	08	00	271.1	LM1
1974	251	03	10	30	280.7	283	1974	251	22	24	00	270.9	
1974	251	03	59	30	280.3		1974	251	22	57	10	270.7	
1974	251	04	09	00	280.2		1974	251	23	32	50	270.4	
1974	251	04	13	00	280.2		1974	252	00	55	30	269.7	
1974	251	04	27	20	280.1		1974	252	01	03	00	269.6	LM1
1974	251	04	53	00	279.8		1974	252	01	28	40	269.4	275
1974	251	05	16	20	279.7		1974	252	03	35	40	268.3	
1974	251	05	25	20	279.6		1974	252	03	43	20	268.2	LM1
1974	251	05	47	00	279.4	LM1	1974	252	03	50	00	268.2	
1974	251	05	49	30	279.4		1974	252	04	06	50	268.0	
1974	251	06	10	30	279.2		1974	252	04	20	00	267.9	
1974	251	06	12	00	279.2		1974	252	04	28	20	267.9	
1974	251	06	19	30	279.1		1974	252	05	09	20	267.5	
1974	251	06	28	00	279.0		1974	252	06	34	50	267.3	
1974	251	07	06	40	278.7		1974	252	06	13	00	267.0	
1974	251	07	25	30	278.6	283	1974	252	06	25	20	266.9	
1974	251	07	35	30	278.5	LM1	1974	252	06	33	00	266.8	LM1
1974	251	07	50	00	278.3		1974	252	07	27	20	266.3	
1974	251	07	51	20	278.3		1974	252	07	52	00	266.1	
1974	251	08	13	20	278.2		1974	252	07	53	55	266.1	
1974	251	08	14	10	278.1		1974	252	07	56	55	266.1	
1974	251	09	18	40	277.6	LM1	1974	252	09	21	40	265.4	LM1

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1974	252	09	30	10	265.3		1974	295	20	40	30	095.1	
1974	252	10	35	00	264.7		1974	295	22	15	10	094.2	
1974	252	11	59	30	264.5		1974	295	22	17	15	094.2	
1974	252	11	01	55	264.5	223	1974	296	00	47	35	094.0	
1974	252	12	23	40	263.8		1974	296	01	01	00	093.4	
1974	252	12	27	15	263.8	LM1	1974	296	01	10	45	092.8	LM2
1974	252	14	16	40	262.9		1974	296	03	32	40	092.4	LM2
1974	252	14	22	25	262.8		1974	296	03	46	30	091.6	
1974	252	14	34	20	262.7		1974	296	05	43	30	091.4	
1974	252	15	40	30	262.2		1974	296	06	46	40	090.5	
1974	252	15	42	55	262.1	LM1	1974	296	06	46	50	090.2	
1974	252	17	53	00	261.0	275	1974	296	07	42	20	089.9	089
1974	252	18	55	30	260.5	LM1	1974	296	07	51	00	089.5	
1974	252	23	29	00	258.2		1974	296	08	28	00	089.4	
1974	253	00	01	30	257.9		1974	296	08	46	20	089.1	
1974	253	01	52	50	257.0	LM1	1974	296	13	25	05	088.9	
1974	253	01	52	50	257.0		1974	296	13	31	20	086.6	LM2
1974	253	02	29	00	256.7		1974	296	14	48	20	086.5	
1974	253	03	24	00	256.2		1974	296	14	55	00	085.8	
1974	253	05	33	10	255.1	223	1974	296	16	02	20	085.2	
1974	253	05	43	20	255.0	LM2	1974	296	16	02	30	085.2	
1974	253	05	47	30	255.0	LM1	1974	296	16	31	50	085.0	083
1974	253	07	33	50	254.1		1974	296	17	38	30	084.9	084
1974	253	08	46	00	253.5	275	1974	296	17	10	20	084.6	083
1974	253	09	43	00	253.0	LM1	1974	296	17	15	10	084.6	
1974	253	10	05	00	252.8		1974	296	17	46	50	084.3	
1974	253	11	11	30	252.2		1974	296	18	00	20	084.2	083
1974	253	13	43	30	250.9	LM1	1974	296	18	46	00	083.8	
1974	253	14	23	30	250.6	232	1974	296	18	55	40	083.8	
1974	253	14	52	40	250.4	LM2	1974	296	19	36	20	083.4	082
1974	253	16	02	40	249.8		1974	296	19	45	50	083.3	
1974	253	18	14	40	248.6	LM1	1974	296	21	13	40	082.6	LM2
1974	253	18	25	30	248.6		1974	296	21	17	30	082.6	
1974	253	19	21	00	248.1		1974	296	21	48	00	082.3	082
1974	253	19	56	55	248.1		1974	296	21	53	30	082.3	LM2
1974	253	21	25	00	246.8		1974	296	23	03	20	081.7	
1974	253	22	25	20	246.5		1974	296	23	19	50	081.5	
1974	253	23	06	00	246.2		1974	297	03	01	50	079.7	
1974	253	23	28	50	246.0	LM2	1974	297	03	05	40	079.6	
1974	295	14	55	50	098.0	LM1	1974	297	04	06	00	079.1	
1974	295	15	01	30	097.9	LM2	1974	297	04	32	00	078.9	
1974	295	16	54	40	097.0		1974	297	05	19	00	078.5	
1974	295	17	37	10	096.6		1974	297	06	07	40	078.1	
1974	295	17	40	35	096.6		1974	297	06	07	20	077.6	
1974	295	18	12	00	096.3	LM2	1974	297	07	22	50	077.4	
1974	295	18	35	30	096.1		1974	297	08	50	20	076.7	LM2
1974	295	18	44	20	096.0	096	1974	297	09	22	00	076.4	057
1974	295	18	54	20	095.9		1974	297	09	55	20	076.2	068
1974	295	19	57	00	095.4	084	1974	297	10	03	10	076.1	

YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1974	297	11	41	40	075.3	LM2
1974	297	12	06	40	075.0	070
1974	297	13	56	30	074.6	057
1974	297	17	08	20	072.5	057
1974	297	17	18	00	072.4	
1974	297	17	53	50	072.1	
1974	297	19	12	30	071.4	LM2
1974	297	19	16	30	071.4	
1974	297	19	20	10	071.4	
1974	297	20	41	50	070.7	
1974	297	21	02	20	070.5	
1974	297	21	26	30	070.3	057
1974	297	21	29	05	070.3	
1974	297	21	34	00	070.2	LM2
1974	297	21	49	50	070.1	070
1974	297	22	14	00	069.9	
1974	297	22	19	00	069.9	
1974	297	22	20	00	069.9	
1974	298	00	24	00	068.8	062
1974	298	01	03	30	068.5	
1974	298	01	31	10	068.2	057
1974	298	02	00	00	068.0	068
1974	298	02	15	40	067.9	
1974	298	03	03	50	067.5	
1974	298	04	16	40	066.8	
1974	298	05	44	10	066.1	
1974	298	06	07	10	065.9	062
1974	298	06	11	30	065.9	057
1974	298	07	41	00	065.1	
1974	298	08	23	20	064.8	
1974	298	08	46	00	064.6	
1974	298	09	04	00	064.4	
1974	298	10	08	30	063.9	054
1974	298	11	07	00	063.4	
1974	298	11	34	05	063.1	057
1974	298	11	45	00	063.1	LM1
1974	298	11	58	00	062.9	062
1974	298	12	28	20	062.7	
1974	298	12	32	30	062.7	
1974	298	15	07	10	061.3	054
1974	298	16	02	50	060.9	
1974	298	16	42	10	060.5	
1974	305	17	27	20	335.1	
1974	305	17	39	30	335.0	
1974	305	17	50	00	334.9	
1974	305	17	57	20	334.9	
1974	305	18	07	00	334.7	
1974	305	18	11	55	334.7	
1974	305	18	38	30	334.5	330
1974	305	18	38	30	334.5	
1974	305	19	06	40	334.2	
1974	305	19	28	20	334.1	
1974	305	19	58	00	333.8	338
1974	305	20	08	20	333.5	33C
1974	305	20	39	50	333.5	33D
1974	305	20	01	20	333.3	
1974	305	21	16	10	333.2	
1974	305	21	44	50	332.9	
1974	305	22	01	10	332.8	
1974	305	22	16	20	332.7	
1974	305	22	42	00	332.5	
1974	305	23	05	00	332.3	
1974	305	23	15	20	332.2	33E
1974	306	00	27	00	332.1	
1974	306	01	26	40	331.6	
1974	306	01	37	20	331.0	
1974	306	01	55	10	330.8	
1974	306	01	59	10	330.8	
1974	306	02	01	05	330.8	
1974	306	02	28	40	330.5	
1974	306	02	33	10	330.5	
1974	306	02	45	40	330.4	329
1974	306	03	48	40	330.4	
1974	306	03	20	40	330.1	
1974	306	03	47	20	329.9	
1974	306	04	37	10	329.5	326
1974	306	05	05	50	329.2	
1974	306	05	11	30	328.9	
1974	306	05	43	30	328.5	
1974	306	06	29	20	328.4	
1974	306	06	45	20	328.4	
1974	306	06	49	30	328.3	328
1974	306	07	25	00	328.0	
1974	306	08	13	00	327.6	
1974	306	08	35	00	327.6	
1974	306	08	38	45	327.4	327
1974	306	08	42	00	327.4	
1974	306	09	52	40	327.3	
1974	306	09	23	20	327.0	
1974	306	09	28	30	327.0	
1974	306	10	04	10	326.7	LM2
1974	306	10	05	50	326.7	
1974	306	10	13	20	326.6	
1974	306	10	24	00	326.5	326
1974	306	10	40	30	326.4	
1974	306	10	45	00	326.4	
1974	306	10	53	10	326.3	
1974	306	11	04	00	326.2	
1974	306	11	07	30	326.2	

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	HR	MNI	AMP	LONG	TYPE	HR	MNI	AMP	LONG
1974	306	11	16	30	326.1		11	16	30	326.1		11	16	30	326.1
1974	306	11	22	50	326.0		11	22	50	326.0		11	22	50	326.0
1974	306	11	29	30	326.0		11	29	30	326.0		11	29	30	326.0
1974	306	11	38	45	325.9	324	11	38	45	325.9		11	38	45	325.9
1974	306	12	05	00	325.7	310	12	05	00	325.7		12	05	00	325.7
1974	306	12	06	00	325.7		12	06	00	325.7		12	06	00	325.7
1974	306	12	50	30	325.3		12	50	30	325.3		12	50	30	325.3
1974	306	12	52	40	325.3		12	52	40	325.3		12	52	40	325.3
1974	306	13	01	00	325.2	320	13	01	00	325.2		13	01	00	325.2
1974	306	13	02	00	325.0		13	02	00	325.0		13	02	00	325.0
1974	306	13	22	00	325.0	325	13	22	00	325.0		13	22	00	325.0
1974	306	13	32	15	324.9	322	13	32	15	324.9		13	32	15	324.9
1974	306	13	57	00	324.7		13	57	00	324.7		13	57	00	324.7
1974	306	14	06	20	324.7		14	06	20	324.7		14	06	20	324.7
1974	306	14	25	30	324.5	321	14	25	30	324.5		14	25	30	324.5
1974	306	14	26	30	324.5		14	26	30	324.5		14	26	30	324.5
1974	306	14	49	10	324.3		14	49	10	324.3		14	49	10	324.3
1974	306	15	09	40	324.1		15	09	40	324.1		15	09	40	324.1
1974	306	15	28	30	324.0		15	28	30	324.0		15	28	30	324.0
1974	306	15	30	50	324.0		15	30	50	324.0		15	30	50	324.0
1974	306	15	32	35	323.9	323	15	32	35	323.9		15	32	35	323.9
1974	306	15	34	10	323.9		15	34	10	323.9		15	34	10	323.9
1974	306	15	54	20	323.7		15	54	20	323.7		15	54	20	323.7
1974	306	16	05	00	323.7		16	05	00	323.7		16	05	00	323.7
1974	306	16	10	25	323.6		16	10	25	323.6		16	10	25	323.6
1974	306	16	12	00	323.6		16	12	00	323.6		16	12	00	323.6
1974	306	16	14	40	323.6		16	14	40	323.6		16	14	40	323.6
1974	306	16	23	30	323.5		16	23	30	323.5		16	23	30	323.5
1974	306	16	27	15	323.5		16	27	15	323.5		16	27	15	323.5
1974	306	16	45	50	323.3		16	45	50	323.3		16	45	50	323.3
1974	306	16	50	15	323.3		16	50	15	323.3		16	50	15	323.3
1974	306	17	10	10	323.1		17	10	10	323.1		17	10	10	323.1
1974	306	17	30	00	322.9		17	30	00	322.9		17	30	00	322.9
1974	306	17	56	00	322.9	326	17	56	00	322.9		17	56	00	322.9
1974	306	18	16	30	322.6	32A	18	16	30	322.6		18	16	30	322.6
1974	306	19	00	40	322.2		19	00	40	322.2		19	00	40	322.2
1974	306	19	38	10	321.9		19	38	10	321.9		19	38	10	321.9
1974	306	19	41	45	321.8	32E	19	41	45	321.8		19	41	45	321.8
1974	306	19	45	30	321.8		19	45	30	321.8		19	45	30	321.8
1974	306	19	50	55	321.7	32C	19	50	55	321.7		19	50	55	321.7
1974	306	20	03	05	321.6		20	03	05	321.6		20	03	05	321.6
1974	306	20	04	30	321.6		20	04	30	321.6		20	04	30	321.6
1974	306	20	28	10	321.4	32B	20	28	10	321.4		20	28	10	321.4
1974	306	20	40	30	321.3		20	40	30	321.3		20	40	30	321.3
1974	306	20	42	40	321.3		20	42	40	321.3		20	42	40	321.3
1974	306	20	49	40	321.1		20	49	40	321.1		20	49	40	321.1
1974	306	21	06	30	321.1	32G	21	06	30	321.1		21	06	30	321.1
1974	306	21	32	30	320.9		21	32	30	320.9		21	32	30	320.9
1974	306	21	43	55	320.8		21	43	55	320.8		21	43	55	320.8

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1974	307	10	30	00	314.3	313	1974	307	20	59	40	309.0	303
1974	307	10	48	10	314.2		1974	307	21	02	00	309.0	30A
1974	307	11	04	00	314.0		1974	307	21	21	35	308.8	LM2
1974	307	11	10	00	314.0	311	1974	307	21	22	00	308.8	
1974	307	11	20	30	313.9	31A	1974	307	21	31	00	308.4	309
1974	307	11	28	00	313.8		1974	307	22	09	50	308.4	308
1974	307	12	04	00	313.5		1974	307	22	15	40	308.3	
1974	307	12	18	00	313.4	LM2	1974	307	22	23	40	308.3	
1974	307	12	23	00	313.4		1974	307	22	25	40	308.2	
1974	307	12	44	05	313.2	LM2	1974	307	22	31	10	308.1	307
1974	307	12	51	30	313.1		1974	307	23	43	10	307.8	
1974	307	13	26	00	312.8	31B	1974	307	23	25	40	307.7	
1974	307	13	33	00	312.8	303	1974	307	23	38	10	307.6	
1974	307	13	41	00	312.7		1974	307	23	43	10	307.5	30A
1974	307	13	50	50	312.6	31C	1974	307	23	57	20	307.5	
1974	307	14	00	00	312.6		1974	308	00	02	50	307.4	
1974	307	14	34	20	312.3		1974	308	00	16	15	307.4	
1974	307	14	40	10	312.2		1974	308	00	32	30	307.3	306
1974	307	14	41	50	312.2		1974	308	01	53	10	307.1	
1974	307	14	43	45	312.2		1974	308	01	05	20	307.0	305
1974	307	14	53	45	312.1		1974	308	01	15	25	306.9	LM2
1974	307	15	03	20	312.0		1974	308	01	43	40	306.6	LMI
1974	307	15	15	30	311.9	31D	1974	308	02	44	30	306.6	
1974	307	15	25	40	311.8		1974	308	02	22	30	306.3	LM2
1974	307	15	27	20	311.8		1974	308	02	25	30	306.2	
1974	307	15	35	00	311.7		1974	308	02	33	00	306.1	
1974	307	15	47	00	311.7		1974	308	02	40	00	306.1	
1974	307	15	57	50	311.6		1974	308	02	42	15	306.1	
1974	307	16	08	40	311.5	31E	1974	308	02	45	00	306.1	
1974	307	16	11	10	311.5		1974	308	03	07	35	305.9	30A
1974	307	16	22	15	311.4		1974	308	03	20	10	305.8	300
1974	307	16	26	15	311.3		1974	308	03	23	30	305.8	302
1974	307	16	50	00	311.1		1974	308	03	42	00	305.6	
1974	307	17	00	50	311.0		1974	308	03	54	20	305.5	
1974	307	17	07	50	311.0		1974	308	04	57	30	305.4	
1974	307	17	10	50	311.0	LM2	1974	308	04	14	40	305.4	304
1974	307	17	35	00	310.7		1974	308	04	16	00	305.3	
1974	307	17	53	30	310.6		1974	308	04	46	00	305.1	LM2
1974	307	18	20	50	310.4		1974	308	04	49	20	305.1	LMI
1974	307	18	32	00	310.3	31G	1974	308	04	50	50	305.0	
1974	307	18	44	30	310.2	31F	1974	308	04	55	35	305.0	301
1974	307	18	49	00	310.1	31H	1974	308	05	58	00	304.9	
1974	307	18	56	10	310.1		1974	308	05	04	00	304.7	
1974	307	19	01	30	310.0		1974	308	05	27	10	304.7	
1974	307	19	26	00	309.8		1974	308	05	37	10	304.6	30B
1974	307	19	49	50	309.6	LM1	1974	308	05	43	00	304.5	
1974	307	20	17	30	309.4		1974	308	06	50	05	304.1	
1974	307	20	41	20	309.2		1974	308	06	40	10	304.0	30A

* * *	YEAR	DAY	HR	MNI	* * *	TYPE	* * *	LONG	* * *	AMP	* * *	LONG	* * *	TYPE	* * *
* * *	1974	346	17	50	* * *		* * *	196.1	* * *	40	* * *	179.0	* * *		* * *
* * *	1974	346	18	56	* * *		* * *	196.0	* * *	50	* * *	178.7	* * *		* * *
* * *	1974	346	21	45	* * *	195	* * *	195.6	* * *	25	* * *	177.7	* * *	LM2	* * *
* * *	1974	346	22	49	* * *	LM2	* * *	194.0	* * *	00	* * *	175.4	* * *		* * *
* * *	1974	346	01	39	* * *		* * *	193.6	* * *	20	* * *	175.2	* * *		* * *
* * *	1974	347	03	28	* * *	LM1	* * *	192.2	* * *	00	* * *	175.1	* * *		* * *
* * *	1974	347	03	41	* * *	221	* * *	191.1	* * *	20	* * *	173.3	* * *	LMI	* * *
* * *	1974	347	06	51	* * *		* * *	191.0	* * *	50	* * *	173.0	* * *	221	* * *
* * *	1974	347	05	47	* * *		* * *	190.0	* * *	15	* * *	171.2	* * *	151	* * *
* * *	1974	347	10	19	* * *	LM2	* * *	189.7	* * *	00	* * *	166.3	* * *		* * *
* * *	1974	347	11	25	* * *		* * *	187.6	* * *	25	* * *	166.0	* * *		* * *
* * *	1974	347	12	11	* * *		* * *	187.2	* * *	00	* * *	164.0	* * *		* * *
* * *	1974	347	13	53	* * *		* * *	186.4	* * *	05	* * *	162.7	* * *		* * *
* * *	1974	347	13	36	* * *		* * *	186.0	* * *	55	* * *	162.3	* * *		* * *
* * *	1974	347	14	47	* * *		* * *	185.9	* * *	00	* * *	161.7	* * *	LMI	* * *
* * *	1974	347	14	04	* * *		* * *	185.8	* * *	10	* * *	160.4	* * *		* * *
* * *	1974	347	16	47	* * *		* * *	185.4	* * *	40	* * *	158.8	* * *		* * *
* * *	1974	347	16	55	* * *		* * *	184.3	* * *	25	* * *	157.5	* * *		* * *
* * *	1974	347	18	01	* * *		* * *	183.8	* * *	10	* * *	155.8	* * *	LM2	* * *
* * *	1974	347	21	28	* * *	LM1	* * *	182.0	* * *	50	* * *	155.6	* * *		* * *
* * *	1974	347	21	43	* * *	221	* * *	181.9	* * *	00	* * *	154.9	* * *	221	* * *
* * *	1974	347	21	50	* * *		* * *	181.8	* * *	20	* * *	152.7	* * *	142	* * *
* * *	1974	347	22	16	* * *	151	* * *	181.6	* * *	05	* * *	152.2	* * *	152	* * *
* * *	1974	347	23	27	* * *		* * *	181.0	* * *	10	* * *	152.0	* * *	151	* * *
* * *	1974	348	00	08	* * *		* * *	180.7	* * *	15	* * *	150.6	* * *		* * *
* * *	1974	348	01	17	* * *	LM2	* * *	180.1	* * *	40	* * *	150.2	* * *	LM	* * *
* * *	1974	348	03	22	* * *		* * *	179.0	* * *		* * *	.	* * *		* * *

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1975	103	15	07	15	152.7	LMI	1975	106	17	20	00	114.9	LM
1975	103	23	15	30	148.5	LMI	1975	106	17	25	20	114.8	LM
1975	104	00	03	00	148.1	LM2	1975	106	17	29	30	114.8	LM
1975	104	09	55	45	143.1	109	1975	106	17	31	20	114.8	LM
1975	104	10	08	05	143.0	142	1975	106	17	43	30	114.7	LM
1975	104	11	10	00	142.5	141	1975	106	17	48	30	114.7	LM
1975	104	13	12	40	141.4	139	1975	106	17	50	10	114.6	LM
1975	104	17	02	45	139.5	139	1975	106	17	57	30	114.6	LM
1975	104	20	29	00	137.7	LMI	1975	106	17	58	10	114.6	LM
1975	105	05	50	15	133.0	133	1975	106	18	04	50	114.5	LM
1975	105	07	04	00	132.3	109	1975	106	18	09	05	114.5	LM
1975	105	11	27	15	130.1	124	1975	106	18	13	00	114.4	LM
1975	105	16	44	30	127.6	124	1975	106	18	21	00	114.4	LM
1975	105	20	44	45	125.4	LM2	1975	106	18	25	50	114.3	LM
1975	105	22	51	05	125.3	LM2	1975	106	18	29	20	114.3	LM
1975	105	22	41	50	124.4	LM2	1975	106	18	38	40	114.2	LM
1975	106	03	23	00	122.0	LM1	1975	106	18	44	10	114.2	LM
1975	106	06	44	00	120.3	109	1975	106	18	48	00	114.1	LM
1975	106	08	24	40	119.4	116	1975	106	19	01	00	114.0	LM
1975	106	13	05	10	117.1	LM1	1975	106	19	12	10	113.9	LM
1975	106	13	20	40	116.9	LM2	1975	106	19	19	50	113.9	LM
1975	106	14	37	00	116.8	LM1	1975	106	19	24	10	113.8	LM
1975	106	14	55	30	116.1	LM2	1975	106	19	36	20	113.6	LM
1975	106	14	57	15	116.1	LM	1975	106	19	49	05	113.6	LM
1975	106	15	04	00	116.0	LM	1975	106	20	57	40	113.5	LM
1975	106	15	08	30	116.0	LM	1975	106	20	03	00	113.3	LM
1975	106	15	11	00	116.0	LM	1975	106	20	23	30	113.3	LM
1975	106	15	13	10	115.9	LM	1975	106	20	26	30	113.2	LM
1975	106	15	19	20	115.9	LM	1975	106	20	39	40	113.2	LM
1975	106	15	25	30	115.9	LM	1975	106	20	45	40	113.2	LM
1975	106	15	36	30	115.8	LM	1975	106	21	04	40	113.0	LM
1975	106	15	41	00	115.7	LM	1975	106	21	09	45	112.9	LM
1975	106	15	46	55	115.7	LM	1975	106	21	22	10	112.8	LM
1975	106	15	52	15	115.6	LM	1975	106	21	38	20	112.7	LM
1975	106	16	03	15	115.5	LM1	1975	106	21	49	20	112.6	112
1975	106	16	08	00	115.5	LM1	1975	106	22	34	00	112.2	109
1975	106	16	15	55	115.4	LM	1975	106	22	32	50	111.7	
1975	106	16	20	40	115.4	LM	1975	107	01	26	45	110.8	
1975	106	16	26	00	115.3	LM	1975	107	02	01	30	110.5	
1975	106	16	31	40	115.3	LM	1975	107	03	00	45	109.0	
1975	106	16	36	40	115.2	LM	1975	107	04	14	30	109.3	
1975	106	16	42	15	115.2	LM	1975	107	04	17	25	109.3	109
1975	106	16	47	40	115.1	LM	1975	107	04	49	50	109.0	
1975	106	16	52	40	115.1	LM	1975	107	05	20	05	108.8	
1975	106	16	58	25	115.0	LM	1975	107	05	32	00	108.7	
1975	106	17	04	45	114.9	LM	1975	107	05	33	45	108.7	
1975	106	17	09	00	114.9	LM	1975	107	05	37	15	107.7	
1975	106	17	14	50	114.9	LM							

YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1975	107	07	49	30	107.5	LM
1975	107	07	51	45	107.5	LM
1975	107	08	54	40	107.0	LM
1975	107	09	47	50	106.8	LM
1975	107	09	48	00	106.5	LM
1975	107	10	11	10	106.3	LM
1975	107	10	55	30	105.9	LM
1975	107	11	25	50	105.7	LM
1975	107	11	47	30	105.5	LM
1975	107	11	51	00	105.3	LM
1975	107	12	11	20	105.1	LM
1975	107	12	39	00	104.8	LM
1975	107	13	09	15	104.7	LM
1975	107	13	27	10	104.7	LM
1975	107	13	42	30	104.5	LM
1975	107	14	16	40	104.2	LM
1975	107	15	57	50	103.4	LM
1975	107	16	03	10	103.3	LM
1975	107	16	30	15	103.1	LM
1975	107	17	00	50	102.8	LM
1975	107	17	43	50	102.5	LM
1975	107	18	01	20	102.3	LM
1975	107	18	14	20	102.2	LM
1975	107	18	17	50	102.2	LM
1975	107	18	35	20	102.0	LM
1975	107	13	30	30	092.4	LM2
1975	107	12	33	40	093.0	LM
1975	107	12	16	40	093.6	LM
1975	107	11	12	50	093.7	LM
1975	107	10	55	15	096.4	LM
1975	107	05	38	00	096.9	LM
1975	107	04	47	25	097.7	LM2
1975	108	02	02	00	098.2	LM
1975	108	01	49	05	098.4	LM
1975	107	22	01	45	100.3	LM
1975	107	21	49	45	100.6	LM
1975	107	21	21	50	100.8	LM
1975	107	21	03	15	100.9	LM
1975	107	20	54	45	100.9	LM
1975	107	20	48	30	101.0	LM
1975	107	20	38	40	101.1	LM
1975	107	20	24	05	101.1	LM
1975	107	20	21	05	101.2	LM
1975	107	19	17	20	101.4	LM
1975	107	19	48	30	101.5	LM
1975	107	19	40	30	101.6	LM
1975	107	19	26	30	101.8	LM
1975	107	19	03	05	101.8	LM

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1975	116	15	20	35	354.1	355	1975	117	09	08	15	345.0	340
1975	116	15	32	30	354.0	355	1975	117	09	19	35	344.9	340
1975	116	15	34	45	353.9	354	1975	117	10	06	25	344.5	340
1975	116	15	47	00	353.8	354	1975	117	10	30	25	344.3	340
1975	116	16	06	10	353.7	354	1975	117	11	23	00	343.9	340
1975	116	16	08	25	353.1	354	1975	117	11	36	40	343.8	340
1975	116	17	10	40	353.0	348	1975	117	11	40	45	343.8	340
1975	116	17	18	50	352.7	348	1975	117	12	38	55	343.3	340
1975	116	18	19	00	352.5	348	1975	117	12	39	50	343.2	340
1975	116	18	23	10	352.5	348	1975	117	13	28	50	342.8	340
1975	116	18	37	00	352.4	348	1975	117	13	32	45	342.8	340
1975	116	18	38	40	352.4	348	1975	117	13	48	40	342.7	341
1975	116	19	03	55	352.2	348	1975	117	13	54	30	342.6	342
1975	116	19	08	30	352.1	348	1975	117	14	03	05	342.5	342
1975	116	19	16	05	352.1	348	1975	117	14	06	20	342.5	342
1975	116	19	17	30	352.1	348	1975	117	15	17	10	341.9	343
1975	116	19	24	05	352.0	348	1975	117	15	38	55	341.7	343
1975	116	19	27	10	351.8	348	1975	117	15	55	40	341.6	343
1975	116	19	51	30	351.7	348	1975	117	16	03	40	341.5	343
1975	116	20	05	20	351.6	348	1975	117	16	06	00	341.5	343
1975	116	20	17	10	351.5	348	1975	117	16	10	30	341.5	343
1975	116	20	22	00	351.5	348	1975	117	16	13	30	341.3	343
1975	116	20	38	50	351.4	348	1975	117	17	30	05	341.0	343
1975	116	21	51	30	351.3	348	1975	117	17	01	05	340.8	337
1975	116	21	09	00	349.8	348	1975	117	17	28	05	340.8	337
1975	116	23	41	55	349.5	348	1975	117	18	46	30	340.7	337
1975	117	00	18	35	349.0	348	1975	117	18	09	50	340.5	337
1975	117	01	24	30	349.0	348	1975	117	18	37	00	340.2	337
1975	117	01	36	40	348.9	344	1975	117	18	40	35	340.2	337
1975	117	02	36	15	348.4	344	1975	117	19	22	45	339.8	339
1975	117	02	56	55	348.2	344	1975	117	19	28	35	339.8	339
1975	117	03	21	20	348.0	344	1975	117	20	17	50	339.4	339
1975	117	03	52	50	347.7	344	1975	117	20	35	15	339.2	339
1975	117	04	36	00	347.3	344	1975	117	21	39	30	339.2	339
1975	117	05	10	40	347.0	344	1975	117	21	25	50	338.8	338
1975	117	05	20	00	347.0	344	1975	117	21	35	50	338.7	338
1975	117	05	39	10	346.8	344	1975	117	22	31	25	338.2	338
1975	117	05	55	00	346.7	344	1975	117	22	32	25	338.2	338
1975	117	05	57	50	346.6	344	1975	117	23	02	50	338.0	336
1975	117	06	02	35	346.6	344	1975	118	00	25	00	337.3	336
1975	117	06	34	30	346.3	344	1975	118	00	52	40	337.1	336
1975	117	06	49	05	346.2	344	1975	118	00	59	10	337.0	336
1975	117	07	16	10	346.0	344	1975	118	01	18	00	336.8	338
1975	117	07	27	45	345.9	344	1975	118	01	31	20	336.7	335
1975	117	07	27	50	345.9	344	1975	118	01	39	50	336.7	335
1975	117	08	12	00	345.5	344	1975	118	01	45	45	336.6	334
1975	117	08	13	40	345.5	344	1975	118	01	51	50	336.6	334
1975	117	08	23	15	345.4	344	1975	118	02	40	30	336.1	333
1975	117	08	38	45	345.3	344	1975	118	02	52	00	336.0	333
1975	117	08	50	05	345.2	344	1975	118	03	03	00	336.0	333

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR
1975	118	03	53	10	335.5		1975	118	21	53	35	326.4	322	1975
1975	118	03	56	00	335.5		1975	118	21	56	25	326.4		1975
1975	118	04	21	05	335.2		1975	118	22	59	30	326.3		1975
1975	118	04	27	00	335.1		1975	118	22	03	30	326.3		1975
1975	118	04	43	00	335.0		1975	118	22	04	10	326.3		1975
1975	118	04	51	00	335.0	332	1975	118	22	31	00	326.1		1975
1975	118	05	56	30	335.0		1975	118	23	49	05	325.4		1975
1975	118	05	01	30	334.9		1975	118	23	44	05	325.4		1975
1975	118	05	04	35	334.7		1975	118	23	56	00	325.4		1975
1975	118	05	33	00	334.7	330	1975	119	00	01	00	325.3		1975
1975	118	06	02	45	334.4	33A	1975	119	00	16	00	325.2		1975
1975	118	06	11	15	334.4		1975	119	00	53	50	324.9		1975
1975	118	06	24	00	334.2		1975	119	00	58	05	324.8		1975
1975	118	07	22	30	333.8		1975	119	01	07	30	324.8		1975
1975	118	07	31	45	333.6	33C	1975	119	01	13	00	324.7		1975
1975	118	07	44	50	333.6		1975	119	01	27	55	324.6		1975
1975	118	08	07	30	333.4		1975	119	01	36	30	324.5		1975
1975	118	08	22	00	333.2		1975	119	01	45	40	324.4		1975
1975	118	08	36	50	333.1		1975	119	02	36	05	324.0		1975
1975	118	08	47	45	333.0		1975	119	02	53	00	323.8		1975
1975	118	08	58	10	333.0		1975	119	03	00	55	323.8		1975
1975	118	09	53	50	332.5		1975	119	03	03	35	323.8		1975
1975	118	12	09	40	331.3	33D	1975	119	03	10	50	323.7		1975
1975	118	13	09	50	330.8	33E	1975	119	03	56	50	323.3		1975
1975	118	13	49	00	330.5		1975	119	04	12	00	323.2		1975
1975	118	14	00	10	330.4		1975	119	04	17	10	323.1		1975
1975	118	15	27	00	329.7		1975	119	04	50	55	322.9		1975
1975	118	15	43	40	329.5		1975	119	05	02	00	322.8		1975
1975	118	15	49	00	329.5		1975	119	05	15	30	322.6		1975
1975	118	16	23	25	329.2		1975	119	05	29	45	322.5		1975
1975	118	16	33	25	329.1		1975	119	06	45	30	322.4		1975
1975	118	16	48	55	329.0	328	1975	119	06	07	10	322.2		1975
1975	118	16	53	00	328.9	329	1975	119	06	19	00	322.1		1975
1975	118	17	16	15	328.9	325	1975	119	06	47	00	322.1		1975
1975	118	17	31	00	328.9		1975	119	07	11	05	322.2		1975
1975	118	18	31	00	328.6		1975	119	07	23	05	322.2		1975
1975	118	18	20	50	328.2		1975	119	07	31	30	321.1		1975
1975	118	18	38	20	328.0		1975	119	07	35	25	321.5		1975
1975	118	19	28	40	327.6		1975	119	07	44	00	321.4		1975
1975	118	19	40	00	327.5		1975	119	07	49	20	321.4		1975
1975	118	19	45	50	327.5	327	1975	119	08	05	15	321.2		1975
1975	118	20	55	00	327.4		1975	119	08	11	15	321.2		1975
1975	118	20	13	10	327.2		1975	119	08	32	55	321.0		1975
1975	118	20	29	00	327.1		1975	119	08	35	10	321.0		1975
1975	118	20	29	45	327.1	310	1975	119	08	45	40	320.7		1975
1975	118	20	41	00	327.0		1975	119	09	04	00	320.7		1975
1975	118	20	53	00	326.9		1975	119	09	17	00	320.6		1975
1975	118	21	03	00	326.8		1975	119	09	38	45	320.4		1975
1975	118	21	07	30	326.8		1975	119	09	57	20	320.3		1975
1975	118	21	27	00	326.6		1975	119	10	05	35	320.2		1975
1975	118	21	39	30	326.5	324	1975	119	10	17	15	320.1		1975

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1975	119	10	19	15	320.1	318	1975	120	01	03	00	312.6	31D
1975	119	10	59	50	319.7	318	1975	120	01	43	30	312.3	31D
1975	119	11	01	35	319.5	318	1975	120	01	45	40	312.2	31D
1975	119	11	31	15	319.4	318	1975	120	02	08	45	312.0	31D
1975	119	11	43	30	319.3	313	1975	120	02	50	15	311.7	31D
1975	119	11	54	00	319.3	313	1975	120	02	56	10	311.6	31D
1975	119	12	56	00	318.8	317	1975	120	02	59	15	311.6	LM2
1975	119	12	46	00	318.5	317	1975	120	03	02	55	311.6	LM2
1975	119	13	24	05	318.4	317	1975	120	03	04	15	311.5	LM2
1975	119	13	41	30	318.2	317	1975	120	03	08	10	311.5	LM2
1975	119	14	00	45	318.2	317	1975	120	03	17	45	311.4	LM2
1975	119	15	00	00	317.2	317	1975	120	03	21	20	311.4	LM2
1975	119	16	01	40	317.2	317	1975	120	03	26	30	311.4	LM2
1975	119	16	03	00	317.0	316	1975	120	03	31	00	311.3	31F
1975	119	16	37	05	316.9	316	1975	120	03	38	10	311.3	31G
1975	119	16	46	00	316.8	316	1975	120	03	49	55	311.2	31G
1975	119	17	53	00	316.8	316	1975	120	04	02	00	311.1	31G
1975	119	17	30	00	316.4	314	1975	120	04	13	00	311.1	31G
1975	119	17	35	00	316.4	314	1975	120	04	25	00	310.9	31G
1975	119	18	29	35	315.9	314	1975	120	04	29	00	310.9	31G
1975	119	18	40	10	315.8	314	1975	120	04	31	30	310.8	31G
1975	119	19	17	00	315.5	314	1975	120	04	38	00	310.8	31G
1975	119	19	35	50	315.4	314	1975	120	04	45	00	310.7	31G
1975	119	19	48	45	315.3	314	1975	120	05	14	00	310.7	31G
1975	119	19	51	50	315.2	314	1975	120	05	23	50	310.4	31H
1975	119	19	58	25	315.2	314	1975	120	05	32	00	310.3	31H
1975	119	20	03	05	315.1	315	1975	120	06	04	15	310.1	31H
1975	119	20	06	20	315.1	315	1975	120	06	13	30	310.0	306
1975	119	20	23	40	315.0	315	1975	120	06	38	00	309.8	306
1975	119	20	25	00	315.0	315	1975	120	06	56	00	309.6	306
1975	119	20	40	40	314.8	315	1975	120	07	01	10	309.6	LM2
1975	119	20	47	00	314.8	315	1975	120	07	04	50	309.6	LM2
1975	119	20	51	30	314.7	LM2	1975	120	07	14	15	309.5	303
1975	119	20	57	20	314.7	311	1975	120	07	32	10	309.3	303
1975	119	20	57	45	314.7	311	1975	120	07	40	00	309.2	303
1975	119	21	49	15	314.2	311	1975	120	07	41	00	309.2	303
1975	119	22	17	45	314.0	312	1975	120	08	54	00	309.1	303
1975	119	22	50	40	313.7	312	1975	120	08	09	00	309.0	309
1975	119	23	16	00	313.5	312	1975	120	08	26	05	308.8	309
1975	119	23	27	50	313.4	312	1975	120	08	35	30	308.8	30A
1975	120	00	01	01	313.4	312	1975	120	08	54	40	308.6	30A
1975	120	00	12	30	313.0	313	1975	120	08	59	50	308.6	308
1975	120	00	16	15	313.0	313	1975	120	09	14	30	308.4	308
1975	120	00	16	30	313.0	313	1975	120	09	21	10	308.4	LM1
1975	120	00	21	50	312.9	31E	1975	120	09	45	20	308.2	307
1975	120	00	32	00	312.8	31E	1975	120	09	51	00	308.1	307
1975	120	00	36	00	312.8	31E	1975	120	09	55	15	308.1	307
1975	120	00	37	00	312.7	31E	1975	120	10	00	35	308.1	307
1975	120	00	55	40	312.7	31E	1975	120	10	05	25	308.0	307

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE
1975	120	10	35	05	307.8		1975	120	10	11	40	303.4	30E
1975	120	10	58	00	307.6		1975	120	19	18	05	303.3	30E
1975	120	11	19	00	307.4		1975	120	19	28	30	303.2	
1975	120	11	20	55	307.4	LM2	1975	120	19	18	00	303.2	30D
1975	120	11	23	35	307.3	LM2	1975	120	19	34	05	303.2	
1975	120	11	32	00	307.3		1975	120	19	36	45	303.2	
1975	120	11	33	35	307.3		1975	120	19	40	20	303.2	
1975	120	11	48	00	307.1	30A	1975	120	19	45	30	303.1	303
1975	120	11	39	40	306.7		1975	120	20	04	00	302.9	
1975	120	12	45	55	306.7		1975	120	20	05	50	302.9	30H
1975	120	13	07	10	306.5	305	1975	120	20	23	00	302.8	LM2
1975	120	13	12	00	306.4	LM1	1975	120	20	33	35	302.7	30G
1975	120	13	12	55	306.4		1975	120	20	39	05	302.6	
1975	120	13	20	05	306.3		1975	120	21	05	00	302.4	
1975	120	13	24	00	306.3		1975	120	21	11	00	302.4	
1975	120	13	35	00	306.2		1975	120	21	15	10	302.3	30F
1975	120	13	41	25	306.2	300	1975	120	21	18	45	302.3	
1975	120	14	56	00	305.6	301	1975	120	21	23	25	302.3	
1975	120	15	03	15	305.5		1975	120	21	41	50	302.1	30I
1975	120	15	07	15	305.5		1975	120	21	46	25	302.1	
1975	120	15	18	30	305.4	30A	1975	120	21	50	25	302.0	
1975	120	15	29	35	305.3	30B	1975	120	21	57	00	302.0	
1975	120	15	51	55	305.1	304	1975	120	22	01	55	302.0	
1975	120	16	12	40	304.9		1975	120	22	07	30	301.9	
1975	120	16	24	35	304.8		1975	120	22	11	25	301.9	
1975	120	16	33	30	304.7		1975	120	22	14	05	301.8	
1975	120	16	41	35	304.7		1975	120	22	16	50	301.8	
1975	120	16	59	45	304.5		1975	120	22	22	40	301.8	
1975	120	17	14	00	304.4		1975	120	22	24	30	301.8	
1975	120	17	47	10	304.1		1975	120	22	27	00	301.7	
1975	120	17	52	45	304.1		1975	120	22	31	10	301.7	
1975	120	18	24	00	303.8		1975	120	22	43	50	301.6	
1975	120	18	28	00	303.8		1975	120	22	55	55	301.5	
1975	120	18	36	30	303.7		1975	120	23	09	00	301.4	LM2
1975	120	18	38	20	303.7	30A	1975	120	23	10	00	301.4	
1975	120	18	49	30	303.6		1975	120	23	26	20	301.2	
1975	120	18	52	00	303.6	LM2	1975	120	23	39	30	301.1	
1975	120	18	56	25	303.5	30C	1975	120	23	48	05	301.0	
1975	120	19	02	00	303.5		1975	120	23	56	30	301.0	
1975	120	19	06	00	303.5		1975	121	00	36	00	300.6	

* P.19

YEAR	DAY	HR	MNI	AMP	LONG	TYPE	YEAR	DAY	HR	MNI	AMP	LONG	TYPE
* 1975	* 132	* 20	* 04	* 35	* 156.3		* 1975	* 134	* 17	* 32	* 50	* 133.1	* 109
* 1975	* 133	* 00	* 14	* 30	* 154.1		* 1975	* 135	* 06	* 18	* 10	* 126.6	
* 1975	* 133	* 00	* 31	* 50	* 154.0		* 1975	* 135	* 09	* 32	* 05	* 124.9	* 124
* 1975	* 133	* 03	* 41	* 00	* 152.4	LM1	* 1975	* 135	* 10	* 45	* 50	* 124.3	
* 1975	* 133	* 13	* 50	* 00	* 147.2	109	* 1975	* 135	* 16	* 27	* 00	* 124.1	
* 1975	* 133	* 14	* 23	* 10	* 146.6		* 1975	* 135	* 18	* 52	* 45	* 120.2	LM2
* 1975	* 133	* 14	* 39	* 50	* 146.8		* 1975	* 135	* 19	* 09	* 00	* 120.0	109
* 1975	* 133	* 14	* 48	* 00	* 146.7	142	* 1975	* 135	* 21	* 06	* 50	* 119.0	LM1
* 1975	* 133	* 21	* 13	* 25	* 143.4	LM2	* 1975	* 135	* 21	* 59	* 50	* 118.6	116
* 1975	* 133	* 23	* 28	* 00	* 142.3	141	* 1975	* 136	* 03	* 32	* 00	* 115.8	LM1
* 1975	* 133	* 23	* 30	* 30	* 142.3	139	* 1975	* 136	* 06	* 07	* 20	* 114.4	112
* 1975	* 134	* 01	* 10	* 30	* 141.4		* 1975	* 136	* 10	* 15	* 40	* 112.3	
* 1975	* 134	* 05	* 42	* 25	* 139.1	133	* 1975	* 136	* 10	* 28	* 30	* 112.2	112
* 1975	* 134	* 09	* 18	* 58	* 137.3	LM1	* 1975	* 136	* 10	* 55	* 25	* 112.2	

```

* 1975 * 145 * 02 * 58 * 10 * 006.2 * TYPE * * * * * P.20
* 1975 * 145 * 02 * 59 * 25 * 006.2 * * * * *
* 1975 * 145 * 04 * 10 * 50 * 005.6 * * * * *
* 1975 * 145 * 04 * 18 * 40 * 005.6 * * * * *
* 1975 * 145 * 04 * 27 * 50 * 005.5 * * * * *
* 1975 * 145 * 04 * 33 * 55 * 005.4 * * * * *
* 1975 * 145 * 04 * 51 * 40 * 005.3 * * * * *
* 1975 * 145 * 05 * 03 * 20 * 005.2 * * * * *
* 1975 * 145 * 05 * 20 * 50 * 005.0 * * * * *
* 1975 * 145 * 05 * 41 * 00 * 004.9 * * * * *

* 1975 * 145 * 06 * 20 * 50 * 004.5 * LONG * * * * *
* 1975 * 145 * 06 * 31 * 50 * 004.4 * * * * *
* 1975 * 145 * 06 * 49 * 20 * 004.3 * * * * *
* 1975 * 145 * 07 * 00 * 45 * 004.2 * * * * *
* 1975 * 145 * 07 * 16 * 25 * 004.1 * * * * *
* 1975 * 145 * 07 * 23 * 05 * 004.0 * * * * *
* 1975 * 145 * 07 * 37 * 10 * 003.9 * * * * *
* 1975 * 145 * 07 * 56 * 50 * 003.7 * * * * *
* 1975 * 145 * 08 * 03 * 05 * 003.7 * * * * *
* 1975 * 145 * 08 * 10 * 00 * 003.6 * * * * *

```