## DIURNAL POLAR MOTION

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## PAUL MCCLURE

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# DIURNAL POLAR MOTION 

Paul McClure<br>Geodynamics Branch Geodynamics Programs Division

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Goddard Space Flight Center
Greenbelt, Maryland

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Paul McClure


#### Abstract

An analytical theory is developed to describe diurnal polar motion in the Earth which arises as a forced response due to lunisolar torques and tidal deformation. Doodson's expansion of the tide generating potential is used to represent the lunisolar torques. Both the magnitudes and the rates of change of perturbations in the Earth's inertia tensor are included in the dynamical equations for the polar motion so as to account for rotational and tidal deformation.

It is found that in a deformable Earth with Love's number $k=0.29$, the angular momentum vector departs by as much as 20 cm from the rotation axis rather than remaining within 1 or 2 cm as it would in a rigid Earth. This 20 cm separation is significant in the interpretation of sub-meter polar motion observations because it necessitates an additional coordinate transformation in order to remove what would otherwise be a 20 cm error source in the conversion between inertial and terrestrial reference systems.


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## DIURNAL POLAR MOTION

## 1. INTRODUCTION

Interaction of the lunar and solar tidal potentials with the Earth's equatorial bulge generates torques on the Earth. Besides producing the well known phenomena of astronomical precession and nutation the lunisolar torques cause the rotation pole to travel within the Eaxth in a nearly diurnal epicycle which has a radius that varies from a few centimeters to a maximum of 62 cm . The diurnal motion of the rotation pole is superimposed upon longer period motions consisting mainly of the 14 month Chandler wobble and the annual and semiannual polar motion.

Unlike the Chandler, annual and semiannual motions the diurnal motion of the pole has not been observed conclusively because of its comparatively small amplitude and high frequency. The astronomical methods for observing polar motion have uncertainties that are about equal to the amplitude of the diurnal polar motion and are able to produce pole positions only at 2 to 5 day intervals. Laser tracking of artificial satellites is now able to give pole positions at intervals of 6 hours [Smith et al., 1972], thus providing an opportunity to observe diurnal motion of the rotation axis within the Earth. The satellite observations have noise levels of about 1 m and it is expected that observational uncertainties can be reduced to 10 cm in the future.

In order to interpret polar motion observations with sub-meter noise levels it is necessary to model the diurnal motion of the pole. Woolard [1953] derived expressions for the diurnal polar motion in a rigid Earth. His results do not include the effects of rotational and tidal deformation and the effect of the lunar and solar mean motion upon the coefficients in the diurnal polar motion terms is neglected. Melchior and Georis [1968] use Doodson's [1922] expansion of the tide generating potential in order to obtain expressions for the lunisolar torques. The effect of mean motion of the disturbing bodies upon the diurnal polar motion is included as a second order correction. Their dynamical equations for the polar motion are for the case of a rigid Earth and do not include the effects of rotational and tidal deformation.

The theory of diurnal polar motion presented here is for the case of a deformable Earth. The response of the Earth to deforming potentials is characterized by Love's number k [Love, 1911]. Terms involving both the magnitudes and the rates of change of perturbations in the Earth's inertia tensor are included in the dynamical equations for the polar motion so as to account for the rotational and
tidal deformation. Doodson's expansion of the tide generating potential is used to represent the lunisolar torques, and mean motion of the tide generating bodies is included in the solution. The polar motion due to lunisolar torques is combined analytically with that due to rotational and tidal deformation in order to form a single set of coefficients for the diurnal polar motion. These coefficients along with the corresponding tesseral diurnal tidal arguments are arranged in tabular form so as to permit rapid computer evaluation of the diurnal polar motion at any instant of time and for a given Love number and set of astronomical constants.

In addition to the results giving the motion of the rotation pole, solutions are obtained for the diurnal motion within the Earth of the angular momentum vector or principal axis of inertia. It is found that, in a deformable Earth, the angular momentum vector departs by as much as 20 cm from the rotation axis rather than remaining within 1 or 2 cm as it would if the Earth were rigid. The pre-cession-nutation theory represented in the Explanatory Supplement to the American Ephemeris and Nautical Almanac [1961, p. 44] includes only the 1 to 2 cm rigid-Earth correction for the departure of the rotation axis from the angular momentum vector. An actual separation of 20 cm is significant in the interpretation of sub-meter polar motion observations because it necessitates an additional coordinate transformation in order to remove what would otherwise be a 20 cm error source in the conversion between inertial and terrestrial reference systems. Three alternative methods for making the additional coordinate transformation are discussed in detail. Each method makes use of tabulated coefficients and arguments and can be readily programmed for use on an automatic computer.

Woolard's theory of precession and nutation gives the best available representation for the direction of the Earth's angular momentum vector in space. The diurnal polar motion theory presented here must be used in conjunction with such a precession-nutation model in order to give the orientation an observatory-fixed terrestrial reference frame. It is important that Woolard's results be understood in the context of the present development and, for this reason, a discussion of his method of solution is included.

## 2. POLAR MOTION DYNAMICS

The rotational motion of a general mass distribution $M$ is described by Liouville's equation

$$
\begin{equation*}
L_{i}=\dot{H}_{i}+\epsilon_{i j k} \omega_{j} H_{k} \tag{2.1}
\end{equation*}
$$

where

$$
\begin{gather*}
H_{i}=I_{i j} \omega_{j}+h_{i}  \tag{2.2}\\
I_{i j}=\int_{M}\left(x_{k} x_{k} \delta_{i j}-x_{i} x_{j}\right) d m  \tag{2.3}\\
h_{i}=\int_{M} \epsilon_{i j k} x_{j} \dot{x}_{k} d m \tag{2.4}
\end{gather*}
$$

Repeated indices indicate summation. The subscripts ( $i=1,2,3$ ) refer to a set of axes ( $x, y, z$ ) having their origin at the center of mass and an angular velocity with components $\omega_{i}, L_{i}$ and $H_{i}$ denote components of the net external torque and the angular momentum respectively. $I_{i j}$ is the inertia tensor. $h_{i}$ is the part of $H_{i}$ arising from motion relative to the $x, y, z$ system.

Although Liouville's equation is valid in any coordinate system, the axes shown in Figure (2.1) are especially useful for geophysical problems. The $x ; y$, $z$ system is attached to a set of terrestrial observatories in some prescribed manner. The $x$ and $y$ axes define the Earth's equator and the $z$ axis is placed so as to remain nearly aligned with the rotation axis.

The following perturbation scheme due to Munk and Macdonald [1960, p. 38] serves to simplify Liouville's equation.

$$
I_{i j}=\left[\begin{array}{ccc}
A+c_{11} & c_{12} & c_{13} \\
c_{12} & A+c_{22} & c_{23} \\
c_{13} & c_{23} & C+c_{33}
\end{array}\right]
$$



Figure 2.1. Terrestrial Coordinate System.

For the Earth, the quantities $\mathrm{m}_{\mathrm{i}}, \mathrm{h}_{\mathrm{i}} / \mathrm{C} \Omega$, and $\mathrm{c}_{\mathrm{ij}} / \mathrm{C}$ are of order $10^{-6}$ or smaller. Neglecting second order terms in the small quantities therefore produces first order equations for the polar motion that are exact to about one part in $10^{6}$. Substitution of (2.5) through (2.8) into (2.1) gives

$$
\begin{align*}
L_{1} & =A \dot{m}_{1} \Omega+\dot{c}_{13} \Omega+\dot{h}_{1} \\
& +m_{2}(C-A) \Omega^{2}-c_{23} \Omega^{2}-h_{2} \Omega  \tag{2.9}\\
L_{2} & =A \dot{m}_{2} \Omega+\dot{c}_{23} \Omega+\dot{h}_{2} \\
& -m_{1}(C-A) \Omega^{2}+c_{13} \Omega^{2}+h_{1} \Omega  \tag{2.10}\\
& L_{3}=\dot{c}_{33} \Omega+C \dot{m}_{3} \Omega+\dot{h}_{3} \tag{2.11}
\end{align*}
$$

Equations (2.9) and (2.10) are written as a single complex equation in the form

$$
\begin{equation*}
\dot{\mathrm{m}}=\mathrm{i} \sigma_{\mathbf{r}}(\mathrm{m}-\psi) \tag{2.12}
\end{equation*}
$$

where

$$
\begin{gather*}
i=\sqrt{-1}  \tag{2.13}\\
m=m_{1}+i m_{2}  \tag{2.14}\\
\sigma_{r}=(\mathrm{C}-\mathrm{A}) \Omega / \mathrm{A} \tag{2.15}
\end{gather*}
$$

The polar motion excitation function $\psi$ is

$$
\begin{align*}
\psi & =\frac{i L}{(C-A) \Omega^{2}}+\frac{c}{C-A}-\frac{i \dot{c}}{(C-A) \Omega} \\
& +\frac{h}{(C-A) \Omega}-\frac{i \dot{h}}{(C-A) \Omega^{2}} \tag{2.16}
\end{align*}
$$

where

$$
\begin{align*}
& \psi=\psi_{1}+i \psi_{2}  \tag{2.17}\\
& \mathrm{~L}=\mathrm{L}_{1}+\mathrm{i} \mathrm{~L}_{2}  \tag{2.18}\\
& \mathrm{c}=\mathrm{c}_{13}+\mathrm{i} \mathrm{c}_{23}  \tag{2.19}\\
& \mathrm{~h}=\mathrm{h}_{1}+\mathrm{i} \mathrm{~h}_{2} \tag{2.20}
\end{align*}
$$

Equation (2.11) is rewritten as

$$
\begin{equation*}
\dot{m}_{3}=\frac{L_{3}}{C \Omega}-\frac{\dot{c}_{33}}{C}-\frac{\dot{h}_{3}}{\mathrm{C} \Omega} \tag{2.21}
\end{equation*}
$$

The complex form of the polar motion equation makes solutions easier to visualize. For example, if there is no excitation then $\psi=0$ and the solution to (2.12) is

$$
\begin{equation*}
m=m_{0} e^{i \sigma_{r} t} \tag{2.22}
\end{equation*}
$$

The motion is a counterclockwise circular path about the origin as shown in Figure (2.2). This "free" or "Eulerian" motion has a period of 10 months. Its amplitude and phase are determined by the initial condition

$$
\begin{equation*}
m(0)=m_{0} \tag{2.23}
\end{equation*}
$$

where $m_{0}$ is a complex constant of integration.


Figure 2.2. Eulerian Motion of the Pole.

The first order expansion of the angular momentum vector is

$$
\begin{gather*}
\mathrm{H}=\mathrm{A} \Omega \mathrm{~m}+\Omega \mathrm{c}+\mathrm{h}  \tag{2.24}\\
\mathrm{H}_{3}=\mathrm{C} \Omega\left(1+\mathrm{m}_{3}\right)+\mathrm{c}_{33} \Omega+\mathrm{h}_{3} \tag{2.25}
\end{gather*}
$$

where

$$
\begin{equation*}
\mathrm{H}=\mathrm{H}_{1}+\mathbf{i} \mathrm{H}_{2} \tag{2.26}
\end{equation*}
$$

The direction cosines of the rotation axis and the angular momentum vector are

$$
\begin{equation*}
\frac{\vec{\omega}}{|\vec{\omega}|}=\left(m_{1}, m_{2}, 1\right) \tag{2.27}
\end{equation*}
$$

$$
\begin{equation*}
\frac{\overrightarrow{\mathrm{H}}}{|\overrightarrow{\mathrm{H}}|}=\left(\frac{\mathrm{H}_{1}}{\mathrm{C} \Omega}, \frac{\mathrm{H}_{2}}{\mathrm{C} \Omega}, 1\right) \tag{2.28}
\end{equation*}
$$

in which second order terms are omitted. The complex numbers $m$ and $H / C \Omega$ represent equatorial projections of unit vectors directed along the rotation axis and the angular momentum vector respectively.

## 3. POLAR MOTION KINEMATICS

It is necessary to define a set of coordinates to relate the terrestrial axes $\mathrm{x}, \mathrm{y}, \mathrm{z}$ to a set of inertial axes X, Y, Z. For this purpose Woolard [1953, p. 15] defines the Euler angles shown in Figure (3.1). The XY plane is that of the ecliptic at a prescribed epoch. The angle $\psi$ is the longitude of the equinox $P_{1 E}$ and is measured in the ecliptic of epoch eastward from $X$. The obliquity of the ecliptic is denoted by $\theta$. The Earth's diurnal rotation is described by the angle $\phi$, which is measured eastward from $\Upsilon_{1 \mathrm{E}}$ to the x axis.


Figure 3.1. Euler Angles.

The rates of change of the Euler angles are related to the angular velocity components by Euler's kinematic equations,

$$
\begin{align*}
\dot{\psi} \sin \theta & =-\omega_{1} \sin \phi-\omega_{2} \cos \phi  \tag{3.1}\\
\dot{\theta} & =-\omega_{1} \cos \phi+\omega_{2} \sin \phi  \tag{3.2}\\
\dot{\phi} & =\omega_{3}-\dot{\psi} \cos \theta \tag{3.3}
\end{align*}
$$

Equations (3.1) and (3.2) are expressed in complex form as

$$
\begin{equation*}
\dot{\theta}+\mathbf{i} \dot{\psi} \sin \theta=-\Omega \mathrm{m} \mathrm{e}^{\mathrm{i} \phi} \tag{3.4}
\end{equation*}
$$

A direction cosine vector $\overrightarrow{\mathrm{u}}$ which is nearly aligned with the z axis can be described by the perturbations $\delta \psi$ and $\delta \theta$ in the Euler angles as shown in Figure (3.2). In order to relate the Euler angle perturbations to the direction cosines $u_{i}$, it is convenient to introduce the node axes $\xi, \eta, \zeta$ shown in Figure (3.3). The $\xi$ axis points toward the equinox $P_{1 E}$ and the $\eta$ axis lies in the equator $90^{\circ}$ to the east. The $\zeta$ axis coincides with z. Neglecting second order terms in $\delta \psi$ and $\delta \theta$, the node axis components of $\overrightarrow{\mathbf{u}}$ are

$$
\begin{gather*}
\mathrm{u}_{\xi}=-\delta \psi \sin \theta  \tag{3.5}\\
u_{\eta}=\delta \theta \tag{3.6}
\end{gather*}
$$

The Euler angle perturbations are defined in the sense

$$
\begin{align*}
& \delta \psi=\psi_{\overrightarrow{\mathbf{u}}}-\psi  \tag{3.7}\\
& \delta \theta=\theta_{\overrightarrow{\mathbf{u}}}-\theta \tag{3.8}
\end{align*}
$$

The $x$ and $y$ components of $\vec{u}$ are related to the node axis components by

$$
\begin{equation*}
u_{\xi}+i u_{\eta}=e^{i \phi}\left(u_{x}+i u_{y}\right) \tag{3.9}
\end{equation*}
$$

Combining (3.9), (3.5), and (3.6) gives

$$
\begin{equation*}
\delta \theta+i \delta \psi \sin \theta=-i e^{i \phi}\left(u_{x}+i u_{y}\right) \tag{3.10}
\end{equation*}
$$



Figure 3.2. Euler Angle Perturbations


Figure 3.3. Node Axis Components of a Unit Vector Related to Euler Angle Perturbations.

## 4. LUNISOLAR TORQUES ON THE EARTH

The torque exerted on the Earth by a point-mass disturbing body with geocentric position $\overrightarrow{\mathbf{r}}_{\mathrm{d}}$ and mass $\mathrm{m}_{\mathrm{d}}$ is

$$
\begin{equation*}
L_{i}=-\epsilon_{i j k} m_{d} r_{d j}\left(\frac{\partial V}{\partial r_{d k}}\right) \tag{4.1}
\end{equation*}
$$

The Earth's gravitational potential $V$ is defined as the integral

$$
\begin{equation*}
V=\int \frac{G d m}{\left|\vec{r}_{d}-\vec{r}\right|} \tag{4.2}
\end{equation*}
$$

over all mass elements dm located at positions $\overrightarrow{\mathrm{r}}$ within the Earth. V is expanded in terms of Legendre functions as

$$
\begin{align*}
V & =\frac{G m_{E}}{r_{d}}\left\{1-\sum_{n=2}^{\infty} J_{n}\left(\frac{a_{E}}{r_{d}}\right)^{n} P_{n}\left(\sin \phi_{d}\right)\right. \\
& +\sum_{n=2}^{\infty} \sum_{m=1}^{n}\left(\frac{a_{E}}{r_{d}}\right)^{n} P_{n}^{m}\left(\sin \phi_{d}\right)\left(C_{n m} \cos m \lambda_{d}\right. \\
& \left.\left.+S_{n m} \sin m \lambda_{d}\right)\right\} \tag{4.3}
\end{align*}
$$

The coordinates $r_{d}, \phi_{d}$, and $\lambda_{d}$ of the disturbing body are shown in Figure (4.1). $\mathbf{P}_{\mathrm{n}}^{\mathrm{m}}(\mu)$ is the Legendre associated function of degree n and order m defined by

$$
\begin{equation*}
P_{n}^{m}(\mu)=\left(1-\mu^{2}\right)^{m / 2} \frac{d}{d \mu^{m}}\left(P_{n}(\mu)\right) \tag{4.4}
\end{equation*}
$$

where the $P_{n}(\mu)$ are Legendre's polynomials,

$$
\begin{equation*}
\mathbf{P}_{\mathbf{n}}(\mu)=\frac{1}{\mathrm{n}!} \frac{\mathrm{d}^{n}}{\mathrm{~d} \mu^{n}}\left(\frac{\mu^{2}-1}{2}\right)^{n} \tag{4.5}
\end{equation*}
$$



Figure 4.1. Disturbing Body Coordinates

The torque components in the $x, y, z$ system are

$$
\begin{align*}
& L_{1}=-m_{d}\left(\frac{\partial V}{\partial \phi_{d}}\right) \sin \lambda_{d}+\frac{m_{d}}{\cos \phi_{d}}\left(\frac{\partial V}{\partial \lambda_{d}}\right) \sin \phi_{d} \cos \lambda_{d} \\
& L_{2}=m_{d}\left(\frac{\partial V}{\partial \phi_{d}}\right) \cos \lambda_{d}+\frac{m_{d}}{\cos \phi_{d}}\left(\frac{\partial V}{\partial \lambda_{d}}\right) \sin \phi_{d} \sin \lambda_{d}  \tag{4.7}\\
& \mathbf{L}_{3}=-m_{d}\left(\frac{\partial V}{\partial \lambda_{d}}\right) \tag{4.8}
\end{align*}
$$

The second degree zonal term in $V$ is larger by a factor of $10^{3}$ than the higher degree zonal terms and the longitude dependent terms. Therefore, only the $\mathrm{J}_{2}$
term is retained for the purpose of computing torques. This simplifies the expressions for the equatorial torques $L_{1}$ and $L_{2}$ and results in zero torque about the z axis. The torques arising from the $\mathrm{J}_{2}$ term are

$$
\begin{gather*}
L_{1}=\frac{G m_{E}}{r_{d}} m_{d} \sin \lambda_{d}\left(\frac{a_{E}}{r_{d}}\right)^{2} J_{2} P_{2}^{1}\left(\sin \phi_{d}\right)  \tag{4.9}\\
L_{2}=-\frac{G m_{E}}{r_{d}} m_{d} \cos \lambda_{d}\left(\frac{a_{E}}{r_{d}}\right)^{2} J_{2} P_{2}^{1}\left(\sin \phi_{d}\right)  \tag{4.10}\\
L_{3}=0 \tag{4.11}
\end{gather*}
$$

In order to use expressions (4.9) and (4.10) in the analysis of polar motion dynamics, the indicated functions of the disturbing body coordinates must be known as explicit functions of time. The time dependence of the lunar and solar coordinates is contained in Equation (A.22) which represents Doodson's expansion of the tide generating potential. A formal spherical harmonic expansion of the tide generating potential is compared with (A.22) in order to determine the time dependence of the functions appearing in (4.9) and (4.10).

The tide generating potential is

$$
\begin{equation*}
U=\frac{G m_{d}}{r_{d}} \sum_{n=2}^{\infty}\left(\frac{r}{r_{d}}\right)^{n} P_{n}\left(\cos \gamma_{d}\right) \tag{4.12}
\end{equation*}
$$

The local zenith angle $\gamma_{d}$ of the disturbing body is given in terms of local latitude $\phi$ and longitude $\lambda$ by

$$
\begin{equation*}
\cos \gamma_{d}=\sin \phi \sin \phi_{d}+\cos \phi \cos \phi_{d} \cos \left(\lambda-\lambda_{d}\right) \tag{4.13}
\end{equation*}
$$

Equation (4.13) is substituted into (4.12) to obtain

$$
\begin{equation*}
U=\frac{G m_{d}}{r_{d}} \sum_{n=2}^{\infty} \sum_{m=0}^{n}\left(\frac{r}{r_{d}}\right)^{n} W_{n m} P_{n}^{m}(\sin \phi) P_{n}^{m}\left(\sin \phi_{d}\right) \cos m\left(\lambda-\lambda_{d}\right) \tag{4.14}
\end{equation*}
$$

where

$$
\begin{gather*}
W_{n m}=\frac{2(n-m)!}{(n+m)!}, \quad(m=1,2, \cdots, n)  \tag{4.15}\\
W_{n 0}=1 \tag{4.16}
\end{gather*}
$$

Linear independence of the spherical harmonic functions,

$$
r^{n} P_{n}^{m}(\sin \phi)\left\{\begin{array}{l}
\cos  \tag{4.17}\\
\sin
\end{array}\right\} m \lambda
$$

implies that the corresponding coefficients in (4.14) and (A.22) are equal. Thus,

$$
\begin{align*}
& \left(\frac{1}{r_{d}}\right)^{n+1} W_{n m} P_{n}^{m}\left(\sin \phi_{d}\right)\left\{\begin{array}{l}
\cos \\
\sin
\end{array}\right\} m \lambda_{d} \\
& \quad=\left(\frac{1}{c_{d}}\right)^{n+1} \sum_{j} A_{n m j d}\left\{\begin{array}{l}
\cos \\
\sin
\end{array}\right\}\left[-\omega_{j} t-\beta_{j}-(n-m) \frac{\pi}{2}\right] \tag{4.18}
\end{align*}
$$

Equation (4.18) is used to write (4.9) and (4.10) in the form

$$
\begin{align*}
& L_{1}=-3 \frac{\mathrm{Gm}_{d}}{c_{d}^{3}} m_{E} a_{E}^{2} J_{2} \sum_{j} A_{21 j} \cos \left(\omega_{j} t+\beta_{j}\right)  \tag{4.19}\\
& L_{2}=3 \frac{G m_{d}}{c_{d}^{3}} m_{E} a_{E}^{2} J_{2} \sum_{j} A_{21 ;} \sin \left(\omega_{j} t+\beta_{j}\right) \tag{4.20}
\end{align*}
$$

The complex form (2.18) of the torque is

$$
\begin{equation*}
L=\sum_{j} A_{j} e^{-i\left(\omega_{j} t+\beta_{j}\right)} \tag{4.21}
\end{equation*}
$$

where

$$
\begin{equation*}
A_{j}=-3 \frac{G m_{d}}{c_{d}^{3}} m_{E} a_{E}^{2} J_{2} A_{21_{j}} \tag{4.22}
\end{equation*}
$$

In evaluating the torques, the variations in $J_{2}$ due to rotational and tidal deformation are neglected. From (5.16) with $\mathrm{C}_{33}=0$,

$$
\begin{equation*}
\mathrm{J}_{2}=\frac{\mathrm{C}-\mathrm{A}}{\mathrm{~m}_{\mathrm{E}} \mathrm{a}_{\mathrm{E}}^{2}} \tag{4.23}
\end{equation*}
$$

and (4.22) is written as

$$
\begin{equation*}
A_{j}=-3 \frac{G m_{d}}{c_{d}^{3}}(C-A) A_{21 j} \tag{4.24}
\end{equation*}
$$

## 5. TIME VARIATIONS IN THE GEOPOTENTIAL AND IN THE INERTIA TENSOR DUE TO ROTATIONAL AND TIDAL DEFORMATION OF THE EARTH

The Earth deforms as a result of tidal forces and the centrifugal force that arises from its spin about a shifting axis of rotation. Such deformations enter into the polar motion excitation function $\psi$ of Equation (2.16) and into the right hand side of (2.21) by causing variations in the inertia tensor perturbations $c_{13}$, $c_{23}$ and $c_{33}$. In order to write formulas for the inertia tensor perturbations, the parts of the Earth's gravitational potential due to rotational and tidal deformation are first found as functions of time. The inertia tensor perturbations are then related to the geopotential coefficients by Equations (D.13) through (D.17) and (D.26).

The Earth's rotation axis moves relative to the terrestrial reference frame $\mathrm{x}, \mathrm{y}, \mathrm{z}$ of Figure (2.1). The resulting centrifugal force exerted on a mass element having the position vector $\vec{r}$ shown in Figure (5.1) is derivable from the rotational disturbing potential

$$
\begin{equation*}
\frac{1}{2}|\vec{\omega} \times \vec{r}|^{2} \tag{5.1}
\end{equation*}
$$

The components of $\vec{\omega}$ are, from (2.6) through (2.8),

$$
\begin{equation*}
m_{1} \Omega, m_{2} \Omega,\left(1+m_{3}\right) \Omega \tag{5.2}
\end{equation*}
$$

in which $\Omega$ is constant and the $m_{j}$ are of order $10^{-6}$. Neglecting second order terms in the $\mathrm{m}_{\mathrm{j}}$, the rotational disturbing potential (5.1) is written as

$$
\begin{align*}
& \frac{1}{3} r^{2} \Omega^{2}\left[1-P_{2}(\sin \phi)\right] \\
& +\frac{2}{3} r^{2} \Omega^{2} m_{3}\left[1-P_{2}(\sin \phi)\right] \\
& \quad-\frac{1}{3} r^{2} \Omega^{2}\left(m_{1} \cos \lambda+m_{2} \sin \lambda\right) P_{2}^{1}(\sin \phi) \tag{5.3}
\end{align*}
$$



Figure 5.1. Generation of the Rotational Disturbing Potential

The Earth responds by deforming so as to change its external potential by the amount

$$
\begin{align*}
V_{R D}= & -\frac{1}{3} k_{s} a_{E}^{2} \Omega^{2}\left(\frac{a_{E}}{r}\right)^{3} P_{2}(\sin \phi) \\
& -\frac{2}{3} k a_{E}^{2} \Omega^{2} m_{3}\left(\frac{a_{E}}{r}\right)^{3} P_{2}(\sin \phi) \\
& -\frac{1}{3} k a_{E}^{2} \Omega^{2}\left(\frac{a_{E}}{r}\right)^{3}\left(m_{1} \cos \lambda+m_{2} \sin \lambda\right) P_{2}^{1}(\sin \phi) \tag{5.4}
\end{align*}
$$

The "secular" or constant part of the response takes place at a vanishingly small frequency and is therefore written in terms of the secular Love number $\mathrm{k}_{\mathrm{s}}$ [Munk and Macdonald, 1960, p.25]. The second and third terms on the right hand side of (5.4) represent response occuring at frequencies associated with the polar
motion and are given in terms of the "tidal-effective" Love number K [ Munk and Macdonald, 1960, p. 27].

The tide generating potential evaluated at $\mathbf{r}=\mathrm{a}_{\mathbf{E}}$ is, from (A.22),

$$
\begin{align*}
\mathrm{U}\left(\mathrm{a}_{\mathrm{E}}, \phi, \lambda\right)= & \frac{\mathrm{Gm}_{d}}{\mathrm{c}_{\mathrm{d}}} \sum_{n=2}^{\infty} \sum_{m=0}^{n}\left(\frac{\mathrm{a}_{\mathrm{E}}}{\mathrm{c}_{\mathrm{d}}}\right)^{\mathrm{n}} \mathrm{P}_{\mathrm{n}}^{m}(\sin \phi) . \\
& \cdot \sum_{j} A_{n m j} \cos \left[\omega_{j} t+\beta_{j}+m \lambda+(n-m) \frac{\pi}{2}\right] \tag{5.5}
\end{align*}
$$

The second degree zonal part of (5.5) contains one lunar and one solar term for which

$$
\begin{equation*}
\omega_{\mathrm{j}_{\text {sec. }}} \mathrm{t}+\beta_{\mathrm{j}_{\text {sec. }}}=0 \tag{5.6}
\end{equation*}
$$

The response to the secular part of $U$ is given in terms of $k_{s}$ rather than $k$. The external potential arising from response to the tide generating potential is

$$
\begin{align*}
V_{T D}= & -\frac{G m_{d}}{c_{d}} k_{s}\left(\frac{a_{E}}{c_{d}}\right)^{2}\left(\frac{a_{E}}{r}\right)^{3} P_{2}(\sin \phi) A_{20 j_{s e c}} . \\
& +\frac{G m_{d}}{c_{d}} \sum_{n=2}^{\infty} \sum_{m=0}^{n} k_{n}\left(\frac{a_{E}}{r}\right)^{n+1}\left(\frac{a_{E}}{c_{d}}\right)^{n} P_{n}^{m}(\sin \phi) \cdot \\
& \cdot \sum_{j \neq j_{s e c} .} A_{n m j} \cos \left[\omega_{j} t+\beta_{j}+m \lambda+(n-m) \frac{\pi}{2}\right] \tag{5.7}
\end{align*}
$$

The external gravitational potential of the Earth is expanded as

$$
\begin{align*}
V= & \frac{G m_{E}}{r}\left\{1-\sum_{n=2}^{\infty} J_{n}\left(\frac{a_{E}}{r}\right)^{n} P_{n}(\sin \phi)\right. \\
& \left.+\sum_{n=2}^{\infty} \sum_{m=1}^{n}\left(\frac{a_{E}}{r}\right)^{n} P_{n}^{m}(\sin \phi)\left(C_{n m} \cos m \lambda+S_{n m} \sin m \lambda\right)\right\} \tag{5.8}
\end{align*}
$$

The second degree zonal part of $V$ is set equal to the combined zonal terms in (5.4) and (5.7)

$$
\begin{align*}
& -\frac{G m_{E}}{r} J_{2}\left(\frac{a_{E}}{r}\right)^{2} P_{2}(\sin \phi)= \\
& -\frac{1}{3} k_{s} a_{E}^{2} \Omega^{2}\left(\frac{a_{E}}{r}\right)^{3} P_{2}(\sin \phi) \\
& -\frac{2}{3} k a_{E}^{2} \Omega^{2} m_{3}\left(\frac{a_{E}}{r}\right)^{3} P_{2}(\sin \phi) \\
& -\frac{G m_{d}}{c_{d}} k_{s}\left(\frac{a_{E}}{c_{d}}\right)^{2}\left(\frac{a_{E}}{r}\right)^{3} P_{2}(\sin \phi) A_{20 j_{s e c}} . \\
& -\frac{G m_{d}}{c_{d}} k\left(\frac{a_{E}}{c_{d}}\right)^{2}\left(\frac{a_{E}}{r}\right)^{3} P_{2}(\sin \phi) . \\
& \cdot \sum_{j \neq j_{s e c}} \tag{5.9}
\end{align*}
$$

The inertia tensor perturbations corresponding to $J_{2}$ are found from Equations (D.13) through (D.17) and Equation (D.26), which take the form

$$
\begin{align*}
& \frac{C-A}{m_{E} a_{E}^{2}}+\frac{2 c_{33}-c_{11}-c_{22}}{2 m_{E} a_{E}^{2}}=J_{2}  \tag{5.10}\\
& c_{22}-c_{11}=0  \tag{5.11}\\
& c_{12}=0  \tag{5.12}\\
& c_{13}=0  \tag{5.13}\\
& c_{23}=0  \tag{5.14}\\
& c_{11}+c_{22}+c_{33}=0 \tag{5.15}
\end{align*}
$$

Equations (5.15) and (5.10) are combined to give

$$
\begin{equation*}
\frac{C-A}{m_{E} a_{E}^{2}}+\frac{3 c_{33}}{2 m_{E} a_{E}^{2}}=J_{2} \tag{5.16}
\end{equation*}
$$

The first term on the right hand side of (5.9) is identified with the term (C-A) $/ \mathrm{m}_{\mathrm{E}} \mathrm{a}_{\mathrm{E}}^{2}$ in $J_{2}$. This effectively defines the secular Love number in terms of $\mathrm{C}-\mathrm{A}$ as

$$
\begin{equation*}
\mathrm{k}_{\mathrm{s}}=\frac{3 \mathrm{G}(\mathrm{C}-\mathrm{A})}{\mathrm{a}_{\mathrm{E}}^{5} \Omega^{2}} \tag{5.17}
\end{equation*}
$$

The remaining terms on the right hand side of (5.9) are associated with the inertia perturbation $c_{33}$. The part of $c_{33}$ due to rotational deformation is

$$
\begin{equation*}
c_{33_{\mathrm{RD}}}=\frac{4}{9} \frac{\mathrm{k} \Omega^{2} \mathrm{~m}_{3} \mathrm{a}_{\mathrm{E}}^{5}}{\mathrm{G}} \tag{5.18}
\end{equation*}
$$

Equation (5.18) is written in terms of the secular Love number as

$$
\begin{equation*}
\mathrm{c}_{33_{\mathrm{RD}}}=\frac{4}{3} \frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}(\mathrm{C}-\mathrm{A}) \mathrm{m}_{3} \tag{5.19}
\end{equation*}
$$

The tidal contribution to $c_{33}$ is

$$
\begin{align*}
\mathrm{c}_{33 \mathrm{TD}}= & \frac{2}{3} \mathrm{k}_{\mathrm{s}} \frac{\mathrm{~m}_{\mathrm{d}} \mathrm{a}_{\mathrm{E}}^{5}}{\mathrm{c}_{\mathrm{d}}^{3}} \mathrm{~A}_{20 \mathrm{j}_{\text {sec }}} . \\
& +\frac{2}{3} \mathrm{k} \frac{\mathrm{~m}_{\mathrm{d}} \mathrm{a}_{\mathrm{E}}^{5}}{\mathrm{c}_{\mathrm{d}}^{3}} \sum_{\mathrm{j} \neq \mathrm{j}_{\mathrm{sec}}} A_{20 \mathrm{j}} \cos \left(\omega_{\mathrm{j}} \mathrm{t}+\beta_{\mathrm{j}}\right) \tag{5.20}
\end{align*}
$$

Equations (5.11) through (5.15) give the remaining perturbations in the inertia tensor as

$$
\begin{equation*}
c_{11_{\mathrm{RD}}}=\mathrm{c}_{22_{\mathrm{RD}}}=-\frac{1}{3} \mathrm{c}_{33_{\mathrm{RD}}} \tag{5.21}
\end{equation*}
$$

$$
\begin{equation*}
c_{11}=c_{22}=-\frac{1}{3} c_{33_{\mathrm{TD}}} \tag{5.22}
\end{equation*}
$$

The second degree tesseral part of V is set equal to a constant term characterized by the fixed coefficients $\mathrm{C}_{2 \mathrm{i}_{c}}$ and $\mathrm{S}_{21_{c}}$ plus time varying terms from (5.4) and (5.7) due to rotational and tidal deformation.

$$
\begin{align*}
& \frac{\mathrm{Gm}_{E}}{\mathrm{r}}\left(\frac{\mathrm{a}_{\mathrm{E}}}{\mathrm{r}}\right)^{2} \mathrm{P}_{2}^{1}(\sin \phi)\left\{\begin{array}{l}
\mathrm{C}_{21} \\
\cos \lambda \\
\mathrm{~S}_{21} \\
\sin \lambda
\end{array}\right\} \\
& =\frac{G m_{E}}{r}\left(\frac{a_{E}}{r}\right)^{2} P_{2}^{1}(\sin \phi)\left\{\begin{array}{ll}
\mathrm{C}_{21}{ }_{c} & \cos \lambda \\
S_{21} c & \sin \lambda
\end{array}\right\} \\
& -\frac{1}{3} k a_{E}^{2} \Omega^{2}\left(\frac{a_{E}}{\mathrm{r}}\right)^{3} P_{2}^{1}(\sin \phi)\left\{\begin{array}{l}
m_{1} \cos \lambda \\
m_{2} \sin \lambda
\end{array}\right\} \\
& -\frac{G m_{d}}{c_{d}} k\left(\frac{a_{E}}{r}\right)^{3}\left(\frac{a_{E}}{c_{d}}\right)^{2} P_{2}^{l}(\sin \phi) . \\
& \cdot \sum_{j} A_{21 j}\left\{\begin{array}{l}
\sin \left(\omega_{j} t+\beta_{j}\right) \cos \lambda \\
\cos \left(\omega_{j} t+\beta_{j}\right) \sin \lambda
\end{array}\right\} \tag{5.23}
\end{align*}
$$

In the case of a second degree tesseral potential, the inertia tensor relations are

$$
\begin{align*}
& C_{21}=-\frac{c_{13}}{m_{E} a_{E}^{2}}  \tag{5.24}\\
& S_{21}=-\frac{c_{23}}{m_{E} a_{E}^{2}} \tag{5.25}
\end{align*}
$$

The inertia tensor perturbations due to rotational deformation are

$$
\begin{equation*}
\mathrm{c}_{13_{\mathrm{RD}}}=\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}(\mathrm{C}-\mathrm{A}) \mathrm{m}_{1} \tag{5.26}
\end{equation*}
$$

$$
\begin{equation*}
\mathrm{c}_{23_{\mathrm{RD}}}=\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}(\mathrm{C}-\mathrm{A}) \mathrm{m}_{2} \tag{5.27}
\end{equation*}
$$

In complex form,

$$
\begin{equation*}
c_{\mathrm{RD}}=\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}(\mathrm{C}-\mathrm{A}) \mathrm{m} \tag{5.28}
\end{equation*}
$$

where

$$
\begin{equation*}
c_{R D}=c_{13_{R D}}+i c_{23_{R D}} \tag{5.29}
\end{equation*}
$$

The tidal perturbations in the inertia tensor are

$$
\begin{align*}
& c_{13}=\frac{k m_{d} a_{E}^{5}}{c_{d}^{3}} \sum_{j} A_{21 j} \sin \left(\omega_{j} t+\beta_{j}\right)  \tag{5.30}\\
& c_{23_{T D}}=\frac{k m_{d} a_{E}^{5}}{c_{d}^{3}} \sum_{j} A_{21 j} \cos \left(\omega_{j} t+\beta_{j}\right) \tag{5.31}
\end{align*}
$$

and these are written in complex form as

$$
\begin{equation*}
c_{T D}=\frac{k m_{d} a_{E}^{5}}{c_{d}^{3}} \sum_{j} A_{21_{j}} e^{-i\left(\omega_{j} t+\beta_{j}-\pi / 2\right)} \tag{5.32}
\end{equation*}
$$

where

$$
\begin{equation*}
\mathrm{c}_{\mathrm{TD}}=\mathrm{c}_{13_{\mathrm{TD}}}+\mathrm{i} \mathrm{c}_{23_{\mathrm{TD}}} \tag{5.33}
\end{equation*}
$$

The constant inertia perturbations are

$$
\begin{align*}
& c_{13_{0}}=-m_{E} a_{E}^{2} c_{21_{c}}  \tag{5.34}\\
& c_{23_{0}}=-m_{E} a_{E}^{2} S_{21_{c}} \tag{5.35}
\end{align*}
$$

In complex form,

$$
\begin{equation*}
c_{0}=-m_{E} a_{E}^{2}\left(C_{21_{c}}+i S_{21_{c}}\right) \tag{5.36}
\end{equation*}
$$

where

$$
\begin{equation*}
c_{0}=c_{13_{0}}+i c_{23_{0}} \tag{5.37}
\end{equation*}
$$

The inertia tensor perturbations $c_{11}, c_{22}$ and $c_{12}$ do not enter explicitly into the excitation function (2.16) and therefore they have no direct effect on the polar motion. The lunisolar torques, however, depend ultimately upon every term in the geopotential. Before neglecting all but the $J_{2}$ term for the purpose of evaluating the lunisolar torques, it is necessary to know the magnitude of the rotational and tidal contributions to the $C_{22}$ and $S_{22}$ terms in the geopotential.

The second degree sectorial term of $V$ is set equal to a constant part plus a time varying part due to tidal deformation.

$$
\left.\begin{array}{l}
\frac{G m_{E}}{r}\left(\frac{a_{E}}{r}\right)^{2} P_{2}^{2}(\sin \phi)\left\{\begin{array}{ll}
C_{22} & \cos 2 \lambda \\
S_{22} & \sin 2 \lambda
\end{array}\right\} \\
\quad=\frac{G m_{E}}{r}\left(\frac{a_{E}}{r}\right)^{2} P_{2}^{2}(\sin \phi)\left\{\begin{array}{ll}
C_{22} & \cos 2 \lambda \\
S_{22} & \sin 2 \lambda
\end{array}\right\} \\
\\
\quad+\frac{G m_{d}}{c_{d}} k\left(\frac{a_{E}}{r}\right)^{3}\left(\frac{a_{E}}{c_{d}}\right)^{2} P_{2}^{2}(\sin \phi)
\end{array}\right\} \begin{aligned}
& \text {. } \sum_{j} A_{22 j}\left\{\begin{array}{c}
\cos \left(\omega_{j}+\beta_{j}\right) \cos 2 \lambda \\
-\sin \left(\omega_{j}+\beta_{j}\right) \sin 2 \lambda i
\end{array}\right\} \tag{5.38}
\end{aligned}
$$

In the case of a second degree sectorial potential, Equations (D.13) through (D.17) and (D.26) take the form

$$
\begin{align*}
& C_{22}=\frac{c_{22}-c_{11}}{4 m_{E} a_{E}^{2}}  \tag{5.39}\\
& S_{22}=-\frac{c_{12}}{2 m_{E} a_{E}^{2}} \tag{5.40}
\end{align*}
$$

$$
\begin{align*}
2 c_{33}-c_{11}-c_{22} & =0  \tag{5.41}\\
c_{13} & =0  \tag{5.42}\\
c_{23} & =0  \tag{5.43}\\
c_{11}+c_{22}+c_{33} & =0 \tag{5.44}
\end{align*}
$$

Relations (5.41) and (5.44) give

$$
\begin{gather*}
c_{33}=0  \tag{5.45}\\
c_{11}=-c_{22} \tag{5.46}
\end{gather*}
$$

From (5.39),

$$
\begin{equation*}
C_{22}=\frac{c_{22}}{2 m_{E} a_{E}^{2}} \tag{5.47}
\end{equation*}
$$

The tidal perturbations in the inertia tensor are

$$
\begin{gather*}
c_{22}=-c_{11}=\frac{2 k m_{d D} a_{E}^{5}}{c_{d}^{3}} \sum_{j} A_{22 j} \cos \left(\omega_{j} t+\beta_{j}\right)  \tag{5.48}\\
c_{12}=\frac{2 k_{\mathrm{TD}}}{c_{d}^{3} a_{E}^{5}} \sum_{j} A_{22 j} \sin \left(\omega_{j} t+\beta_{j}\right) \tag{5.49}
\end{gather*}
$$

The constant inertia tensor perturbations are

$$
\begin{gather*}
c_{22_{0}}=-c_{11_{0}}=2 m_{E} a_{E}^{2} C_{22_{c}}  \tag{5.50}\\
c_{12_{0}}=-2 m_{E} a_{E}^{2} S_{22_{c}} \tag{5.51}
\end{gather*}
$$

## 6. POLAR MOTION IN A RIGID EARTH WITH LUNISOLAR TORQUES

The polar motion excitation function due to lunisolar torques is, from (2.16),

$$
\begin{equation*}
\psi_{L}=\frac{i L}{(C-A) \Omega^{2}} \tag{6.1}
\end{equation*}
$$

where the complex torque $L$ is given by (4.21).
In a rigid Earth there is no rotational or tidal deformation to affect the inertia tensor, but constant products of inertia still enter into the excitation function (2.16). The excitation due to the constant products of inertia $c_{0}$ of Equation (5.36) is

$$
\begin{equation*}
\psi_{0}=\frac{\mathrm{C}_{0}}{\mathrm{C}-\mathrm{A}} \tag{6.2}
\end{equation*}
$$

The differential equation (2.12) for the polar motion takes the form

$$
\begin{equation*}
\dot{m}=i \sigma_{r}\left[m-\frac{c_{0}}{C-A}-\sum_{j} \frac{i A_{j}}{(C-A) \Omega^{2}} e^{-i\left(\omega_{j} t+\beta_{j}\right)}\right] \tag{6.3}
\end{equation*}
$$

where $A_{j}$ is given by (4.24). The general solution for the position of the rotation axis is

$$
\begin{equation*}
m=m_{0} e^{i \sigma_{r} t}+\frac{c_{0}}{C-A}+\sum_{j} \frac{i A_{j}}{A \Omega\left(\omega_{j}+\sigma_{r}\right)} e^{-i\left(\omega_{j} t+\beta_{j}\right)} \tag{6.4}
\end{equation*}
$$

The complex constant of integration $m_{0}$ may be written in terms of an amplitude $\gamma_{0}$ and a phase $\Gamma_{0}$ as

$$
\begin{equation*}
m_{0}=\gamma_{0} e^{i \Gamma_{0}} \tag{6.5}
\end{equation*}
$$

The axis of figure in a rigid Earth remains fixed relative to the $x, y, z$ coordinate system. From Equation (E.9), the axis of figure is

$$
\begin{equation*}
\psi_{f}=\frac{c_{0}}{C-A} \tag{6.6}
\end{equation*}
$$

The direction cosines of the angular momentum vector are found by substituting the solution (6.4) for $m$ into Equation (2.24).

$$
\begin{align*}
\frac{H}{C \Omega}= & \frac{A}{C} m_{0} e^{i \sigma_{r} t}+\frac{C_{0}}{C-A} \\
& +\sum_{j} \frac{i A_{j}}{C \Omega\left(\omega_{j}+\sigma_{r}\right)} e^{-i\left(\omega_{j} t+\beta_{j}\right)} \tag{6.7}
\end{align*}
$$

The solutions for the rotation axis, the axis of figure and the angular momentum vector are shown in Figure (6.1). The axis of figure is fixed in the $\mathrm{x}, \mathrm{y}$ plane. The constant $\mathrm{m}_{0}$ represents the initial displacement of the Eulerian pole from the axis of figure. The Eulerian pole moves in a counterclockwise circular path about the axis of figure, completing one cycle in a period of $2 \pi / \sigma_{\mathrm{r}} \cong 10$ months. Lunisolar excitation causes the rotation axis to move in a clockwise epicycle about the Eulerian pole position. Only one term of the summation in (6.4) is represented by the circular epicycle shown in Figure (6.1). The frequencies $\omega_{j}$ are grouped around the siderial frequency of $15.04107 \mathrm{de}-$ grees per hour and lie in the range

$$
\begin{equation*}
11.76554 \text { deg. hr. }{ }^{-1} \leq \omega_{\mathrm{j}} \leq 17.69937 \text { deg. } \mathrm{hr}^{-1} \tag{6.8}
\end{equation*}
$$

The different terms alternately reinforce and cancel one another so that the rotation axis follows a path like the one shown in Figure (6.2). The amplitude of the nearly diurnal epicycle reaches about 62 cm when the principal terms are in phase.

Ultimately it is necessary to locate the rotation axis, the axis of figure and the z axis of the "observatory-fixed" $\mathrm{x}, \mathrm{y}, \mathrm{z}$ system relative to inertially directed axes. The position of a particular axis in inertial space is specified by the Euler angles $\theta$ and $\psi$ defined in Figure (3.1). The Euler angles for the angular momentum vector obtained by integrating Poisson's equations [Woolard, 1953, p. 34] are denoted by $\theta_{H}$ and $\psi_{H}$. The Euler angles $\theta_{r}$ and $\psi_{r}$ for the rotation axis are given in complex form by

$$
\begin{equation*}
\theta_{\mathbf{r}}+i \psi_{\mathrm{r}} \sin \theta=\theta_{\mathrm{H}}+i \psi_{\mathbf{H}} \sin \theta+\delta \theta_{\mathbf{r}}+i \delta \psi_{\mathrm{r}} \sin \theta \tag{6.9}
\end{equation*}
$$

where, from the kinematical relationship (3.10),

$$
\begin{equation*}
\delta \theta_{r}+i \delta \psi_{r} \sin \theta=-i e^{i \phi}\left(m-\frac{H}{\mathrm{C} \Omega}\right) \tag{6.10}
\end{equation*}
$$



Figure 6.1. Polar Motion in a Rigid Earth


Figure 6.2. Diurnal Polar Motion

The Euler angle perturbations for the axis of figure and the terrestrial $z$ axis are given respectively by

$$
\begin{gather*}
\delta \theta_{f}+i \delta \psi_{f} \sin \theta=-i e^{i \phi}\left(\psi_{f}-\frac{\mathrm{H}}{\mathrm{C} \Omega}\right)  \tag{6.11}\\
\delta \theta_{z}+i \delta \psi_{z} \sin \theta=i \mathrm{e}^{\mathrm{i} \phi} \frac{\mathrm{H}}{\mathrm{C} \Omega} \tag{6.12}
\end{gather*}
$$

Substitution of (6.4), (6.6) and (6.7) into (6.10), (6.11) and (6.12) gives

$$
\begin{align*}
\delta \theta_{r}+i \delta \psi_{r} \sin \theta= & -i\left(\frac{C-A}{C}\right) m_{0} e^{i\left(\sigma_{r} t+\phi\right)} \\
& +\left(\frac{C-A}{C}\right) \sum_{j \cdot} \frac{A_{j}}{A \Omega\left(\omega_{j}+\sigma_{r}\right)} e^{-i\left(\omega_{j} t+\beta_{j}-\phi\right)}  \tag{6.13}\\
\delta \theta_{f}+i \delta \psi_{f} \sin \theta= & i \frac{A}{C} m_{0} e^{i\left(\sigma_{r} t+\phi\right)} \\
& -\sum_{j} \frac{A_{j}}{C \Omega\left(\omega_{j}+\sigma_{r}\right)} e^{-i\left(\omega_{j} t+\beta_{j}-\phi\right)}  \tag{6.14}\\
\delta \theta_{z}+i \delta \psi_{z} \sin \theta= & i \frac{A}{C} m_{0} e^{i\left(\sigma_{r} t+\phi\right)}+i\left(\frac{c_{0}}{C-A}\right) e^{i \phi} \\
& -\sum_{j} \frac{A_{j}}{C \Omega\left(\omega_{j}+\sigma_{r}\right)} e^{-i\left(\omega_{j} t+\beta_{j}-\phi\right)} \tag{6.15}
\end{align*}
$$

The arguments $\omega_{j+} \mathrm{t}+\beta_{\mathrm{j}+}$ and $\omega_{\mathrm{j}-} \mathrm{t}+\beta_{\mathrm{j}-}$ - correspondingly to certain pairs of distinct indices $j+$ and $j$ - are symmetric with respect to the Greenwich mean sidereal hour angle $\phi_{M}$ in the sense that

$$
\begin{align*}
& \omega_{j+} t+\beta_{j+}=\left(\phi_{M}+\pi\right)+a_{j}  \tag{6.16}\\
& \omega_{j-} t+\beta_{j-}=\left(\phi_{M}+\pi\right)-a_{j} \tag{6.17}
\end{align*}
$$

where $a_{j}$ denotes a linear combination of Doodson's standard variables defined by Equation (B.13). The Euler angle perturbations (6.13) through (6.15) contain sums of the general form

$$
\begin{equation*}
\sum_{j} \tilde{A}_{j} e^{-i\left(\omega_{j} t+\beta_{j}-\phi\right)} \tag{6.18}
\end{equation*}
$$

considered in Appendix F. The terms in (F.5) that contain $\sin \left(\phi_{M}-\phi\right)$ involve products of the small difference ( $\phi_{M}-\phi$ ) given by Equation (C.5) with factors the size of the Euler angle perturbations themselves. The $\sin \left(\phi_{M}-\phi\right)$ terms are neglected and the resulting expressions are

$$
\begin{align*}
& \delta \theta_{r}=\gamma_{0}\left(\frac{C-A}{C}\right) \sin \left(\sigma_{r} t+\phi+\Gamma_{0}\right) . \\
& -\sum_{\mathbf{j}}\left(\frac{\mathbf{C}-\mathbf{A}}{\mathbf{C}}\right) \frac{1}{\mathbf{A} \Omega}\left(\frac{\mathbf{A}_{\mathbf{j}+}}{\omega_{\mathbf{j}+}+\sigma_{\mathbf{r}}}+\frac{\mathbf{A}_{\mathbf{j}-}}{\omega_{\mathrm{j}-}+\sigma_{\mathrm{r}}}\right) \cos \alpha_{\mathrm{j}}  \tag{6.19}\\
& \delta \psi_{\mathrm{r}} \sin \theta=-\gamma_{0}\left(\frac{\mathrm{C}-\mathrm{A}}{\mathrm{C}}\right) \cos \left(\sigma_{\mathrm{r}}^{\mathrm{t}}+\phi+\Gamma_{0}\right) \\
& +\sum_{j}\left(\frac{C-A}{C}\right) \frac{1}{A \Omega}\left(\frac{A_{j+}}{\omega_{j+}+\sigma_{r}}-\frac{A_{j-}}{\omega_{j-}+\sigma_{r}}\right) \sin a_{j}  \tag{6.20}\\
& \delta \theta_{f}=-\gamma_{0} \frac{A}{C} \sin \left(\sigma_{f} t+\phi+\Gamma_{0}\right) \\
& +\sum_{j} \frac{1}{C \Omega}\left(\frac{\mathbf{A}_{j+}}{\omega_{j+}+\sigma_{r}}+\frac{\mathbf{A}_{j-}}{\omega_{j-}+\sigma_{r}}\right) \cos \alpha_{j} \tag{6.21}
\end{align*}
$$

$\delta \psi_{\mathrm{f}} \sin \theta=\gamma_{0} \frac{\mathrm{~A}}{\mathrm{C}} \cos \left(\sigma_{\mathrm{r}} \mathrm{t}+\phi+\Gamma_{0}\right)$

$$
\begin{equation*}
-\sum_{j} \frac{1}{\mathrm{C} \Omega}\left(\frac{\mathrm{~A}_{\mathrm{j}+}}{\omega_{\mathrm{j}+}+\sigma_{\mathrm{r}}}-\frac{\mathrm{A}_{\mathrm{j}-}}{\omega_{\mathrm{j}-}+\sigma_{\mathrm{r}}}\right) \sin \alpha_{\mathrm{j}} \tag{6.2.2}
\end{equation*}
$$

$$
\delta \theta_{z}=-\gamma_{0} \frac{\mathrm{~A}}{\mathrm{C}} \sin \left(\sigma_{\mathrm{r}} \mathrm{t}+\phi+\Gamma_{0}\right)
$$

$$
-\left(\frac{\mathrm{C}_{23_{0}}}{\mathrm{C}-\mathrm{A}}\right) \cos \phi-\left(\frac{\mathrm{c}_{13_{0}}}{\mathrm{C}-\mathrm{A}}\right) \sin \phi
$$

$$
\begin{equation*}
+\sum_{j} \frac{1}{\mathbf{C} \Omega}\left(\frac{\mathbf{A}_{j+}}{\omega_{j+}+\sigma_{s}}+\frac{\mathbf{A}_{j-}}{\omega_{j-}+\sigma_{r}}\right) \cos a_{j} \tag{6.23}
\end{equation*}
$$

$\delta \psi_{\mathrm{z}} \sin \theta=\gamma_{0} \frac{\mathrm{~A}}{\mathrm{C}} \cos \left(\sigma_{\mathrm{r}} \mathrm{t}+\phi+\Gamma_{0}\right)$

$$
\begin{align*}
& +\left(\frac{c_{13_{0}}}{C-A}\right) \cos \phi-\left(\frac{c_{23_{0}}}{C-A}\right) \sin \phi \\
& -\sum_{j} \frac{1}{C \Omega}\left(\frac{A_{j+}}{\omega_{j+}+\sigma_{r}}-\frac{A_{j-}}{\omega_{j-}+\sigma_{r}}\right) \sin \alpha_{j} \tag{6.24}
\end{align*}
$$

## 7. POLAR MOTION IN A DEFORMABLE EARTH WITH LUNISOLAR TORQUES AND TIDAL DEFORMATION

The polar motion excitation functions due to lunisolar torques and to the constant part of the inertia tensor perturbations are

$$
\begin{equation*}
\psi_{L}=\frac{i L}{(C-A) \Omega^{2}} \tag{7.1}
\end{equation*}
$$

and

$$
\begin{equation*}
\psi_{0}=\frac{c_{0}}{C-A} \tag{7.2}
\end{equation*}
$$

respectively, where $L$ is given by (4.21) and $c_{0}$ is given by (5.36). In a rigid Earth $\psi_{L}$ and $\psi_{0}$ are the only sources of excitation and their effects are considered in Section 6.

In a non-rigid Earth rotational deformation gives rise to the products of inertia $\mathrm{c}_{\mathrm{RD}}$ given by (5.28) and tidal deformation causes the products of inertia $\mathrm{c}_{\mathrm{TD}}$ given by (5.32). The excitation functions due to rotational and tidal deformation are, from (2.16),

$$
\begin{align*}
& \psi_{\mathrm{RD}}=\frac{\mathrm{c}_{\mathrm{RD}}}{\mathrm{C}-\mathrm{A}}-\frac{\mathrm{i}}{\Omega} \frac{\dot{\mathrm{c}}_{\mathrm{RD}}}{\mathrm{C}-\mathrm{A}}  \tag{7.3}\\
& \psi_{\mathrm{TD}}=\frac{\mathrm{c}_{\mathrm{TD}}}{\mathrm{C}-\mathrm{A}}-\frac{\mathrm{i}}{\Omega} \frac{\dot{\mathrm{c}}_{\mathrm{TD}}}{\mathrm{C}-\mathrm{A}} \tag{7.4}
\end{align*}
$$

The rotation pole $m$ satisfies (2.12), which takes the form

$$
\begin{equation*}
\dot{\mathrm{m}}=\mathrm{i} \sigma_{\mathrm{F}}\left(\mathrm{~m}-\psi_{0}-\psi_{\mathbf{L}}-\psi_{\mathbf{R D}}-\psi_{\mathrm{TD}}\right) \tag{7.5}
\end{equation*}
$$

Substituting for $c_{R D}$ from (5.28) into (7.3) gives

$$
\begin{equation*}
\psi_{\mathrm{RD}}=\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}} \mathrm{~m}-\frac{\mathrm{i}}{\Omega} \frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}} \dot{\mathrm{~m}} \tag{7.6}
\end{equation*}
$$

Since $\psi_{\mathrm{RD}}$ contains terms in $m$ and $\dot{\mathrm{m}}$, it is necessary to modify the form of (7.5). Equation (7.5) becomes

$$
\begin{equation*}
\dot{\mathrm{m}}=\mathrm{i} \sigma_{0} \mathrm{~m}-\frac{\mathrm{i} \sigma_{\mathrm{r}}}{\left(1+\frac{\sigma_{\mathrm{r}}}{\Omega} \frac{\mathbf{k}}{\mathbf{k}_{\mathbf{s}}}\right)}\left(\psi_{0}+\psi_{\mathrm{L}}+\psi_{\mathrm{TD}}\right) \tag{7.7}
\end{equation*}
$$

where

$$
\begin{equation*}
\sigma_{0}=\frac{\sigma_{r}\left(1-\frac{\mathrm{k}}{\mathrm{k}_{s}}\right)}{\left(1+\frac{\sigma_{\mathrm{r}}}{\Omega} \frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right)} \tag{7.8}
\end{equation*}
$$

Separate solutions of (7.7) are first obtained for each of the excitations $\psi_{0}$, $\psi_{L}$, and $\psi_{T D}$. These solutions are then combined with the free motion to form the complete solution. The solution due to $\psi_{0}$ is

$$
\begin{equation*}
\underset{\substack{\text { of inertia }}}{m_{\text {fixed product }}} \frac{\left(\frac{c_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)} \tag{7.9}
\end{equation*}
$$

The polar motion due to lunisolar torques is

$$
\begin{equation*}
\underset{\substack{\text { tunigolar } \\ \text { torques }}}{\mathrm{m}_{\mathrm{j}}} \sum_{\left(1+\frac{\sigma_{\mathrm{r}}}{\Omega} \frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right) \mathrm{A} \Omega\left(\omega_{j}+\sigma_{0}\right)} e^{-i\left(\omega_{j} t+\beta_{j}\right)} \tag{7.10}
\end{equation*}
$$

The coefficients $\mathrm{A}_{\mathrm{j}}$ are given by (4.24),

$$
\begin{equation*}
A_{j}=-3 \frac{G m_{d}}{c_{d}^{3}}(C-A) A_{21 j} \tag{7.11}
\end{equation*}
$$

and the $\mathrm{A}_{21 \mathrm{j}}$ are related to Doodson's coefficients by Equation (A.23). The solution due to tidal deformation is

$$
\begin{equation*}
\mathrm{m}_{\substack{\text { tidal } \\ \text { deformation }}}=\sum_{j} \frac{i \sigma_{\mathrm{r}} \mathrm{~A}_{\mathrm{Tj}}}{\left(1+\frac{\sigma_{\mathrm{r}}}{\Omega} \frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right)\left(\omega_{\mathrm{j}}+\sigma_{0}\right)} e^{-\mathrm{i}\left(\omega_{\mathrm{j}} \mathrm{t}+\beta_{\mathrm{j}}\right)} \tag{7.12}
\end{equation*}
$$

where

$$
\begin{equation*}
A_{T j}=k a_{E}^{5}\left(\frac{m_{d}}{c_{d}^{3}}\right)\left(\frac{1}{C-A}\right)\left(\frac{n_{j}}{\sqrt{2}}\right) A_{21_{j}} \tag{7.13}
\end{equation*}
$$

The frequency $n_{j}$ is defined as

$$
\begin{equation*}
\mathbf{n}_{\mathbf{j}}=\Omega-\omega_{\mathbf{j}} \tag{7.14}
\end{equation*}
$$

and is related to Doodson's arguments by Equation (B.14). The free motion is

$$
\begin{equation*}
m_{\text {free }}=m_{0} e^{i \sigma_{0} t} \tag{7.15}
\end{equation*}
$$

where $m_{0}$ is a complex constant of integration which is written in terms of its amplitude $\gamma_{0}$ and phase $\Gamma_{0}$ as

$$
\begin{equation*}
m_{0}=\gamma_{0} \mathrm{e}^{\mathrm{i} \Gamma_{0}} \tag{7.16}
\end{equation*}
$$

It is useful to make some observations about the separate solutions before combining them to form the complete solution. If the Love number k is zero, Equations (7.9), (7.10) and (7.15) reduce to the corresponding terms in (6.4) for a rigid Earth, and the polar motion (7.12) due to tidal deformation vanishes. Each coefficient $A_{T_{j}}$ is proportional to the angular rate $n_{j}$ by which the frequency of the jth tidal component differs from the Earth's rotation rate. When $\omega_{\mathrm{j}}=\Omega$, the jth tidal component stands still in space and has no effect on the polar motion. The period of the free motion is lengthened by rotational deformation from 10 months to $2 \pi / \sigma_{0} \cong 14$ months.

Tidal deformation produces polar motion having components with the same frequencies as those for the motion excited by lunisolar torques. The ratio of the jth coefficient in (7.12) to the corresponding coefficient in (7.10) is

$$
\begin{equation*}
\frac{m_{j, ~ t i d a l ~ d e ̀ f o r m a t i o n ~}}{m_{j, ~ l u n i s o l a r ~ t o r q u e s ~}}=-\frac{n_{j}}{\Omega} \frac{k}{k_{s}} \tag{7.17}
\end{equation*}
$$

Equation (7.17) shows that the motion due to tidal deformation is smaller than that due to lunisolar torques and directed in the opposite sense.

The solutions (7.9), (7.10), (7.12) and (7.15) are combined to give

$$
\begin{equation*}
m=m_{0} e^{i \sigma_{0} t}+\frac{\left(\frac{c_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)}+\sum_{j} \frac{i}{A \Omega^{2}} A_{j R A} e^{-i\left(\omega_{j} t+\beta_{j}\right)} \tag{7.18}
\end{equation*}
$$

where

$$
\begin{equation*}
\mathbf{A}_{\mathrm{jRA}}=\left[\frac{\mathbf{A}_{j}}{\frac{\left(\frac{\omega_{j}}{\Omega}\right)}{\left(1-\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)^{2}}+\frac{\sigma_{r}}{\Omega}}\right] \tag{7.19}
\end{equation*}
$$

The axis of figure is from Equation (E.9),

$$
\begin{equation*}
\psi_{\mathrm{f}}=\frac{\mathrm{c}}{\mathrm{C}-\mathrm{A}} \tag{7.20}
\end{equation*}
$$

where $c$ is the total of all perturbations in the inertia tensor.

$$
\begin{equation*}
\mathrm{c}=\mathrm{c}_{0}+\mathrm{c}_{\mathrm{RD}}+\mathrm{c}_{\mathrm{TD}} \tag{7.21}
\end{equation*}
$$

The motion of the axis of figure due solely to rotational deformation is

$$
\begin{equation*}
\psi_{f_{R D}}=\frac{c_{R D}}{\mathrm{C}-\mathrm{A}}=\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}} \mathrm{~m} \tag{7.22}
\end{equation*}
$$

Substituting for $m$ from (7.18) into (7.22) gives

$$
\begin{equation*}
\frac{c_{0}}{C-A}+\psi_{f_{R D}}=\frac{\left(\frac{c_{0}}{C-A}\right)}{\left(1-\frac{k}{k}\right)}+\frac{k}{k_{s}} m_{0} e^{i \sigma_{0} t}+\sum_{j} \frac{k}{k_{s}} \frac{i}{A \Omega^{2}} A_{j R A} e^{-i\left(\omega_{j} t+\beta_{j}\right)} \tag{7.23}
\end{equation*}
$$

The motion of the axis of figure due to tidal deformation alone is

$$
\begin{equation*}
\psi_{\mathrm{f}_{\mathrm{TD}}}=\frac{\mathrm{C}_{\mathrm{TD}}}{\mathrm{C}-\mathrm{A}} \tag{7.24}
\end{equation*}
$$

Substituting for $\mathrm{c}_{\mathrm{TD}}$ from (5.32) into (7.23) and making use of (7.11) and (5.17) gives

$$
\begin{equation*}
\psi_{f}=\sum_{j}-i \frac{k}{k_{g}} \frac{A_{j}}{\Omega^{2}(C-A)} e^{-i\left(\omega_{j} t+\beta_{j}\right)} \tag{7.25}
\end{equation*}
$$

The appearance of C-A in the denominator of (7.25) causes each term in $\psi_{\mathrm{f}}$ to be much larger than the corresponding term in $\psi_{f_{R D}}$. The jth terms of (7.25) and (7.23) have the ratio

$$
\begin{equation*}
\frac{\psi_{f_{j}, \text { tidal deformation }}}{\psi_{f_{j, ~ r o t a t i o n a l ~ d e f o r m a t i o n ~}}}=-\left(\frac{A}{C-A}\right)\left[\frac{\left(\frac{\omega_{j}}{\Omega}\right)}{\left(1+\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)}+\frac{\sigma_{r}}{\Omega}\right] \tag{7.26}
\end{equation*}
$$

The tidal effect is therefore oppositely directed from the rotational effect and about 300 times larger. In terms of displacement at the Earth's surface, the amplitudes are of order

$$
\begin{equation*}
\left|\psi_{\mathrm{f}_{\mathrm{j}, ~ t i d a l ~ d e f o r m a t i o n ~}}\right| \sim 60 \mathrm{~m} \tag{7.27}
\end{equation*}
$$

and

$$
\begin{equation*}
\left|\psi_{\mathrm{f}_{\mathrm{j}}, \text { rotational deformation }}\right| \sim 20 \mathrm{~cm} \tag{7.28}
\end{equation*}
$$

The combined motion of the axis of figure is, from (7.23) and (7.25),

$$
\begin{equation*}
\psi_{f}=\frac{\left(\frac{c_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)}+\frac{k}{k_{s}} m_{0} e^{i \sigma_{0} t}-\sum_{j} \frac{i}{A \Omega \sigma_{r}} \frac{k}{k_{s}} A_{j A F} e^{-i\left(\omega_{j} t+\beta_{j}\right)} \tag{7.29}
\end{equation*}
$$

where

$$
\begin{equation*}
A_{j_{A_{A F}}}=\left[\frac{\left(\frac{\left(\frac{\omega_{j}}{\Omega}\right)}{\left(1-\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)}\right.}{\left.\frac{\left(\frac{\omega_{j}}{\Omega}\right)}{\left(1-\frac{n_{j}}{\Omega} k\right.} k_{s}\right)}+\frac{\sigma_{r}}{\Omega}\right] \tag{7.30}
\end{equation*}
$$

When $\mathrm{k}=0$, the solution (7.30) reduces to that given by (6.6) for a rigid Earth.
The direction cosines (2.28) of the angular momentum vector are found by substituting the solution (7.18) for m into Equation (2.24).

$$
\begin{align*}
\frac{H}{C \Omega} & =\frac{\left(\frac{C_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)}+\left[\frac{A}{C}+\left(\frac{C-A}{C}\right) \frac{k}{k_{s}}\right] m_{0} e^{i \sigma_{0} t} \\
& +\sum_{j} \frac{i}{C \Omega^{2}}\left(1-\frac{k}{k_{s}}\right) A_{j A M} e^{-i\left(\omega_{j} t+\beta_{j}\right)} \tag{7.31}
\end{align*}
$$

where

$$
\begin{equation*}
A_{j A M}=\left[\frac{\frac{A_{j}}{\left(1-\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)}}{\frac{\left(\frac{\omega_{j}}{\Omega}\right)}{\left(1-\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)}+\frac{\sigma_{r}}{\Omega}}\right] \tag{7.32}
\end{equation*}
$$

The factor $\left(1-\mathrm{k} / \mathrm{k}_{\mathrm{s}}\right) \cong 2 / 3$ in the summation term of (7.31) causes the diurnal terms for the angular momentum vector to be only $2 / 3$ the size of the corresponding terms in the motion of the rotation axis. The result is that the angular momentum vector and the rotation axis are separated by as much as $1 / 3 \times 62 \mathrm{~cm}$, or 21 cm , instead of remaining within 1 or 2 cm of each other as they do in the case of a rigid Earth.

The solutions for $\mathrm{m}, \psi_{\mathrm{f}}$, and $\mathrm{H} / \mathrm{C} \Omega$ are shown in Figure (7.1). The fixed pole $\Psi$ is defined by

$$
\begin{equation*}
\Psi=\frac{\left(\frac{c_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)} \tag{7.33}
\end{equation*}
$$

An interesting consequence of Equation (2.24) is

$$
\begin{equation*}
\frac{\mathrm{H}}{\mathrm{C} \Omega}-\psi_{\mathrm{f}}=\frac{\mathrm{A}}{\mathrm{C}}\left(\mathrm{~m}-\psi_{\mathrm{f}}\right) \tag{7.34}
\end{equation*}
$$

so that the angular momentum vector, the axis of rotation and the axis of figure all lie in the same plane, for either a rigid or a deformable Earth.

In the case of a deformable Earth, the Euler angle perturbations defined by Equations (6.10) through (6.12) are


Figure 7.1. Polar Motion in a Deformable Earth

$$
\begin{align*}
\delta \theta_{r}+i \delta \psi_{r} \sin \theta & =-i\left(\frac{C-A}{C}\right)\left(1-\frac{k}{k_{s}}\right) m_{0} e^{i\left(\sigma_{0} t+\phi\right)} \\
& +\sum_{j} \frac{1}{A \Omega^{2}}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j}}{\Omega}\right)\right] . \\
& \cdot A_{j A M} e^{\infty i\left(\omega_{j} t+\beta_{j}-\phi\right)} \tag{7.35}
\end{align*}
$$

$$
\begin{align*}
\delta \theta_{f}+i \delta \psi_{f} \sin \theta= & i \frac{A}{C}\left(1-\frac{k}{k_{s}}\right) m_{0} e^{i\left(\sigma_{0} t+\phi\right)} \\
& -\sum_{j} \frac{1}{A \Omega \sigma_{r}}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j}}{\Omega}\right)\right] \cdot \\
& \cdot A_{j A M} e^{-i\left(\omega_{j}\right.}{ }^{\left.t+\beta_{j}-\phi\right)}  \tag{7.36}\\
\delta \theta_{z}+i \delta \psi_{z} \sin \theta= & i \frac{\left(\frac{c_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)} e^{i \phi}+i\left[\frac{A}{C}+\left(\frac{C-A}{C}\right) \frac{k}{k_{s}}\right] m_{0} e^{i\left(\sigma_{0} t+\phi\right)} \\
& -\sum_{j} \frac{1}{C \Omega^{2}}\left(1-\frac{k}{k_{s}}\right) A_{j_{A M}} e^{-i\left(\omega_{j} t+\beta_{j}-\phi\right)} \tag{7.37}
\end{align*}
$$

The Euler angle perturbations contain sums of the form

$$
\begin{equation*}
\sum_{j} \widetilde{\mathrm{~A}}_{\mathrm{j}} \mathrm{e}^{-\mathrm{i}\left(\omega_{\mathrm{j}} \mathrm{t}+\beta_{\mathrm{j}}-\phi\right)} \tag{7.38}
\end{equation*}
$$

considered in Appendix F. The terms in Equation (F.5) that contain $\sin \left(\phi_{M}-\phi\right)$ involve products of the small difference ( $\phi_{M}-\phi$ ) given by (C.5) with factors the size of the Euler angle perturbations themselves. The $\sin \left(\phi_{M}-\phi\right)$ terms are therefore of second order. When they are neglected the resulting expressions for the Euler angle perturbations are

$$
\begin{align*}
& \delta \theta_{r}=\left(\frac{C-A}{C}\right)\left(1-\frac{k}{k_{s}}\right) \gamma_{0} \sin \left(\sigma_{0} t+\phi+\Gamma_{0}\right) \\
& -\sum_{j} \frac{1}{A \Omega^{2}}\left\{A_{j+}{ }_{A M}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j+}}{\Omega}\right)\right]\right. \\
& \left.+A_{j-A M}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j-}}{\Omega}\right)\right]\right\} \cos \alpha_{j}  \tag{7.39}\\
& \delta \psi_{\mathrm{r}} \sin \theta=-\left(\frac{\mathrm{C}-\mathrm{A}}{\mathrm{C}}\right)\left(1-\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right) \gamma_{0} \cos \left(\sigma_{0} \mathrm{t}+\phi+\Gamma_{0}\right) \\
& +\sum_{j} \frac{1}{A \Omega^{2}}\left\{A_{j+A M}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j t}}{\Omega}\right)\right]\right. \\
& \left.-A_{j-}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j}-}{\Omega}\right)\right]\right\} \sin a_{j}  \tag{7.40}\\
& \delta \theta_{f}=-\frac{A}{C}\left(1-\frac{k}{k_{s}}\right) \gamma_{0} \sin \left(\sigma_{0} t+\phi+\Gamma_{0}\right) \\
& +\sum_{j} \frac{1}{A \Omega \sigma_{r}}\left\{A_{j+}^{A M}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j+}}{\Omega}\right)\right]\right. \\
& \left.+A_{j-}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j-}}{\Omega}\right)\right]\right\} \cos \alpha_{j} \tag{7.41}
\end{align*}
$$

$$
\begin{align*}
& \delta \psi_{\mathrm{f}} \sin \theta=\frac{\mathrm{A}}{\mathrm{C}}\left(1-\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right) \gamma_{0} \cos \left(\sigma_{\sigma} \mathrm{t}+\phi+\Gamma_{\mathrm{o}}\right) \\
& -\sum_{j} \frac{1}{A \Omega \sigma_{r}}\left\{A_{j+}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j+}}{\Omega}\right)\right]\right. \\
& \left.-A_{j-A M}\left[\left(\frac{C-A}{C}\right)+\frac{k}{k_{s}}\left(\frac{A}{C}-\frac{n_{j}-}{\Omega}\right)\right]\right\} \sin a_{j}  \tag{7.42}\\
& \delta \theta_{z}=-\left[\frac{\mathrm{A}}{\mathrm{C}}+\left(\frac{\mathrm{C}-\mathrm{A}}{\mathrm{C}}\right) \frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right] \gamma_{0} \sin \left(\sigma_{0} \mathrm{t}+\phi+\Gamma_{0}\right) \\
& -\left(\frac{1}{1-\frac{k}{k_{s}}}\right)\left[\left(\frac{{ }^{c} 23_{0}}{C-A}\right) \cos \phi+\left(\frac{{ }^{c} 13_{0}}{C-A}\right) \sin \phi\right] \\
& +\sum_{j} \frac{1}{C \Omega^{2}}\left(1-\frac{k}{k_{f}}\right)\left[A_{j+}+A_{j-}{ }_{A N}\right] \cos \alpha_{j}  \tag{7.43}\\
& \delta \psi_{z} \sin \theta=\left[\frac{\mathrm{A}}{\mathrm{C}}+\left(\frac{\mathrm{C}-\mathrm{A}}{\mathrm{C}}\right) \frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right] \gamma_{0} \cos \left(\sigma_{0} \mathrm{t}+\phi+\Gamma_{0}\right) \\
& +\left(\frac{1 .}{1-\frac{k}{k_{s}}}\right)\left[\left(\frac{{ }^{C} 13_{0}}{C-A}\right) \cos \phi-\left(\frac{c_{23_{0}}}{C-A}\right) \sin \phi\right] \\
& -\sum_{j} \frac{1}{C \Omega^{2}}\left(1-\frac{k}{k_{s}}\right)\left[A_{j+}{ }_{A M}-A_{j-}{ }_{A M}\right] \sin \alpha_{j} \tag{7.44}
\end{align*}
$$

The arguments $\alpha_{j}$ are defined by Equation (B.13).

## 8. COMPARISON OF THE PRESENT THEORY WITH WOOLARD'S DEVELOPMENT OF THE EULERIAN AND LUNISOLAR CORRECTION TERMS

Doodson's [1922] expansion of the tide generating potential is based upon Brown's [1905] lunar theory and Newcomb's [1898] theory of the sun. Woolard [1953] also uses the theories of Brown and Newcomb in his development of the lunisolar precession and nutation. Since the solutions obtained in Sections 6 and 7 are derived using Doodson's expansion, they may be expected to agree closely with Woolard's equations for the Eulerian and lunisolar correction terms. Woolard's paper is devoted mainly to a very elegant analytical integration of Poisson's equations in order to obtain the Euler angles $\theta_{\mathrm{H}}$ and $\psi_{\mathrm{H}}$ of the angular momentum vector. No attempt is made here to comment on this part of his procedure. Of primary interest are the corrections $\delta \theta_{z}$ and $\delta \psi_{\mathrm{z}}$ that must be added to $\theta_{\mathrm{H}}$ and $\psi_{\mathrm{H}}$ in order to obtain the position of the observatory-fixed z axis.

Woolard assumes a rigid Earth, so his results must be compared with the solutions for a rigid Earth in Section 6, or equivalently, with those of Section 7 after setting the Love number $k$ equal to zero. In order to make a detailed comparison, Woolard's development of the Euler angle perturbations and the polar motion coordinates is reproduced using the present notation. References to equation numbers in Woolard's paper are of the form (1), (2), etc.

In a rigid Earth, the axis of figure remains fixed. Woolard takes advantage of this at the outset by aligning the $z$ axis of the terrestrial coordinate system with the axis of figure. In Section 6, the axis of figure is given a fixed displacement $c_{0} /(\mathrm{C}-\mathrm{A})$ relative to the z axis, and a corresponding term is carried along in the restatement of Woolard's equations for purposes of comparison.

Woolard begins by differentiating his form (3) of Euler's kinematic equations. In the present notation, the kinematic equations are

$$
\begin{equation*}
\dot{\theta}+\mathrm{i} \dot{\psi} \sin \theta=-\Omega m \mathrm{e}^{\mathrm{i} \phi} \tag{8.1}
\end{equation*}
$$

and their derivative is

$$
\begin{equation*}
\frac{\mathrm{d}}{\mathrm{dt}}(\dot{\theta})+\mathrm{i} \frac{\mathrm{~d}}{\mathrm{dt}}(\dot{\psi} \sin \theta)=-\Omega \ddot{m} \mathrm{e}^{\mathrm{i} \phi}-\Omega \mathrm{mi} \dot{\phi} \mathrm{e}^{\mathrm{i} \phi} \tag{8.2}
\end{equation*}
$$

The time derivative $\dot{\mathrm{m}}$ is replaced using the dynamical equation,

$$
\begin{equation*}
\dot{m}=i \sigma_{r}\left[m-\frac{c_{0}}{C-A}-\frac{i L}{\Omega^{2}(C-A)}\right] \tag{8.3}
\end{equation*}
$$

which is equivalent to Woolard's (6). Substitution of (8.3) into (8.2) gives

$$
\begin{equation*}
\frac{\mathrm{d}}{\mathrm{dt}}(\dot{\theta})+\mathrm{i} \frac{\mathrm{~d}}{\mathrm{dt}}(\dot{\psi} \sin \theta)=-\Omega \mathrm{e}^{\mathrm{i} \phi} \mathbf{i m}\left(\sigma_{\mathrm{r}}+\dot{\phi}\right)-\mathrm{e}^{\mathrm{i} \phi} \frac{\mathrm{~L}}{\mathrm{~A}}+\mathrm{i} \mathrm{e}^{\mathrm{i} \phi} \frac{\mathrm{c}_{0} \Omega^{2}}{\mathrm{~A}} \tag{8.4}
\end{equation*}
$$

The derivative $\ddot{\phi}$ in (8.4) is eliminated using Euler's third kinematic relation,

$$
\begin{equation*}
\omega_{3}=\dot{\phi}+\dot{\psi} \cos \theta \tag{8.5}
\end{equation*}
$$

Since there are no torques about the $z$ axis, the third dynamical equation reduces to

$$
\begin{equation*}
\omega_{3}=\Omega=\text { constant } \tag{8.6}
\end{equation*}
$$

The polar motion $m$ on the right hand side of (8.4) is eliminated using (8.1). The resulting expression is written as

$$
\begin{align*}
\dot{\theta} & +i \dot{\psi} \sin \theta=-\frac{i \mathrm{~L}}{\mathrm{C} \Omega}-\frac{\mathrm{iA}}{\mathrm{C} \Omega}\left[\frac{\mathrm{~d}}{\mathrm{dt}}(\dot{\theta})+\mathrm{i} \frac{\mathrm{~d}}{\mathrm{dt}}(\dot{\psi} \sin \theta)\right] \\
& +\frac{\mathrm{A}}{\mathrm{C} \Omega}(\dot{\theta}+\dot{\psi} \sin \theta) \ddot{\psi} \cos \theta-\frac{\mathrm{c}_{0} \Omega}{\mathrm{C}} \mathrm{e}^{\mathrm{i} \phi} \tag{8.7}
\end{align*}
$$

It is convenient to introduce the complex number $\chi$ defined by

$$
\begin{equation*}
x=\dot{\theta}+i \dot{\psi} \sin \theta \tag{8.8}
\end{equation*}
$$

In terms of $\chi$, Equation (8.7) is

$$
\begin{equation*}
\ddot{x}-i \frac{C \Omega}{A} x+i_{\chi}(\Omega+\dot{\phi})=-\frac{L}{A} e^{i \phi}+i \frac{c_{0} \Omega^{2}}{A} e^{i \phi} \tag{8.9}
\end{equation*}
$$

Equation (8.9) is equivalent to Woolard's (19).
Poisson's equations for the motion of the angular momentum vector are stated as Woolard's Equation (30). In (8.9) the first and third terms on the left are neglected in comparison with the second to obtain Poisson's equations in the form

$$
\begin{equation*}
x_{\mathrm{H}}=-\frac{i L}{\mathrm{C} \Omega} \mathrm{e}^{\mathrm{i} \phi} \tag{8.10}
\end{equation*}
$$

where the subscript $H$ means that the corresponding Euler angles $\theta_{\mathrm{H}}$ and $\psi_{\mathrm{H}}$ refer to the angular momentum vector.

The terms "free" or "Eulerian" motion refer to the motion that would occur in the absence of the lunisolar torque L. Woolard's derivation of the Eulerian terms is discussed on pages 130 and 131 of his paper. The torque $L$ is set equal to zero and it is assumed that

$$
\begin{equation*}
\dot{\phi}=\Omega \tag{8.11}
\end{equation*}
$$

in (8.9). Then

$$
\begin{equation*}
\ddot{x}=\mathrm{i} \frac{\mathrm{C} \Omega}{\mathrm{~A}} x=\mathrm{i} \frac{\mathrm{c}_{0} \Omega^{2}}{\mathrm{~A}} \mathrm{e}^{\mathrm{i} \phi} \tag{8.12}
\end{equation*}
$$

which is the same as Woolard's (52). The solution to (8.12) is

$$
\begin{equation*}
x=x_{0} e^{i(C \Omega / A) t}-\frac{c_{0} \Omega}{C-A} e^{i \phi} \tag{8.13}
\end{equation*}
$$

in which $\chi_{0}$ is a complex constant of integration and the assumption (8.11) is used for an approximate integration of the $c_{0}$ term. Equation (8.13) may be written in terms of the polar motion $m$ by making use of Euler's kinematic equation (8.1) in the form

$$
\begin{equation*}
x=-\Omega m \mathrm{e}^{\mathrm{i} \phi} \tag{8.14}
\end{equation*}
$$

The polar motion solution corresponding to (8.13) is

$$
\begin{equation*}
m=m_{0} e^{i[(c \Omega / A) t-\phi+\phi(0)]}+\frac{c_{0}}{c-A} \tag{8.15}
\end{equation*}
$$

where

$$
\begin{equation*}
m_{0}=-\frac{X_{0}}{\Omega} e^{-i \phi(0)} \tag{8.16}
\end{equation*}
$$

The assumption (8.11) implies that the Euler angle $\phi$ varies linearly with time. Thus

$$
\begin{equation*}
\phi=\Omega \mathbf{t}+\phi(0) \tag{8.17}
\end{equation*}
$$

Equation (8.17) is substituted into Equation (8.14) to obtain

$$
\begin{equation*}
\mathrm{m}=\mathrm{m}_{\mathrm{e}} \mathrm{e}^{\mathrm{i} \sigma_{\mathrm{r}} \mathrm{t}}+\frac{\mathrm{c}_{0}}{\mathrm{C}-\mathrm{A}} \tag{8.18}
\end{equation*}
$$

which agrees exactly with the Eulerian terms in Equation (6.4).
It is of interest that Equation (8.18) is in complete agreement with the Eulerian terms of (6.4) even though the assumption (8.11) is not used at all in deriving (6.4) while it is used twice in deriving (8.18) from (8.9). To see what is happening, Woolard's simplified differential equation (8.12) is written in polar motion form as

$$
\begin{equation*}
\ddot{m}=i\left(\sigma_{r}+\Omega-\ddot{\phi}\right) m-i \frac{c_{0} \Omega}{A} \tag{8.19}
\end{equation*}
$$

Equation (8.19) differs from the Eulerian part of (6.3) by the inclusion of the extraneous term

$$
\begin{equation*}
\mathbf{i}(\Omega-\dot{\phi}) \mathrm{m} \tag{8.20}
\end{equation*}
$$

If the term $i_{X}(\Omega-\dot{\phi})$ in (8.9) is retained at the outset, then the polar motion form is

$$
\begin{equation*}
\ddot{\mathrm{m}}=\mathbf{i} \sigma_{\mathbf{r}} \mathrm{m}-\mathrm{i} \frac{\mathrm{c}_{0} \Omega}{\mathrm{~A}} \tag{8.21}
\end{equation*}
$$

which is identical with the Eulerian part of (6.3). Woolard's initial neglect of the term $\mathrm{i}_{\chi}(\Omega-\dot{\phi})$ in simplifying (8.9) is effectively cancelled by his assumption (8.17) of a linearly varying Euler angle $\phi$.

In order to obtain the Euler angle perturbations, (8.10) is substituted for the torque term in (8.9) and the latter equation is written as

$$
\begin{equation*}
x-x_{H}=-\frac{i \mathrm{~A}}{\mathrm{C} \Omega} \ddot{x}+\frac{\mathrm{A}}{\mathrm{C} \Omega} x(\Omega-\ddot{\phi})-\frac{\mathrm{c}_{0} \Omega}{\mathrm{C}} \mathrm{e}^{\mathrm{i} \phi} \tag{8.22}
\end{equation*}
$$

From the definition (8.8),

$$
\begin{equation*}
x-x_{H}=\delta \ddot{\theta}_{z}+i \delta \dot{\psi}_{z} \sin \theta \tag{8.23}
\end{equation*}
$$

where $\delta \theta_{z}$ and $\delta \psi_{z}$ are the perturbations that must be added to the Euler angles of the angular momentum vector in order to obtain those of the $z$ axis. The integrals of the $\ddot{x}$ and $\chi$ terms on the right hand side of (8.22) are added to the integral of the Eulerian solution (8.13) to give

$$
\begin{align*}
\delta \theta_{z}+i \delta \psi_{z} \sin \theta=K & +i \frac{m_{0} A}{C \Omega} e^{i(C \Omega / A) t} \\
& +i \frac{c_{0}}{C-A} e^{i(\Omega t+\phi(0))} \\
& -i \frac{A}{C \Omega} \chi+\frac{A}{C \Omega} \int x(\Omega-\dot{\phi}) d t \tag{8.24}
\end{align*}
$$

In performing this integration, sine $\theta$ is taken to be constant on the left hand side and the assumption (8.11) is used for the $c_{0}$ term. Equation (8.24) is equivalent to Woolard's Equations (53). Woolard's solution $\chi_{\mathrm{H}}$ of Poisson's equations is then substituted for $x$ in Equation (8.24) in order to obtain his expressions (54) for $\delta \theta_{z}$ and $\delta \psi_{z}$.

Woolard uses his solution for the Euler angle perturbations $\delta \theta_{z}$ and $\delta \psi_{z}$ in order to obtain the diurnal terms in the polar motion. From (2.24) the jth ${ }^{2}$ diurnal term $m_{j}$ in the polar motion is related to the corresponding angular momentum term by

$$
\begin{equation*}
H_{j}=A \Omega m_{j} \tag{8.25}
\end{equation*}
$$

in the case of a rigid Earth with $c_{0}=0$. Substituting (8.25) into (6.12) gives

$$
\begin{equation*}
m_{j}=-i \frac{C}{A} e^{-i \phi}\left(\delta \theta_{z_{j}}+i \delta \psi_{z_{j}} \sin \theta\right) \tag{8.26}
\end{equation*}
$$

which is the complex form of Woolard's (69). Woolard's Equations (70) for the main diurnal polar motion terms are obtained by substituting his solution (54) for the Euler angle perturbations into Equation (8.26).

The arguments in Woolard's diurnal polar motion expressions contain the Euler angle $\phi$. The arguments $\omega_{j} \mathbf{t}+\beta_{\mathrm{j}}$ in the polar motion solution (6.4) contain the Greenwich mean sidereal time $\phi_{M}$ instead of $\phi$ since, from (B.12),

$$
\begin{equation*}
\omega_{j} \mathbf{t}+\beta_{\mathrm{j}}=\left(\phi_{\mathrm{M}}+\pi\right)+\alpha_{\mathrm{j}} \tag{8.27}
\end{equation*}
$$

Woolard's polar motion equations contain $\phi$ in place of $\phi_{M}$ because the terms involving $\sin \left(\phi_{M}-\phi\right)$ are neglected in his Equation (54) for the Euler angle perturbations upon which the polar motion is based. The same approximation involving the $\sin \left(\phi_{M}-\phi\right)$ terms is made in deriving the Euler angle perturbations of Section 6, but there the polar motion is found first rather than being based upon the Euler angle perturbations as it is in Woolard's paper, and so the difficulty does not arise. The diurnal part of the Euler angle perturbations $\delta \theta_{z}$ and $\delta \psi_{z}$ is of the form

$$
\begin{equation*}
\sum_{j}\left(\delta \theta_{z_{j}}+i \delta \psi_{z_{j}} \sin \theta\right)=\sum_{j} \widetilde{A}_{j} e^{-i\left(\omega_{j} t+\beta_{j}-\phi\right)} \tag{8.28}
\end{equation*}
$$

Substituting (8.27) for the argument gives

$$
\begin{equation*}
\sum_{j}\left(\delta \theta_{z_{j}}+i \dot{\delta} \psi_{z_{j}} \sin \theta\right)=-e^{-i\left(\phi_{M}-\phi\right)} \sum_{j} \tilde{A}_{j} e^{-i a_{j}} \tag{8.29}
\end{equation*}
$$

The diurnal terms in Woolard's Equations (54) are equivalent to (8.29) with ( $\phi_{M}-\phi$ ) neglected. Substituting (8.29) with the ( $\phi_{M}-\phi$ ) term included into (8.26) gives

$$
\begin{equation*}
\sum_{j} m_{j}=i \frac{C}{A} \sum_{j} \tilde{A}_{j} e^{-i\left(\phi_{M}+a_{j}\right)} \tag{8.30}
\end{equation*}
$$

which contains $\phi_{M}$ in the arguments as it should. All terms involving the difference ( $\phi_{\mathrm{M}}-\phi$ ) multiplied by the polar motion components or the Euler angle perturbations are of second order and are therefore negligible in numerical computations. However, it is more exact as well as more straightforward to use $\phi_{\mathrm{M}}$ in the polar motion arguments instead of the Euler angle $\phi$.

The integral term in Equation (8.24) gives rise to the secular terms in Woolard's Equation (54) for $\delta \psi_{z}$ and in his polar motion equation (70). This integral term is effectively accounted for already by Woolard's use of the correct Eulerian solution (8.18) in place of (8.13) and therefore should not be
included in (8.24) at all. Fortunately, the small size of the secular term makes it negligible in numerical computations.

Of greater importance for numerical computations is the error inherent in Woolard's procedure for combining the complementary and particular solutions of the basic Equation (8.9). The nature of this error is readily determined by obtaining the particular integral of (8.9) with the help of expansion (4.21) of the lunisolar torque. Equation (8.9) is written as

$$
\begin{equation*}
\dot{x}-\mathbf{i}\left(\sigma_{r}+\ddot{\phi}\right) x=-\frac{1}{A} \sum_{j} A_{j} e^{-i\left(\omega_{j} t+\beta_{j}-\phi\right)} \tag{8.31}
\end{equation*}
$$

where $\sigma_{r}$ is given by (2.15). The particular integral of (8.31) is

$$
\begin{equation*}
x=-\sum_{\mathrm{j}} \frac{\mathrm{i} \mathrm{~A}_{\mathrm{j}}}{\mathbf{A}\left(\omega_{\mathrm{j}}+\sigma_{\mathrm{r}}\right)} \mathrm{e}^{-\mathrm{i}\left(\omega_{\mathrm{j}} \mathrm{t}+\beta_{\mathrm{j}}-\phi\right)} \tag{8.32}
\end{equation*}
$$

Woolard's procedure is to neglect the $\ddot{\chi}$ term in (8.31) and solve for $x_{\chi}$ to obtain

$$
\begin{equation*}
x \cong-\sum \frac{i A_{j}}{A\left(\ddot{\phi}+\sigma_{r}\right)} e^{-i\left(\omega_{j} t+\beta_{j}-\phi\right)} \tag{8.33}
\end{equation*}
$$

The effect is to replace each tidal frequency $\omega_{j}$ in the coefficients of (8.33) by the frequency $\ddot{\phi}$, thus neglecting the motion of the sun and moon. The frequency $\dot{\phi}$ is

$$
\begin{equation*}
\ddot{\phi}=\Omega-\dot{\psi} \cos \theta=\left[1+0\left(10^{-7}\right)\right] \Omega \tag{8.34}
\end{equation*}
$$

so that solutions equivalent to Woolard's can be obtained to within 1 part in $10^{7}$ by substituting

$$
\begin{equation*}
\omega_{j}=\Omega \tag{8.35}
\end{equation*}
$$

into the coefficients of the rigid-Earth formulas of Section 6. Equations (6.4) and (6.7) with approximation (8.35) are used to obtain the coefficients given in Table (8.1) for the diurnal motion of the rotation axis and the angular momentum vector. The six largest terms in the diurnal polar motion given by Woolard's Equations (70) correspond to the terms listed in Table (8.2). The amplitudes in Table (8.2) agree to within the number of significant figures retained by Woolard, thus providing a check on the rigid-Earth polar motion theory of Section 6.

Table 8.1
Coefficients for the Diurnal Motion of the Rotation Axis,
the Axis of Figure and the Angular Momentum Vector in
a Rigid Earth with Wollard's Approximation; $\omega_{j}=\Omega$.

COEFFICIENTS IN SECONOS DF ARC SIME FOR X-CTMPONENTS COSINE FOR Y-COMPONENTS

| Index | TIDAL ARG. CODE NUMBER | $\begin{aligned} & \text { BIST. } \\ & \text { BODY } \end{aligned}$ | PHIM | $\underset{\mathrm{L}}{\mathrm{COEF}}$ | $\begin{gathered} \mathrm{FFIC} \\ \mathrm{LP} \end{gathered}$ | $\begin{gathered} \text { IENTS } \\ \mathrm{F} \end{gathered}$ | $\begin{gathered} \text { DF } \\ \text { D. } \end{gathered}$ | (19 | ROTATION AXIS | AXIS MF FIGURE | ANGULAR MOMENTUM |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
| 1 | 105.955 | M | 1 | -4 | 0 | -2 | 0 | -2 | -0.0000018 | 0.0 | -0.0000018 |
| 2 | 107.755 | M | 1 | -2 | 0 | -2 | -2 | -2 | -0.0000075 | 0.0 | -0.0000075 |
| 3 | 109.555 | M | 1 | 0 | 0 | -2 | -4 | -2 | -0.0000046 | 0.0 | -0.0000046 |
| 4 | 115.845 | M | 1 | -3 | 0 | -2 | 0 | -1 | -0.0000034 | 0.0 | -0.0000034 |
| 5 | 115.855 | M | 1 | -3 | 0 | -2 | 0 | -2 | -0.0000177 | 0.0 | -0.0000177 |
| $t$ | 117.645 | M | 1 | -1 | 0 | -2 | -2 | -1 | -0.0000087 | 0.0 | -0.0000087 |
| 7 | 117.655 | M | 1 | -1 | 0 | -2 | -2 | -2 | -0.00004 56 | 0.0 | -0.0000455 |
| 8 | 118.654 | M | 1 | -1 | 1 | -2 | -2 | -2 | -0.0000034 | 0.0 | -0.0000034 |
| 9 | 119.445 | 14 | 1 | 1 | 0 | -2 | -4 | -1 | -0.0000016 | 0.0 | -0.0000016 |
| 10 | 119.455 | M | 1 | 1 | 0 | -2 | -4. | -2 | -0.0000089 | 0.0 | -0.0000088 |
| 11 | 124.756 | M | 1 | -2 | -1 | -2 | 0 | -2 | 0.0000021 | 0.0 | 0.0000021 |
| 12 | 125.745 | 14 | 1 | -2 | 0 | -2 | 0 | -1 | -0.0000295 | 0.0 | -0.0000294 |
| 13 | 125.755 | $\stackrel{1}{ }$ | 1 | -2 | 0 | -2 | 0 | -2 | -0.0001567 | 0.0 | -0.0001562 |
| 14 | $126.55 t$ | i. | 1 | 0 | -1 | -2 | -2 | -2 | 0.0000028 | 0.0 | 0.0000026 |
| 15 | 126.655 | i | 1 | -1 | 0 | -2 | -1 | -2 | 0.0000018 | 0.0 | 0.0000018 |
| $1 t$ | 126.754 | M | 1 | -2 | 1 | -2 | 0 | -2 | -0.0000025 | 0.0 | -0.0000025 |
| 17 | 127.545 | M | 1 | 0 | 0 | -2 | -2 | -1 | -0.0000358 | 0.0 | -0.0000357 |
| 18 | 127.555 | $\mathrm{M}_{4}$ | 1 | 0 | 0 | -2 | -2 | -2 | -0.0001892 | 0.0 | -0.0001886 |
| 19 | 12.8 .544 | M | 1 | 0 | 1 | -2 | -2 | -1 | -0.0000023 | 0.0 | -0.0000023 |
| 20 | 128.550 | M | 1. | 0 | 5 | -6 | 2 | -6 | -0.0000130 | 0.0 | -0.0000129 |
| 21 | 129.355 | M | 1 | 2 | 0 | -2 | -4 | -2 | -0.00000057 | 0.0 | -0.0000057 |
| 22 | 133.855 | M | 1 | -3 | 0 | -2 | 2 | -2 | 0.0000038 | 0.0 | 0.0000038 |
| 23 | 134.656 | $\cdots$ | 1 | -1 | -1 | -2 | 0 | -2 | 0.0000100 | 0.0 | 0.0000100 |
| 24 | 135.435 | M | 1 | 1 | 0 | -4 | 0 | -2 | 0.0000046 | 0.0 | 0.0000046 |
| 25 | 135.635 | M | 1 | -1 | 0 | -2 | 0 | 0 | $0.00000 \times 9$ | 0.0 | 0.0000069 |
| $2 t$ | 135.645 | M | 1 | -1 | 0 | -2 | 0 | -1 | -0.0002232 | 0.0 | -0.0002224 |
| 27 | 135.655 | M | 1 | -1 | 0 | -2 | 0 | -2 | -0.0011841 | 0.0 | -0.0011802 |
| 28 | 135.855 | 14 | 1 | -3 | 0 | 0 | 0 | 0 | 0.0000031 | 0.0 | 0.0000031 |
| 29 | 136.456 | M | 1 | 1 | -1 | -2 | -2 | -2 | 0.0000021 | 0.0 | 0.0000021 |
| 30 | 136.555 | M | 1 | 0 | 0 | -2 | -1 | -2 | 0.0000084 | 0.0 | 0.0000064 |
| 31 | 136.644 | M | 1 | -1 | 1 | -2 | 0 | -1 | -0.0000018 | 0.0 | -0.0000018 |
| 32 | 136.654 | M | 1 | -1 | 1 | -2 | 0 | -2 | -0.0000112 | 0.0 | -0.0000111 |
| 33 | 137.445 | M | 1 | 1 | 0 | -2 | -2 | -1 | -0.0000423 | 0.0 | -0.0000422 |
| 34 | 137.455 | M | 1 | 1 | 0 | -2 | -2 | -2 | -0.0002250 | 0.0 | -0.0002242 |
| 35 | 137.655 | M | 1 | -1 | 0 | $\bigcirc$ | -2 | 0 | 0.0000128 | 0.0 | 0.0000128 |
| 36 | 137.665 | M | 1 | $-1$ | 0 | 0 | -2 | -1 | -0.0000039 | 0.0 | -0.0000039 |
| 37 | 138.444 | ${ }^{\text {H }}$ | 1 | 1 | 1 | -2 | -2 | -1 | -0.0000018 | 0.0 | -0.0000018 |
| 38 | 138.454 | M | 1 | 1 | 1 | -2 | -2 | -2 | -0.0000105 | 0.0 | -0.0000105 |
| 39 | 139.455 | M | 1 | 1 | 0 | 0 | -4 | 0 | 0.0000023 | 0.0 | 0.0000023 |
| 40 | 143.535 | M | 1 | 0 | 0 | -4 | 2 | -2 | 0.0000028 | 0.0 | 0.0000029 |
| 41 | 143.745 | M | 1 | -2 | 0 | -2 | 2 | -1 | 0.0000033 | 0.0 | 0.0000033 |
| 42 | 143.755 | M | 1 | -2 | 0 | -2 | 2 | -2 | 0.0000185 | 0.0 | 0.0000185 |
| 43 | 144.546 | M | 1 | 0 | -1 | -2 | 0 | -1 | 0.0000025 | 0.0 | 0.0000025 |
| 44 | 144.55 t | M | 1 | 0 | -1 | -2 | 0 | -2. | 0.0000213 | 0.0 | 0.0000213 |
| 45 | 145.535 | M | 1 | 0 | 0 | -2 | 0 | 0 | 0.0000358 | 0.0 | 0.0000357 |
| $4 t$ | 145.545 | M | 1 | 0 | 0 | -2 | 0 | -1 | -0,0011659 | 0.0 | -0.0011621 |
| 47 | 145.555 | M | 1 | 0 | 0 | -2 | 0 | -2 | -0.0061846 | 0.0 | -0.006164.4 |
| 48 | 145.755 | M | 1 | -2 | 0 | 0 | 0 | 0 | 0.0000399 | 0.0 | 0.0000397 |
| 49 | 145.765 | M | 1 | -2 | 0 | 0 | 0 | -1 | 0.0000066 | 0.0 | 0.0000065 |
| 50 | 146.544 | M | 1 | 0 | 1 | -2 | 0 | -1 | -0.0000020 | 0.0 | -0.0000020 |
| 51 | 146.554 | M | 1 | 0 | 1 | -2 | 0 | -2 | -0.0000189 | 0.0 | -0.000018R |
| 52 | 147.355 | M | 1 | 2 | 0 | -2 | -2 | -2 | 0.0000034 | 0.0 | 0.0000034 |
| 53 | 147.545 | M | 1 | 0 | 0 | 0 | -2 | 1 | -0.0000023 | 0.0 | -0.0000023 |
| 54 | 147.555 | M | 1 | 0 | 0 | 0 | -2 | 0 | 0.0000806 | 0.0 | 0.0000803 |
| 55 | 147.565 | M | 1 | 0 | 0 | 0 | -2 | -1 | -0.0000176 | 0.0 | -0.0000175 |
| $5 t$ | 148.554 | 4 | 1 | 0 | 1 | 0 | -2 | 0 | 0.0000054 | 0.0 | 0.0000054 |
| 57 | 152.656 | M | 1 | -1 | -1 | -2 | 2 | -2 | 0.0000073 | 0.0 | 0.0000023 |
| 58 | 153.645 | H | 1 | -1 | 0 | -2 | 2 | -1 | 0.0000103 | 0.0 | 0.0000103 |
| 59 | 153.655 | M | 1 | -1 | 0 | -2 | 2 | -2 | 0.0000456 | 0.0 | 0.0000455 |
| $t 0$ | 154.656 | H | 1 | -1 | -1 | 0 | 0 | 0 | -0.0000025 | 0.0 | -0.0000025 |

Table 8.1-(continued)
COEFFICIENTS IN SECONÖS DF AÖC SINE FOR $X$-COMPONENTS COSINE FOR Y-CDMPONENTS

| INDEX | TIDAL ARG. CDDE NUMBER | $\begin{aligned} & \text { DIST. } \\ & \text { BODY } \end{aligned}$ | PHIM | COEFFICIENTS |  |  | OF |  | $\begin{aligned} & \text { ROTATION } \\ & \text { AXIS } \end{aligned}$ | AXIS MF FIGURE | ANGULAR <br> MDMENTUM |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  |  |  | 1 L | LP | F | D | OM |  |  |  |
| 81 | 155.435 | M | 1 | 1 | 0 | -2 | 0 | 0 | -0.0000028 | 0.0 | -0.0000028 |
| $t 2$ | 155.445 | M | 1 | 1 | 0 | -2 | 0 | -1 | 0.0000323 | 0.0 | 0.0000322 |
| 63 | 155.455 | M | 1 | 1 | 0 | -2 | 0 | -2 | 0.0001748 | 0.0 | 0.0001742 |
| $t 4$ | 155.645 | M | 1 | -1 | 0 | 0 | 0 | 1 | -0.0000139 | 0.0 | -0.0000139 |
| 65 | 155.655 | M | 1 | -1 | 0 | 0 | 0 | 0 | 0.0004864 | 0.0 | 0.0004848 |
| t $\epsilon$ | 155.665 | M | 1 | -1 | 0 | 0 | 0 | -1 | 0.0000975 | 0.0 | 0.0000972 |
| 67 | 155.675 | M | 1 | -1 | 0 | 0 | 0 | -2 | -0.0000028 | 0.0 | -0.0000028 |
| 68 | 156.555 | M | 1 | 0 | 0 | 0 | -1 | 0 | -0.0000026 | 0.0 | -0.0000026 |
| 69 | 156.654 | M | 1 | -1 | 1 | 0 | 0 | 0 | 0.0000030 | 0.0 | 0.0000029 |
| 70 | 157.445 | M | 1 | 1 | 0 | 0 | -2 | 1 | -0.0000026 | 0.0 | $-0.0000026$ |
| 71 | 157.455 | M | 1 | 1 | 0 | 0 | -2 | 0 | 0.0000929 | 0.0 | 0.0000926 |
| 72 | 157.465 | M | 1 | 1 | 0 | 0 | -2 | -1 | 0.0000203 | 0.0 | 0.0000203 |
| 73 | 158.454 | M | 1 | 1 | 1 | 0 | -2 | 0 | 0.0000039 | 0.0 | 0.0000039 |
| 74 | $1 \in 1.557$ | S | 1 | 0 | -2 | -2 | 2 | -2 | -0.0000069 | 0.0 | -0.0000069 |
| 75 | $1 \in 2.556$ | S | 1 | 0 | -1 | -2 | 2 | -2 | -0.0001692 | 0.0 | -0.0001686 |
| 76 | 163.535 | M | 1 | 0 | 0 | -2 | 2 | 0 | -0.0000023 | 0.0 | -0.0000023 |
| 77 | 163.545 | M | 1 | 0 | 0 | -2 | 2 | -1 | 0.0000327 | 0.0 | 0.0000325 |
| 78 | 143.555 | M | 1 | 0 | 0 | -2 | 2 | -2 | -0.0000049 | 0.0 | -0.0000049 |
| 79 | 163.555 | 5 | 1 | 0 | 0 | -2 | 2 | -2 | -0.0028864 | 0.0 | -0.0028769 |
| 80 | 143.557 | 5 | 1 | 0 | -2 | 0 | 0 | 0 | 0.0000018 | 0.0 | 0.0000018 |
| 81 | 163.755 | M | 1 | -2 | 0 | 0 | 2 | 0 | 0.0000043 | 0.0 | 0.0000043 |
| 82 | 164.554 | 5 | 1 | 0 | 1 | -2 | 2 | -2 | 0.0000241 | 0.0 | 0.0000241 |
| 83 | 164.556 | S | 1 | 0 | -1 | 0 | 0 | 0 | 0.0000696 | 0.0 | 0.0000693 |
| 84 | 145.545 | M | 1 | 0 | 0 | 0 | 0 | 1. | -0.0001723 | 0.0 | -0.0001717 |
| 85 | 145.555 | S | 1 | 0 | 0 | 0 | 0 | 0 | 0.0027652 | 0.0 | 0.0027561 |
| 86 | 165.555 | M | 1 | 0 | 0 | 0 | 0 | 0 | 0.0059457 | 0.0 | 0.0059262 |
| 87 | 165.565 | M | 1 | 0 | 0 | 0 | 0 | -1 | 0.0011785 | 0.0 | 0.0011747 |
| 88 | 165.575 | M | 1 | 0 | 0 | 0 | 0 | -2 | -0.0000253 | 0.0 | -0.0000252 |
| 89 | 166.554 | S | 1 | 0 | 1 | 0 | 0 | 0 | 0.0000696 | 0.0 | 0.0000693 |
| 90 | 167.355 | M | 1 | 2 | 0 | 0 | -2 | 0 | 0.0000043 | 0.0 | 0.0000043 |
| 91 | 167.553 | S | 1 | 0 | 2 | 0 | 0 | 0 | 0.0000018 | 0.0 | 0.0000019 |
| 92 | 167.555 | S | 1 | 0 | 0 | 2 | -2 | 2 | 0.0001243 | 0.0 | 0.0001239 |
| 93 | 167.565 | M | 1 | 0 | 0 | 2 | -2 | 1 | -0.0000048 | 0.0 | -0.0000047 |
| 94 | 167.575 | M | 1 | 0 | 0 | 2 | -2 | 0 | -0.0000023 | 0.0 | -0.0000023 |
| 95 | 168.554 | S | 1 | 0 | 1 | 2 | -2 | 2 | 0.0000073 | 0.0 | 0.0000072 |
| 96 | 172.656 | M | 1 | -1 | -1 | 0 | 2 | 0 | 0.0000039 | 0.0 | 0.0000039 |
| 97 | 173.445 | M | 1 | 1 | 0 | -2 | 2 | -1 | 0.0000028 | 0.0 | 0.0000028 |
| 98 | 173.645 | M | 1 | -1 | 0 | 0 | 2 | 1 | -0.0000030 | 0.0 | -0.0000029 |
| 99 | 173.655 | M | 1 | -1 | 0 | 0 | 2 | 0 | . 0.0000929 | 0.0 | 0.0000926 |
| 100 | 173.665 | M | 1 | -1 | 0 | 0 | 2 | -1 | 0.0000184 | 0.0 | 0.0000183 |
| 101 | 174.456 | M | 1 | 1 | -1 | 0 | 0 | 0 | D. 0000030 | 0.0 | 0.0000029 |
| 102 | 174.555 | M | 1 | 0 | 0 | 0 | 1 | 0 | -0.0000026 | 0.0 | -0.0000026 |
| 103 | 175.445 | M | 1 | 1 | 0 | 0 | 0 | 1 | -0.0000143 | 0.0 | -0.0000142 |
| 104 | 175.455 | M | 1 | 1 | 0 | 0 | 0 | 0 | 0.0004864 | 0.0 | 0.0004848 |
| 105 | 175.465 | M | 1 | 1 | 0 | 0 | 0 | -1 | 0.0000963 | 0.0 | 0.0000960 |
| $10 t$ | 175.475 | M | 1 | 1 | 0 | 0 | 0 | -2 | -0.0000021 | 0.0 | -0.0000021 |
| 107 | 175.655 | M | 1 | -1 | 0 | 2 | 0 | 2 | -0.00000 75 | 0.0 | -0.0000075 |
| 108 | 175.665 | M | 1 | -1 | 0 | 2 | 0 | 1 | -0.0000048 | 0.0 | -0.0000047 |
| 109 | 175.675 | M | 1 | -1 | 0 | 2 | 0 | 0 | -0.0000028 | 0.0 | -0.0000028 |
| 110 | 176.454 | M | 1 | 1 | 1 | 0 | 0 | 0 | -0.0000025 | 0.0 | -0.0000025 |
| 111 | 177.455 | M | 1 | 1 | 0 | 2 | -2 | 2 | -0.0000020 | 0.0 | -0.0000020 |
| 112 | 182.556 | M | 1 | 0 | -1 | 0 | 2 | 0 | 0.0000053 | 0.0 | 0.0000052 |
| 113. | 183.545 | M | 1 | 0 | 0 | 0 | 2 | 1 | 0.0000026 | 0.0 | 0.0000026 |
| 114 | 183.555 | M | 1 | 0 | 0 | 0 | 2 | 0 | 0.0000807 | 0.0 | 0.0000805 |
| 115 | 183.565 | M | 1 | 0 | 0 | 0 | 2 | -1 | 0.0000158 | 0.0 | 0.0000157 |
| 116 | 185.355 | M | 1 | 2 | 0 | 0 | 0 | 0 | 0.0000394 | 0.0 | 0.0000393 |
| 117 | 185.365 | M | 1 | 2 | 0 | 0 | 0 | -1 | 0.0000079 | 0.0 | 0.0000079 |
| 118 | 185.555 | M | 1 | 0 | 0 | 2 | 0 | 2 | 0.0002683 | 0.0 | 0.0002655 |
| 119 | 185.565 | M | 1 | 0 | 0 | 2 | 0 | 1 | 0.0001705 | 0.0 | 0.0001699 |
| 120 | 185.575 | M | 1 | 0 | 0 | 2 | 0 | 0 | 0.0000358 | 0.0 | 0.0000357 |
| 121 | 185.585 | M | 1 | 0 | 0 | 2 | 0 | -1 | 0.0000023 | 0.0 | 0.0000023 |
| 122 | 191.655 | M | 1 | $-1$ | 0 | 0 | 4 | 0 | 0.0000025 | 0.0 | 0.0000025 |
| 123 | 193.455 | M | 1 | 1 | 0 | 0 | 2 | 0 | 0.0000128 | 0.0 | 0.0000128 |
| 124 | 193.465 | M | 1 | 1 | 0 | 0 | 2 | -1 | 0.0000025 | 0.0 | 0.0000025 |
| 125 | 193.655 | M | 1 | -1 | 0 | 2 | 2 | 2 | 0.0000097 | 0.0 | 0.0000096 |
| 126 | 193.665 | M | 1 | -1 | 0 | 2 | 2 | 1 | 0.0000062 | 0.0 | 0.0000062 |
| 127 | 195.255 | M | 1 | 3 | 0 | 0 | 0 | 0 | 0.0000031 | 0.0 | 0.0000031 |
| 128 | 195.455 | M | 1 | 1 | 0 | 2 | 0 | 2 | 0.0000510 | 0.0 | 0.0000509 |
| 129 | 195.465 | M | 1 | 1 | 0 | 2 | 0 | 1 | 0.0000327 | 0.0 | 0.0000325 |

## Table 8.1-(continued)

# COEFFICIENTS IN SECONDS OF ARC <br> SINE FOR $X$-COMPDNFNTS 

COSINE FDR Y-COMPONENTS

| IMDEX | TIDAL ARG. CODE NUMBER | $\begin{aligned} & \text { DIST. } \\ & \text { BDDY } \end{aligned}$ | PHIM | $\underset{L}{\operatorname{COEF}}$ | $L P$ | $\underset{F}{\text { NTS }}$ | $\begin{array}{r} \text { OF } \\ \mathrm{O} \end{array}$ | DM | ROTATION AXIS. | AXIS OF FIGURE | ANGULAR MOMENJUM |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
| 130 | 195.475 | M | 1 | 1 | 0 | 2 | 0 | 0 | 0.0000069 | 0.0 | 0.0000069 |
| 131 | $1 \times 3.555$ | M | 1 | 0 | 0 | 2 | 2 | 2 | 0.0000082 | 0.0 | 0.0000082 |
| 132 | $1 \times 3.565$ | M | 1 | 0 | 0 | 2 | 2 | 1 | 0.0000053 | 0.0 | 0.0000052 |
| 133 | $1 \times 5.355$ | M | 1 | 2 | 0 | 2 | 0 | 2 | 0.0000067 | 0.0 | 0.0000067 |
| 134 | $1 \times 5.365$ | M | 1 | 2 | 0 | 2 | 0 | 1 | 0.0000044 | 0.0 | 0.0000044 |
| 135 | 1E3.455 | M | 1 | 1 | 0 | 2 |  | 2 | 0.0000020 | 0.0 | 0.0000020 |

```
Explanation of symbola
    PHIM la the Greenwich mesn alderial time; $%
    L, LP, F, D and OM are:Brown's fundamental
    argumente; l, l', F, D and D
Constants
    k=0
    K, =-7!552430 }\times1\mp@subsup{0}{}{3}\mp@subsup{\textrm{Jullan}}{\mathrm{ century }}{}\mp@subsup{}{}{-1
    K
    (C-A)/C = 3.272930 * 10-3
    n=360:9856 day }\mp@subsup{}{}{-1
    (C0}/2)=2.343852\times10
    J}=1,082645\times1\mp@subsup{0}{}{-3
    (m}/\mp@subsup{m}{\textrm{g}}{}/\mp@subsup{m}{E}{})=3,334320\times1\mp@subsup{0}{}{0
```

The effect of Woolard's approximation (8.35) on numerical computations of the diurnal variation of latitude for Goddard Space Flight Center is shown in Figure (8.1). The diurnal terms in Equation (6.4) are substituted into Equation (9.14) in order to compute the latitude variation. The amplitudes differ by at most $0!0008$ which is $3.7 \%$ of the diurnal polar motion amplitude or 2.5 cm at the Earth's surface.

Table 8.2
Comparison of Woolard's Diurnal Polar Motion Amplitudes with Those from Table 8.1

| Tidal Argument <br> Code Number | Amplitude from <br> Woolard's Equations <br> $(70)$ | Amplitude from <br> the Column Headed <br> "Rotation Axis"in <br> Table 8.1 |
| :---: | :---: | :---: |
| 135.655 | $-0!\prime 0012$ | -0.0011841 |
| 145.545 | -0.0012 | -0.0011659 |
| 145.555 | -0.0062 | -0.0061846 |
| 163.555 | -0.0029 | -0.0028913 |
| 165.555 | 0.0087 | 0.0087109 |
| 165.565 | 0.0012 | 0.0011785 |

The Euler angle perturbations in Table 8.3 are obtained from Equations (6.19) through (6.24) with $\omega_{j}=\Omega$ so as to make the results equivalent to Woolard's Equation (54). The amplitudes of the largest terms in $\delta \theta_{z}$ and $\delta \psi_{z} \sin \theta$ from Table 8.3 are compared with Woolard's Equation (54) in Table 8.4. The amplitudes agree to as many significant figures as Woolard retains, thus providing a check on the Euler angle perturbation theory of Section 6.


Figure 8.1. Effect of Woolard's Approximation on the Diurnal Variation of Latitude for
Goddard Space Flight Center. Both Curves are for a Rigid Earth.

Table 8-3
Coefficients for Perturbations in the Euler Angles of the Rotation Axis, the Axis of Figure and the Angular Momentum Vector in a Rigid Earth with Woolard's Approximation; $\omega_{j}=\Omega$

CDEFFICIENTS FOR EULER ANGLE PERTURBATIONS
IN SECONOS OF ARC
cosines for del.ta theta
SINES FMR DELTA PSI * SINITHETAI

| TIDAL COOE | ARGUMENT NUMBERS | CDEFFICIENTS |  |  |  | DF | $\begin{gathered} \text { ROTATIOM } \\ \text { AXIS } \end{gathered}$ |  | AXIS OF FIGURE |  | TERRESTRIAL Z AXIS |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  | $L$ | LP | F | D | OM | cos | S jN | cos | SIN | $\cos$ | SIM |
| 105.955 | - | 4 | 0 | 2 | 0 | 2 | -0.00000001 | -0.0000000-1 | 0.00000180 | 0.00000180 | 0.00000180. | 0.00000180 |
| 107.755 | - | 2 | 0 | 2 | 2 | 2 | -0.00000002 | -0.00000002 | 0.00000752 | 0.00000752 | 0.00000752 | $0.0000075 ?$ |
| 109.555 |  | 0 | 0 | 2 | 4 | 2 | -0.00000002 | -0.00000002 | 0.00000458 | 0.00000458 | 0.00000458 | 0.00000458 |
| 115.845 |  | 3 | 0 | 2 | 0 | 1 | -0.00000001 | -0.00000001 | 0.00000343 | 0.00000343 | 0.00000343 | 0.00000343 |
| 115.855 |  | 3 | 0 | 2 | 0 | 2 | -0.00000006 | -0.00000006 | 0.00001766 | 0.00001766 | 0.00001766 | 0.00001766 |
| 117.645 |  | 1 | 0 | 2 | 2 | 1 | -0.00000003 | -0.00000003 | 0.00000867 | 0.00000887 | 0.00000867 | 0.00000867 |
| 117.655 | . | 1 | 0 | 2 | 2 | 2 | -0.00000015 | -0.00000015 | 0.00004547 | 0.00004547 | 0.00004547 | 0.00004547 |
| 118.654 | - | 1 | -1 | 2 | 2 | 2 | -0.00000001 | -0.00000001 | 0.00000343 | 0.00000343 | 0.00000343 | 0.00000343 |
| 119.445 |  | -1 | 0 | 2 | 4 | 1 | -0.00000001 | -0.00000001 | 0.00000164 | 0.00000164 | 0.00000164 | 0.00000164 |
| 119.455 | . | -1 | 0 | 2 | 4 | 2 | -0.00000003 | -0.00000003 | 0.00000883 | 0.00000883 | 0.00000883 | 0.00000883 |
| 124.75 t | - | 2 | 1 | 2 | 0 | 2 | 0.00000001 | 0.00000001 | -0.00000213 | -0.00000213 | -0.00000213 | -0.00000213 |
| 125.745 | $1 \times 5.365$ | 2 | 0 | 2 | 0 | 1 | -0.00000008 | -0.00000011 | 0.00002502 | 0.00003386 | 0.00002502 | 0.000033 ab |
| 125.755 | $1 \times 5.355$ | 2 | 0 | 2 | 0 | 2 | -0.00000049 | -0.00000053 | 0.00014949 | 0.00016290 | 0.00014949 | 0.00016291 |
| $12 t .55 t$ | . | 0 | 1 | 2 | 2 | 2 | 0.00000001 | 0.00000001 | -0.00000262 | -0.00000262 | -0.00000262 | -0.00000262 |
| 12t. 655 | . | 1 | 0 | 2 | 1 | 2 | 0.00000001 | 0.00000001 | -0.00000180 | -0.00000180 | -0.00000180 | -0.00000180 |
| 12 t. 754 | . | 2 | -1 | 2 | 0 | 2 | -0.00000001 | -0.00000001 | 0.00000245 | 0.00000245 | 0.00000245 | 0.00000245 |
| 127.545 | $1 \times 3.565$ | 0 | 0 | 2 | 2 | 1 | -0.00000010 | -0.00000013 | 0.00003042 | 0.00004089 | 0.00003042 | 0.00004089 |
| 127.555 | $1 \times 3.555$ | 0 | 0 | 2 | 2 | 2 | -0.00000059 | -0.00000006 | 0.00018040 | 0.00019676 | 0.00018041 | 0.00019676 |
| 128.544 |  | 0 | -1 | 2 | 2 | 1 | -0.00000001 | -0.0000000 1 | 0.00000229 | 0.00000229 | 0.00000229 | 0.00000229 |
| 128.550 |  | 0 | -5 | 6 | -2 | $t$ | -0.00000004 | -0.00000004 | 0.00001292 | 0.00001792 | 0.00001292 | 0.00001292 |
| 129.355 | . | -2 | 0 | 2 | 4 | 2 | -0.00000002 | -0.0000000? | 0.00000572 | 0.00000572 | 0.00000572 | 0.00000572 |
| 133.855 |  | 3 | 0 | 2 | -2 | 2 | 0.00000001 | 0.00000001 | -0.00000376 | -0.00000376 | -0.00000376 | -0.00000376 |
| 134.656 |  | 1 | 1 | 2 | 0 | 2 | 0.00000003 | 0.00000003 | -0.00000998 | -0.00000998 | -0.00000998 | -0.00000998 |
| 135.435 |  | -1 | 0 | 4 | , | 2 | 0.00000002 | 0.00000002 | -0.00000458 | -0.00000458 | -0.00000458 | -0.00000458 |
| 135.t35 | 195.475 | 1 | 0 | 2 | 0 | 0 | 0.00000005 | 0.0 | -0.00001374 | 0.0 | -0.00001374 | 0.0 |
| 135.645 | 195.465 | 1 | 0 | 2 | 0 | 1 | -0.00000062 | -0.00000084 | 0.00018989 | 0.00025499 | 0.00018989 | 0.00025499 |
| 135.655 | 195.455 | 1 | 0 | 2 | 0 | 2 | -0.00000371 | -0.00000404 | 0.00112936 | 0.00123109 | 0.00112938 | 0.00123111 |
| 135.855 | 195.255 | 3 | 0 | 0 | 0 | 0 | 0.00000002 | 0.0 | -0.00000622 | 0.0 | -0.00000622 | 0.0 |
| 13 t.45t |  | $-1$ | 1 | 2 | 2 | 2 | 0.00000001 | 0.00000001 | -0.00000213 | -0.00000213 | -0.00000213 | -0.00000213 |
| 136.555 |  | 0 | 0 | 2 | 1 | 2 | 0.00000002 | 0.00000002 | -0.00000638 | -0.0000063B | -0.00000638 | -0.00000638 |
| 136.644 |  | 1 | -1 | 2 | 0 | 1 | -0.0000000 1 | -0.00000001 | 0.00000180 | 0.00000180 | 0.00000180 | 0.00000180 |
| 136.654 |  | 1 | -1 | 2 | 0 | 2 | -0.00000004 | -0.00000004 | 0.00001112 | 0.00001112 | 0.00001112 | 0.00001112 |
| 137.445 | 193.665 | -1 | 0 | 2 | 2 | 1 | -0.00000012 | -0.00000016 | 0.00003598 | 0.00004841 | 0.00003598 | 0.00004841 |
| 137.455 | 193.655 | -1 | 0 | 2 | 2 | 2 | -0.00000070 | -0.0n0000 77 | 0.00021459 | 0.00023389 | 0.00021459 | 0.00023389 |
| 137.t55 | 193.455 | 1 | 0 | 0 | 2 | 0 | 0.00000008 | 0.0 | -0.00002551 | 0.0 | -0.00002552 | 0.0 |
| 137.665 |  | 1 | 0 | 0 | 2 | 1 | -0.00000001 | -0.00000001 | 0.00000393 | 0.00000393 | 0.00000393 | 0.00000393 |
| 138.444 |  | -1 | -1 | 2 | 2 | 1 | -0.00000001 | -0.0000000.1 | 0.00000180 | 0.00000180 | 0.00000180 | 0.00000180 |
| 138.454 | . | -1 | -1 | 2 | 2 | 2 | -0.00000003 | -0.00000003 | 0.00001047 | 0.00001047 | 0.00001047 | 0.00001047 |
| 139.455 | 191.655 | -1 | 0 | 0 | 4 | 0 | 0.00000002 | -0.00000000 | -0.00000474 | 0.00000016 | -0.00000474 | 0.00000016 |
| 143.535 |  | 0 | 0 | 4 | -2 | 2 | 0.00000001 | 0.00000001 | -0.00000278 | -0.0000027\% | -0.00000278 | -0.00000278 |

Table 8-3-(continued)
COEFFICIFATS FIR EILER ANGLE PFRTIRAATIONS
IN SECMNDS OF ARC
cosines for nelta theta
SINES FIR MELTA PSI * SIN(THFTA)

| $\begin{array}{r} \text { TIOAL } \\ \text { CODE } \end{array}$ | ARGUMENT NUABERS | COEFFICIENTS |  |  |  | OF | $\begin{aligned} & \text { ROTATIOM } \\ & \text { AXIS } \end{aligned}$ |  | AXIS nf FIGURF |  | $\begin{aligned} & \text { TERRESTRIAL } \\ & Z \text { AXIS } \end{aligned}$ |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  | L | LP | $F$ | D | OM | CDS | SIM | cos | SIM | cos | 515 |
| 143.745 |  | 2 | 0 | 2 | -2 | 1 | 0.00000001 | 0.00000001 | -0.00000327 | -0.00n00327 | -0.00000327 | -0.00000327 |
| 143.755 |  | 2 | 0 | 2 | -2 | 2 | 0.00000006 | 0.00000006 | -0.00001848 | -0.00nol84R | -0.00001848 | -0.00001848 |
| 144.546 |  | 0 | 1 | 2 | 0 | 1 | 0.00000001 | 0.00000001 | -0.00000245 | -0.00000245 | -0.00000245 | -0.000cmoz 5 |
| -144.55t |  | 0 | 1 | 2 | 0 | 2 | 0.00000007 | 0.00000007. | -0.00002126 | -0.00002126 | -0.00002126 | -0.00no?126 |
| 145.535 | 185.575 | 0 | 0 | 2 | 0 | 0 | 0.00000023 | 0.0 | -0.00007131 | 0.0 | -0.00007131 | 0.0 |
| 145.545 | 185.565 | 0 | 0 | 2 | 0 | 1 | -0.00000326 | -0.00000437 | 0.00099214 | 0.00133201 | 0.00099215 | 0.00133203 |
| 145.555 | 185.555 | 0 | 0 | 2 | 0 | 2 | -0.00001937 | -0.00002111 | 0.00589886 | 0.00642977 | 0.00589893 | 0.0n64ア934 |
| 145.755 | 185.355 | 2 | 0 | 0 | 0 | 0 | 0.00000026 | 0.00000000 | -0.000.07900 | -n.000000449 | -0.00007900 | -n.00rumors |
| 145.765. |  | 2 | 0 | 0 | 0 | 1 | 0.00000007 | 0.00000002 | -0.00000654 | -0.00000654 | -0.00000654 | -0.000)00654 |
| $14 t .544$ |  | 0 | -1 | 2 | 0 | 1 | -0.00000001 | $-0.00000001$ | 0.00000196 | 0.00000196 | 17.00000196 | 0.00000196 |
| 146.554 |  | 0 | -1 | 2 | 0 | 2 | -0.00000006 | -0.00000006 | 0.00001881 | 0.00001821 | 0.00001 trl | 0.00nolatil |
| 147.355 |  | -2. | 0 | 2 | 2 | 2 | 0.00000001 | 0.00000001 | -0.00000343 | -0.00000343 | -0.00000343 | -0.0000034.3 |
| 147.545 | 183.565 | 0 | 0 | 0 | 2 | -1 | 0.00000004 | -0.00000006 | -0.00001341 | 0.00001799 | -0.00001341 | 0.00001799 |
| 147.555 | 183.555 | 0 | 0 | 0 | 2 | 0 | 0.00000053 | -0.00000000 | -0.00016078 | $0.00000 n 16$ | -0.00016078 | 0.0000 mal 6 |
| 147.565 | 183.545 | 0 | 0 | 0 | 2 | 1 | -0.00000005 | -0.00000007 | 0.00001483 | 0.00002012 | 0.00001488 | $0.0000 \%$ 012. |
| 148.554 | 182.55 t | 0 | -1 | 0 | 2 | 0 | 0.00000003 | 0.00000000 | -0.00001063 | -0.00000016 | -0.00001u63 | -r.onoonolo |
| 152.456 | . | 1 | 1 | 2 | -2 | 2 | 0.00000001 | 0.00000001 | -0.00000229 | -0.00000229 | -0.00000229 | -0.00090229 |
| 153.645 | - | 1 | 0 | 2 | -2 |  | 0.00000003 | 0.00000003 | -0.00001030 | -0.00001030 | -0.00001030 | -0.00001030 |
| 153.t55 | 177.455 | 1 | 0 | 2 | -2 | 2 | 0.00000014 | 0.00000016 | -0.00004351 | -0.0000474.3 | -0.00004351 | -0.00004743 |
| $154.65 t$ | 176.454 | 1 | 1 | 0 | 0 | D | -0.00000002 | 0.0 | 0.00000491 | 0.0 | 0.00000491 | 0.0 |
| 155.435 | 175.675 | -1 | 0 | 2 | 0 | 0 | -0.00000002 | 0.0 | 0.00000556 | 0.0 | 0.00000556 | 0.0 |
| 155.445 | 175.66.5 | -1 | 0 | 2 | 0 | 1 | 0.00000009 | 0.000000012 | -0.00002748 | -0.00003886 | -0.00002748 | -0.00003696 |
| 155.455 | 175.655 | -1 | 0 | 2 | 0 | 2 | 0.00000055 | 0.00000060 | -0.00016666 | -0.00018171 | -0.00016667 | -0.01019171 |
| 155.645 | 175.485 | 1 | 0 | 0 | 0 | -1 | 0.00000027 | -0.00000036 | -0.00002?11 | 0.00010991 | -0.00008č11 | V.0001(1991 |
| 155.t55 | 175.455 | 1 | 0 | 0 | 0 | 0 | 0.00000318 | 0.0 | -0.00096957 | 0.0 | -0.00096958 | 0.0 |
| 155.6.65 | 175.445 | 1 | 0 | 0 | 0 | 1 | 0.00000027 | 0.00000037 | -0.00008292 | -0.00011138 | -0.00008292 | $-0.0001113 \%$ |
| 155.t75 | - | 1 | 0 | 0 | 0 | 2 | -0.00000001 | -0.00000001 | 0.00000278 | 0.000000278 | 0.00000278 | $0.0001027 \%$ |
| $15 t .555$ | 174.555 | 0 | 0 | 0 | 1 | 0 | -0.00000002 | 0.0 | 0.00000523 | 0.0 | 0.00000523 | 0.0 |
| 156.654 | 174.45 t | 1 | -1 | 0 | 0 | 0 | 0.00000002 | 0.0 | -0.00000589 | 0.0 | -0.00000589 | 0.0 |
| 157.445 | 173.685 | -1 | 0 | 0 | 2 | -1 | 0.00000005 | -0.00000007 | -0.00001570 | 0.00002094 | -0.00001570 | 0.00002094 |
| 157.455 | 173.655 | -1 | 0 | 0 | 2 | 0 | 0.00000061 | 0.0 | -0.00018515 | 0.0 | -0.00018515 | . 0.0 |
| 157.465 | 173.645 | -1 | 0 | 0 | 2 | 1 | 0.00000006 | 0.00000008 | -0.00001734 | -0.00002323 | -0.00001734 | -0.00002323 |
| 158.454 | 172.65t | -1 | - | 0 | 2 | 0 | 0.00000003 | 0.0 | -0.00000785 | 0.0 | -0.00000785 | 0.0 |
| 161.557 |  | 0 | 2 | 2 | -2 | 2 | -0.00000002 | -0.001000002 | 0.00000687 | 0.00000687 | 0.00000687 | 0.00000687 |
| 162.556 | 168.554 | 0 | 1 | 2 | -2 | 2 | -0.00000053 | -0.00000058 | 0.00016140 | 0.00017588 | 0.00016140 | 0.00017588 |
| 163.535 | 167.575 |  | 0 | 2 | -2 | 0 | -0.00000002 | 0.0 | 0.00000458 | 0.0 | 0.00000458 | 0.0 |
| 163.545 | 187.545 |  | 0 | 2 | -2 | 1 | 0.00000009 | 0.00000012 | -0.00002780 | -0.00003729 | -0.00002781 | -0.000033729 |
| 163.555 |  | 0 | 0 | 2 | -2 | 2 | -0.00000002 | -0.00000002 | 0.00000491 | 0.00000491 | D.00000491 | 0.00000491 |
| 163.555 | 167.555 | 0 | 0 | 2 | -2 | 2 | -0.00000904 | -0.00000985 | 0.00275300 | 0.00300079 | 0.00275303 | 0.00300082 |
| 163.557 | 167.553 | 0 | 2 | 0 |  | 0 | 0.00000001 | 0.0 | -0.00000362 | 0.0 | -0.00000362 | 0.0 |

Table 8-3-(continued)
COEFFICIENTS FDR EILER ANGLE PERTURBATIONS IN SECONDS DF ARC

COSINES fOR DELTA THETA
SINES FOR DELTA PSI * SIMTHETAJ

| TIOAL CODE | ARGUMENT NUMBERS | L | CoEFFICIENTS |  |  | OF | ROTATION$\Delta \times I S$ |  | AXIS MF FIGURE |  | TERRESTRIAL $z$-axis |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  |  | LP | F | D | OM | $\cos$ | SIN | cos | SIN | cos | SIN |
| 163.755 | 167.355 | 2 | 0 | 0 | -2 | 0 | 0.00000003 | 0.0 | -0.00000R.50 | 0.0 | -0.00000851 | 0.0 |
| 164.554 | - | 0 | -1 | 2 | -2 | 2 | 0.00000008 | 0.00000008 | -0.00002407 | -0.00002407 | -0.00002407 | $-0.00002407$ |
| 164.556 | 166.554 | 0 | 1 | 0 | 0 | 0 | D. 00000046 | 0.0 | -0.00013R68 | 0.0 | -0.00013869 | 0.0 |
| 165.545 | 165.5 t-5 | 0 | 0 | 0 | 0 | -1 | 0.00000329 | -0.00000442 | -0.00100293 | 0.00134640 | -0.00100294 | 0.00134642 |
| 165.555 | * | 0 | 0 | 0 | 0 | 0 | 0.00000905 | 0.00000905 | -0.00275609 | -0.00275609 | -0.00275613 | -0.00275613 |
| 165.555 | 105* | 0 | 0 | 0 | 0 | 0 | 0.00001946 | 0.00001946 | -0.00592617 | -0.00592617 | -0.00592624 | -0.00592624 |
| . | 165.575 | 0 | 0 | 0 | D | -2 | -0.00000008 | 0.00000008 | 0.0000251 .9 | -0.00002519 | . 0.00002519 | -0.00002519 |
| - | 173.445 | 1 | 0 | -2 | 2 | -1 | 0.00000001 | -0.00000001 | -0.00000278 | 0.00000278 | -0.00000278 | 0.00000273 |
| - | 175.475 | 1 | 0 | 0 | 0 | -2 | -0.00000001 | 0.00000001 | 0.00000213 | -0.00000213 | 0.00000213 | -0.00000213 |
| - | 185.365 | 2 | 0 | 0 | 0 | -1 | 0.00000003 | -0.00000003 | -0.00000785 | 0.00000785 | -0.00000785 | 0.00000785 |
| - | 185.585 | 0 | 0 | 2 | 0 | -1 | 0.00000001 | -0.00000001 | -0.00000229 | 0.00000229 | -0.00000229 | 0.00000229 |
| - | 193.465 | 1 | 0 | 0 | 2 | -1 | 0.00000001 | -0.00000001 | -0.00000245 | 0.00000245 | -0.00000225 | 0.00000245 |
| - | 163.455 | 1 | 0 | 2 | 2 | 2 | 0.00000001 | -0.00000001 | -0.00000196 | 0.00000196. | -0.00000196 | 0.00000196 |

## Explanntion of symbols

L, LP, F, D and OM are Brown'a fundamental arguments; $\ell^{\prime}, \ell^{\prime}, F, D$ and $\cap$

$\mathrm{K}_{\mathrm{s}}=-7,7552430 \times 10^{3}$ Jullan century ${ }^{-1}$
$K_{e}=-3!484150 \times 10^{3} \mathrm{Juliman}^{2}$ century ${ }^{-1}$
$(\mathrm{C}-\mathrm{A}) / \mathrm{C}=3.272990 \times 10^{-3}$
$\mathrm{n}=360: 9856 \mathrm{day}^{-1}$
$\begin{aligned}\left(\mathrm{C}_{\varepsilon} / \mathrm{a}_{\mathrm{\Sigma}}\right) & =2.343852 \times 10^{4} \\ J_{2} & =1.082645 \times 10^{-3}\end{aligned}$
$\mathrm{J}_{2}=1.082645 \times 10^{-3}$
$\left(\mathrm{m}_{\mathrm{e}} / 2 \mathrm{n}_{\mathrm{E}}\right)=3.334320 \times 10^{6}$

Table 8.4
Comparison of Woolard's Euler Angle Perturbations with those from Table 8.3

| Code Numbers of Symmetric Tidal Arguments | Terrestrial z Axis Euler Angle Perturbations from Table 8.3 $\delta \theta_{z} \quad \delta \psi_{z} \sin \theta$ |  | Euler Angle Perturbations from Woolard's Equation (54) $\delta \theta_{z} \quad \delta \psi_{z} \sin \theta$ |  |
| :---: | :---: | :---: | :---: | :---: |
| 117.655 | $00^{\prime} .00004547$ | 0.00004547 |  | 0':00005 |
| $125.755 \quad 1 \times 5.355$ | 0.00014949 | 0.00016291 | 0.00015 | 0.00016 |
| $127.5451 \times 3.565$ | 0.00003042 | 0.00004089 |  | 0.00004 |
| $127.5551 \times 3.555$ | 0.00018041 | 0.00019676 | 0.00018 | 0.00020 |
| 135.645195 .465 | 0.00018989 | 0.00025499 | 0.00019 | 0.00025 |
| 135.655195 .455 | 0.00112938 | 0.00123111 | 0.00113 | 0.00123 |
| 137.445193 .665 | 0.00003598 | 0.00004841 |  | 0.00005 |
| 137.455193 .655 | 0.00021459 | 0.00023389 | 0.00021 | 0.00023 |
| 145.545185 .565 | 0.00099215 | 0.00133203 | 0.00133 | 0.00099 |
| 145.555185 .555 | 0.00589893 | 0.00642984 | 0.00590 | 0.00643 |
| 147.555183 .555 | -0.00016078 | 0.00000016 | -0.00016 |  |
| 153.655177 .455 | -0.00004351 | -0.00004743 |  | -0.00005 |
| 155.455175 .655 | -0.00016667 | -0.00018171 | -0.00017 | -0.00018 |
| $155.645 \quad 175.465$ | -0.00008211 | 0.00010991 |  | -0.00011 |
| 155.655175 .455 | -0.00096958 | 0.0 | -0.00097 |  |
| 155.665175 .445 | -0.00008292 | -0.00011138 |  | -0.00011 |
| 157.455173 .655 | -0.00018515 | 0.0 | -0.00018 |  |
| 162.556168 .554 | 0.00016140 | 0.00017588 | 0.00016 | 0.00018 |
| 163.555167 .555 | 0.00275303 | 0.00300082 | 0.00275 | 0.00300 |
| 164.556166 .554 | -0.00013869 | 0.0 | -0.00014 |  |
| 165.545165 .565 | -0.00100294 | 0.00134642 | -0.00100 | -0.00135 |
| 165.555 | -0.00868237 | 0.0 | -0.00868 |  |

## 9. COMPUTATIONAL FORMULAS AND APPLICATIONS

For numerical computations the diurnal terms in Equations (7.18), (7.29), (7.31) and (7.39) to (7.44) for $\mathrm{m}, \psi_{\mathrm{f}}, \mathrm{H} / \mathrm{C} \Omega$, and the Euler angle perturbations are written in terms of the common multipliers of the lunar and solar terms that arise in the theory of precession and nutation [Woolard, 1953, pp. 124-125]. The common multiplier of the solar terms is

$$
\begin{align*}
K_{\odot} & =-3\left(\frac{G m_{\odot}}{c_{\odot}^{3}}\right)\left(\frac{C-A}{C \Omega}\right)  \tag{9.1}\\
& =-3484!15(\text { Julian century })^{-1}
\end{align*}
$$

and the common multiplier of the lunar terms is

$$
\begin{align*}
K_{s} & =-3\left(\frac{G m_{\odot}}{a_{E}^{3}}\right)\left(\frac{C-A}{C \Omega}\right)\left(\frac{m_{9}}{m_{\odot}}\right)\left(\frac{a_{E}}{c_{3}}\right)^{3}  \tag{9.2}\\
& =-7552.4295(\text { Julian century })^{-1}
\end{align*}
$$

From Equation (4.24),

$$
\begin{equation*}
\left(\frac{A_{j}}{A \Omega^{2}}\right)_{d}=-3 \frac{G m_{d}}{\Omega^{2} c_{d}^{3}}\left(\frac{C-A}{A}\right) A_{21 j} \tag{9.3}
\end{equation*}
$$

which is written in terms of the constants $K_{\odot}$ and $K_{\rho}$ as

$$
\begin{equation*}
\left(\frac{A_{j}}{A \Omega^{2}}\right)_{d}=K_{d} \frac{C}{A \Omega} A_{21 j} \tag{9.4}
\end{equation*}
$$

In order to make the value of the secular Love number $\mathrm{k}_{\mathrm{s}}$, given by (5.17), consistent with the values (9.1) and (9.2) of $K_{\odot}$ and $K_{p}, k_{s}$ is written as

$$
\begin{equation*}
k_{s}--\left(\frac{K_{\odot}}{\Omega}\right)\left(\frac{c_{\odot}}{a_{E}}\right)^{3}\left(\frac{C-A}{m_{E} a_{E}^{2}}\right)\left(\frac{C}{C-A}\right)\left(\frac{m_{E}}{m_{\odot}}\right) \tag{9.5}
\end{equation*}
$$

The computational form of Equation (7.18) for the polar motion is

$$
\begin{equation*}
m=m_{0} e^{i \sigma_{0} t}+\frac{\left(\frac{c_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)}+\sum_{j} \frac{i\left(\frac{A_{j}}{A \Omega^{2}}\right) e^{-i\left(\omega_{j} t+\beta_{j}\right)}}{\left[\frac{\left(\frac{\omega_{j}}{\Omega}\right)}{\left(1-\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)}+\frac{\sigma_{r}}{\Omega}\right]} \tag{9.6}
\end{equation*}
$$

where ( $A_{j} / A \Omega^{2}$ ) and $k_{s}$ are computed from (9.4) and (9.5) respectively. The direction cosines of the axis of figure are computed using Equation (7.29) in the form

$$
\begin{align*}
& \psi_{f}=\frac{k}{k_{s}} m_{0} e^{i \sigma_{0} t}+\frac{\left(\frac{c_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)} \\
& -\sum_{j} \frac{i\left(\frac{A_{j}}{A \Omega^{2}}\right)\left(\frac{\Omega}{\sigma_{r}}\right)\left(\frac{k}{k_{s}}\right)\left(\frac{\omega_{j}}{\Omega}\right) e^{-i\left(\omega_{j} t+\beta_{j}\right)}}{\left(1-\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)\left[\frac{\left(\frac{\omega_{j}}{\Omega}\right)}{\left(1-\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)}+\frac{\sigma_{r}}{\Omega}\right]} \tag{9.7}
\end{align*}
$$

The computational form of Equation (7.31) for the angular momentum vector is

$$
\begin{align*}
\frac{H}{C \Omega}= & {\left[\frac{A}{C}+\left(\frac{C-A}{C}\right) \frac{k}{k_{s}}\right] m_{0} e^{i \sigma_{0} t}+\frac{\left(\frac{C_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)} } \\
& +\sum_{j} \frac{i\left(\frac{A_{j}}{A \Omega^{2}}\right)\left(\frac{A}{C}\right)\left(1-\frac{k}{k_{s}}\right) e^{-i\left(\omega_{j} t+\beta_{j}\right)}}{\left(1-\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)\left[\frac{\left(\frac{\omega_{j}}{\Omega}\right)}{\left(1-\frac{n_{j}}{\Omega} \frac{k}{k_{s}}\right)}+\frac{\sigma_{r}}{\Omega}\right]} \tag{9.8}
\end{align*}
$$

The computational forms of Equations (7.39) to (7.44) for the Euler angle perturbations are obtained similarly in terms of the dimensionless ratio ( $\mathrm{A}_{\mathrm{j}} / \mathrm{A} \Omega^{2}$ ).

Equations (9.6), (9.7) and (9.8) are used to compute the coefficients listed in Table 9.1 for the diurnal motion of $m, \psi_{f}$, and $\mathrm{H} / \mathrm{C} \Omega$ in a rigid Earth. Table 9.2 gives coefficients for evaluating the diurnal motion of $m, \psi_{f}$, and $\mathrm{H} / \mathrm{C} \Omega$ in a deformable Earth with $\mathrm{k}=0.29$. The computational forms of Equations (7.39) to (7.44) are used to obtain the coefficients given in Tables 9.3 and 9.4 for the Euler angle perturbations. Table 9.3 is for a rigid Earth and Table 9.4 is for a deformable Earth with $\mathrm{k}=0.29$.

The effect of polar motion on latitude and time is shown in Figure 9.1. The true equinox of date $\tau_{T}$ corresponds to the ascending node of the true equator of date on the mean ecliptic of date. The true sidereal system has its x axis directed toward $\boldsymbol{r}_{\mathrm{T}}$, its y axis $90^{\circ}$ eastward from $\boldsymbol{\gamma}_{\mathrm{T}}$ in the true equator of date and its z axis along the rotation axis $\vec{\omega}$. The transformation from the true sidereal system into the terrestrial or "observatory-fixed" system $x, y, z$ is given by

$$
\begin{equation*}
\vec{x}_{\text {terrestrial }}=R_{2}\left(-m_{1}\right) R_{1}\left(m_{2}\right) R_{3} \text { (GASTI) } \vec{x}_{\text {true siderial }} \tag{9.9}
\end{equation*}
$$

where $R_{j}(\alpha)$ denotes the rotation of a coordinate system about its $j$ axis through the angle $\alpha$, and GAST1 is the Greenwich apparent sidereal time corrected for polar motion. Let $\mathrm{x}_{\mathrm{T}}, \mathrm{y}_{\mathrm{T}}$, and $\mathrm{z}_{\mathrm{T}}$ correspond to the system obtained by rotating the true sidereal system through GAST1 about $\vec{\omega}$. Then

$$
\begin{equation*}
\vec{x}_{\text {terrestrial }}=R_{2}\left(-m_{1}\right) R_{1}\left(m_{2}\right) \vec{x}_{T} \tag{9.10}
\end{equation*}
$$

Latitude and longitude are denoted respectively by $\Phi_{T}$ and $\Lambda_{T}$ in the $x_{T}, y_{T}, z_{T}$ system and by $\Phi$ a $\Lambda$ in the terrestrial system. Equation (9.10) is used to obtain

$$
\begin{align*}
& \cos \Phi_{T} \cos \Lambda_{T}=\cos \Phi \cos \Lambda-m_{1} \sin \Phi  \tag{9.11}\\
& \cos \Phi_{T} \sin \Lambda_{T}=\cos \Phi \sin \Lambda-m_{2} \sin \Phi  \tag{9.12}\\
& \sin \Phi=\sin \Phi+\cos \Phi\left(m_{1} \cos \Lambda+m_{2} \sin \Lambda\right) \tag{9.13}
\end{align*}
$$

in which second and higher order terms in $m_{1}$ and $m_{2}$ are neglected. From (9.13), the first order expression for the latitude variation is

$$
\begin{equation*}
\Phi_{T}-\Phi=m_{1} \cos \Lambda+m_{2} \sin \Lambda \tag{9.14}
\end{equation*}
$$

Equations (9.11) and (9.12) are combined to give

$$
\begin{equation*}
\Lambda_{T}-\Lambda=\tan \Phi\left(m_{1} \sin \Lambda-m_{2} \cos \Lambda\right) \tag{9.15}
\end{equation*}
$$

# Coefficients for the Diurnal Motion of the Rotation Axis, the Axis of Figure and the Angular Momentum Vector in a Rigid Earth 



Table 9-1-(continued)

COFFFICIFNTS IM SFCOMAS OF ARC
SINE FOR $X$-CNMPONFNTS
COSINE FOR Y-CTMPINFNTS


## Table 9-1-(continued)

|  |  |  |  |  |  |  |  |  | COFFFICIENTS IN SFCIMDS MF <br> SINE FOR $X$-CTMPINFNTS <br> COSINE FDR Y-COMPONENTS |  |  | $\triangle R C$ |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
| INDEX. | TIDAL ARG. CDDE NUMBER | DIST. <br> RחDY | COEFFICIENTS |  |  |  | [ F | $n \mathrm{n}$ | $\begin{gathered} \text { ROTATION } \\ \text { AXIS } \end{gathered}$ | $\Delta X I S$ TF FIfita F | ANGULAR MIIMFNTIIR |  |
| 133 | $1 \times 5.355$ | M | 1 | 2 | 0 | 2 | 0 | 2 | 0.0000059 | 0.0 |  | .0000059 |
| 134 | $1 \times 5.365$ | M | 1 | 2 | 0 | 2 | 0 | 1 | 0.0000029 | 0.0 |  | . 00000078 |
| 135 | 1E3.455 | 1 | 1 | 1 | 0 | 2 | 2 | 2 | 0.0000017 | 0.0 |  | . 0000017 |

## Explanation of symbols

PHIM ts the Greenwich mean sidertal time; $\phi$,
L, LP, F, D and OM are Brown's fundamental
arguments; $l, l^{\prime}, F, D$ and $\Omega$

```
Constanta
            k=0
    K, =-7!552430 \times 10 3
    K}=-3!484150 \times 10 3 Jultan century -1
    (C-A)/C = 3.272930 }\times1\mp@subsup{0}{0}{-3
        n=380:9856 duy }\mp@subsup{}{}{-1
    (C6}/\mp@subsup{C}{2}{\prime})=2.343852\times1\mp@subsup{0}{}{4
        J
    (mo/m
```

Table 9-2

## Coefficients for the Diurnal Motion of the Rotation Axis, the Axis of Figure and the Angular Momentum Vector in a Deformable Earth

| INDEX | TIDAL akg. CCDE NUABER | $\begin{aligned} & \text { DIST. } \\ & \text { BNDY } \end{aligned}$ | PHiM | COEFFICIENTS |  |  | $\begin{aligned} & {[1 F} \\ & \mathrm{D} \end{aligned}$ | 0 Mi | CDFFFICIFNTS IN SFGONDS MF ARC SINF FIIR $x$-CMMPINAFNTS COS IME FIR $Y$-GTIMPGMFNTS |  |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  |  |  |  |  |  |  |  |  |  |
|  |  |  |  | 4 | LP | F |  |  | AXIS | AXIS $\cap F$ <br> FIGIIRF | ANGIIAR mithanation |
| 1 | 105.955 | M | 1 | -4 | 0 | -2 |  |  |  |  |  |  |
| 2 | 107.755 | M | 1 | -2 | 0 | -2 | -2 | -2 | -0.0000022 | 0.0001699 | -0.0000016 |
|  | 109.555 | M | 1 | 0 | 0 | -2 | -4 | -2 | -0.0000054 | 0.00071105 | -0.00000th |
| 4 | 115.845 | M | 1 | -3 | 0 | -2 | 0 | -1 | -0.0000054 | 0.0004325 | -0.0000040 |
| 5 | 115.455 | M | 1 | -3 | 0 | -2 | () | -2 | -0.0000 204 | 0.0003744 | - 1.00000079 |
| 6 | 117.645 | M | 1 | $-1$ | 0 | -2 | -2 | -2 | -0.0000204 -0.0000100 | 0.0016684 | -0.0000149 |
| 7 | 117.655 | $\cdots$ | 1 | -1 | 0 | -2 | -2 | -1 -2 | -0.0000100 | 0.0008184 | -0.00000073 |
| 8 | 118.654 | M | 1 | -1 | 1 | -2 | -2 | - | -0.0000524 | 0.00042946 | -0.0000381 |
| 9 | 119.445 | M | 1 | - | 0 | -2 | -2 | -1 | -0.0000079 -0.0000019 | 10.0003744 0.0001545 | -0.000000? |
| 10 | 119.455 | M | 1 | 1 | 0 | -2. | -4 | -2 | -0.00000101 | 0.0001545 | $-0.0000014$ |
| 11 | 124.756 | M |  | -2 | -1 | -2 | -4 | -2 | -0.0000101 0.0000024 | 1.00018242 -0.0002008 | $-0.00140074$ |
| 12 | 125.745 | M | $i$ | -2 | 0 | -2 | 0 | -1 | - 0.0000330 | $\begin{array}{r} -0.0002008 \\ 0.0027810 \end{array}$ | $\begin{gathered} 0.0000017 \\ -0.0000>39 \end{gathered}$ |
| 13 | 125.755 | M | 1 | -2 | 0 | -2 | 0 | -2 | -0.0001751 |  | $\text { -0. } 0000739$ |
| 14 | 126.556 | M | 1 | 0 | -1 | -2 | -2 | -2 | -0.0001751 0.0000079 | 0.0147545 -0.0012472 | $-0.0(1) 1262$ $0.0 n 00031$ |
| 15 | 126.655 126.754 | M in | 1 | -1 -2 | 0 | -2 | -1 | -2 | 0.0000020 | -0.00002472 -0.0001599 | 0.04100071 0.0000014 |
| 17 | 127.545 | M | 1 | -2 | 1 | -2 | 0 | -2 | - $0.000 \mathrm{coz7}$ | 0.10102717 | - 0.00000000 |
| 14 | 127.555 | M | 1 | 0 | 0 | -2 | -2 | - | -0.000030y | 0.0033681 | -0.000028*. |
| 19 | 128.544 | M | 1 | 0 | 1 | -2 | -2 | -2 | -0.000) -0.00105 | $0.017 \times 1.39$ | -0.00015515 |
| 20 | 128.550 | M | 1 | 0 | 5 | -6 | 2 | - -1 | - -0.0000026 | 0.0002163 | -0.000vola |
| 21 | 129.355 | M | , | 2 | 0 | -2 | -4 | -2 | -0.0000144 | 0.0012706 | -0. nu00] 113 |
| 22 | 133.855 | , | 1 | -3 | 0 | -2 | -4 | -2 | -0.0000064 | 0.0005408 | -0.0000046 |
| 23 | 134.656 | M | 1 | -1 | -1 | -2 | 0 | -2 | 0.0000041 | -0.01003554 | 6. 0 Outon>o |
| 24 | 1.35.435 | M | 1 | 1 | 0 | -2 | 0 | -2 | 0.0000109 0.000050 | -0.00094? | $0.0060 n 78$ |
| 25 | 135.635 | 1. | 1 | -1 | 0 | -2 | 0 | -2 | 0.0000050 | -C.0004326 | 0.nuthonza |
| 2 h | 135.645 | M | 1 | -1 | 0 | - | , | -1 | -0.0000075 | -0.0006490 | 0.00001053 |
| 27 | 135.655 | M | 1 | -1 | 0 | -2 | 0 | -1 | -0.0002470 -0.0012840 | \$.0210140 | -0.0001729 |
| 2 R | 135.855 | M | 1 | -3 | 0 | 0 | 0 | - | -0.0012840 | 1.1114976 | -0.0009149 |
| 29 | 136.456 | M | 1 | 1 | -1 | -2 | -2 | -2 | n.0nono 04 | -0.010012926 | 0.0001)(02\% |
| 30 | 136.555 | M | 1 | 0 | 0 | -2 | -1 | -2 | 0.0000023 | -0.000200\% | 0.10 mog 1 c |
| 31 | 136.644 | M | - | -1 | 1 | -2 | -1 | - | -0.00000 29 | -0.000n076 | 0.01000049 |
| 32 | 136.554 | M | 1. | -1 | 1 | - ? | $1)$ | - | - -0.00001 ? | 0.0001700 | -0.0600n14 |
| 33 | 137.445 | M | 1 | - | 0 | -? | -2 | -2 | -0.0000nl ${ }^{-0.000047}$ | 0.00105017 | -0.0mmonra |
| 34 | 137.455 | M | 1 | 1 | 0 | -2 | -2 | -1 | -0.0000457 | 0.0039865 | -0.0000325 |
| 35 | 137.655 | M | 1 | -1 | 0 | -2 | -2 | -2 | -0.0002430 | 0.0211842 | - 0.00001720 |
| 36 | 137.665 | M | 1 | $-1$ | 0 | 0 | -2 | -1 | O.000n134 | -0.01012052 | 0.0000098 |
| 37 | 138.444 | N | 1 | 1 | 1 | -2 | -2 | -1 | -0.0000043 | 0.0003708 | -0.000003n |
| 38 | 138.454 | M | 1 | 1 | 1 | -2 | -2 | -1 | -0.00000) 9 | 0.0001700 | -0.00กtyol4 |
| 39411 | 139.455 | M | 1 | 1 | 0 | -2 | -2 | -2 | -0.0000 113 | 0.0009 Rrg | -0.0noonan |
|  | 143.535 | M | 1 | 11 | 0 | -4 | -4 | - | 0.0000025 0.0000230 | -0.0002163 | 0.0000018 |
| 41 | 143.745 | M | 1 | -2 | 0 | -2 |  |  | -1000 | -0.0002627 | $0.006 m 021$ |
| 42 | 143.755 | M | 1 | -2 |  | -2 |  | -1 | 0.0000035 | -0.0003091 | 0.0000025 |
| 43 | 14.4 .546 | M | 1 | -2 | -1 | -2 | ${ }^{2}$ | -2 | 0.0000196 | -0.0017462 | 0.0100138 |
| 44 | 144.556 | M | 1 | 0 | -1 | -2 -2 | 0 | -1 | 0.0000026 | -0.0002318 | U.0000018 |
| 45 | 145.535 | M | 1 | 0 | - | -2 | 0 | -2 | 0.0000275 | -0.002u0rg | 0.0000159 |
| 46 | 145.545 | M | 1 | 0 | 0 | -2 | O | -1 | -0.0000377 | -0.0033687 | 0.0000276 |
| 47 | 145.555 | M | 1 | 0 | 0 | -2 | 0 | -1 | -0.0012292 | 0.1097934 | -0.0008t50 |
| 48 | 145.755 | M | 1 | -2 | 0 | . | 0 | -2 | -0.0065199 | 0.5824076 | -0, 01145924 |
| 49 | 145.765 | M | 1 | -2 | 0 | 0 | 0 | -1 | 0.0000420 | -0.0037551 | 1. nonospoh |
| 50 | 146.544 | M | 1 | 0 | 1 | -2 |  | -1 | 0.0000069 | -0.0006191 | 0.0000049 |
| 51 | 146.554 | H | . | 0 | 1 | -2 | 0 | -1 | -0.00000 -0.0000 | 0.0001854 | -0.00000015 |
| 52 | 147.355 | M | 1 | 2 | 0 | -2 | -2 | -2 | -0.0000199 | 0.0017771 | -0.0000140 |
| 53 | 147.545 | M | 1 | 2 | U | -2 | -2 | -2 | O.0n00076 | -0.0003245 | 0.0 (inoinc 5 |
| 54 | 147.555 | M | 1 | , | 0 |  | -2 | , | -0.0000024 | 0.0002163 | - -0.1000017 |
| 55 | 147.565 | N | 1 | 0 | 0 | 0 | -2 | -1 | O.0nonrat | -0.0075975 | 0.10000595 |
| 56 | 148.554 | M | 1 | 0 | 1 | 0 | -2 | $-1$ | -0.0000184 | 0.0016535 | -0.0000130 |
| 57 | 152.656 | M | 1 | -1 | -1 | -2 | -2 | - | n.000n057 | -0.0005 100 | 0.00000480 |
| 58 | 153.645 | M | 1 | -1 | -1 | -2 | 2 | -2 | 0.0000024 | -0.0002144 | 0.0000017 |
| 59 | 153.655 | M | 1 | -1 | 0 | -2 | 2 | -1 | 0.0000186 | -0.0009736 | 0.0500074 |
| 60 | 154.656 | M | 1 | - | -1 | - | 2 | -2 | -0.0000470 | -0.0042963 | 0.0000328 |
| 61. | 155.435 | M | 1 | -1 | -1 | -2 | 0 | 0 | -0.000n025 | 0.0002718 | -0.0000019 |
| 62 | 155.445 | M | 1 | 1 | 0 | -2 | 0 | 0 | -0.0000029 | 0.0002627 | -0.0000020 |
| 63 | 155.455 | M | 1 | 1 | 0 | -2 | 0 | -1 | 0.01100372 | -0.00311445 | 0.0000231 |
|  |  |  |  |  |  |  | 0 | -2 | 0.0001794 | -0.0164599 | 0.0001249 |

Table 9-2-(continued)
COFFFICIENTS IA SFCONDS OF ARC
SINF FOR X-C OMPIINFNTS
COSINE FQR Y-COMPONENTS


## Table 9-2-(continued)

```
COEFFICIFMTS IN SFCONTS GF ARC SINE FIR \(X\)-COMPINMFNTS COSINF FGR Y-COMPINENTS
```

| INDEX | TIDAL ARG. CODE NUMBER | $\begin{aligned} & \text { OIST. } \\ & \text { BODY } \end{aligned}$ | PHIM | COEF | $\begin{gathered} \text { FICI } \end{gathered}$ | $\begin{gathered} \text { ENTS } \\ \text { F } \end{gathered}$ | $\mathrm{OF}$ | nm | $\begin{aligned} & \text { ROTATION } \\ & \text { AXIS } \end{aligned}$ | AXIS OF Figlirf | ANGIJLAR <br>  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
| 132 | $1 \times 3.565$ | M | 1 | 0 | 0 | 2 | 2 | 1 | $0.00000 \angle 4$ | -0.0004947 | 0.0000032 |
| 133 | $1 \times 5.355$ | M | 1 | 2 | 0 | 2 | 0 | 2 | 0.0000081 | -0.0006339 | 0.0000040 |
| 134 | $1 \times 5.365$ | M | 1 | 2 | 0 | 2 | 0 | 1 | 0.0000040 | -0.00014 174 | 0.9000027 |
| 135 | 1153.455 | M | 1 | 1 | 0 | 2 | 2 | 2 | 0.0000018 | -0.0001855 | O.0000012 |

Explanation of symbols
PFIM is the Greenwich mean sidertal time; 中
L, LP, F, D and OM are Brown's fundamental arguments; $\ell, \ell^{\prime}, \mathrm{F}, \mathrm{D}$ and $\Omega$

Constants
$k=0$
$K_{7}=-7!552430 \times 10^{3}$ Jultan century ${ }^{-1}$
$K_{a}=-315484150 \times 10^{3}$ Julian century ${ }^{-1}$
$(\mathrm{C}-\mathrm{A}) / \mathrm{C}=3.272930 \times 10^{-3}$
$\mathrm{n}=36009856 \mathrm{day}^{-1}$
$\left(\mathrm{C}_{5} / \mathrm{a}_{\mathrm{E}}\right)=2.343852 \times 10^{4}$
$J_{2}=1.082645 \times 10^{-3}$
$\left(m_{\sigma} / m_{R}\right)=3.334320 \times 10^{6}$

## Table 9－3

# Coefficients for Perturbations in the Euler Angles of the Rotation Axis，the Axis of Figure and the Angular Momentum Vector in a Rigid Earth 



C．asinks fig nftta thfta
SINFS FOR NFLTA HST：SIMITHFTAS

| tionl CIOE | ARGHMENT NUMBERS | COEFFICIFATS |  |  |  | OF | $\begin{gathered} \text { ent } \triangle T I B N \\ \text { AXIS } \end{gathered}$ |  | $\Delta \times I S$ nF FITHIPF |  | $\begin{gathered} \text { TFKKFSTKIAL } \\ z \text { nxlS } \end{gathered}$ |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  | L | LP | F | 1） | UP1 | crss | SIN | cns | STM | Lns | $\bigcirc \mathrm{S}$ |
| 105.955 | － | 4 | 0 | 2 | 0 | 2 | －0．00000601 | －0．0nounome | 0．0nuonez30 | 0.00 0maフa | 9．（1）notor30 |  |
| 1117.755 |  | 2 | 0 | 2 | 2 | 2 | －0．nonunnoz | －0．01000003 | $0.0 n u 00955$ | O． 0 M00n0 55 | 11.006010955 |  |
| 109.555 |  | 0 | 0 | 2 | 4 | 2 | －c．oonunam？ |  | 0． 10 numas 78 | 1）．00000n57a |  |  |
| 115.845 |  | 3 | 0 | 2 | 0 | 1 | －0．no00000］ | －0．0160000 | 1）． 0 nu00419 |  | 1．060606414 | $0.000004{ }^{\text {a }}$ |
| 115.855 |  | 3 | 0 | 2 | 0 | 2 | － 1.000 momoz | －0．0mbornori | 0.00 0n2157 | O．00naplat | （1．160）（18） 157 | H．（rimioltio |
| 117.645 | － | 1 | 0 | 2 | 2 | 1 | － 2 －nunonma | － 0.004000078 | n．nouninsz | 0．0000105？ | 1）．0000105\％ |  |
| 117.655 | － | 1 | 0 | 2 | 2 | 2 | －0．00000n 18 | －0．00sunfila | 0.001805519 | n．0ncm5519 | 0.000 ¢5bis |  |
| 118.654 |  | 1 | －1 | 2 | 2 | 2 | －0．Dunuormi | －0．00000001 | O．nomonemin | 6． 0 nonoti4 1 A | 6．choteri4 4 |  |
| 119.445 |  | －1 | 0 | 2 | 4 | 1 | －0．000uname |  | $0.000 n \mathrm{n} 197$ | 0.0 mon 197. | a．OrMentic ${ }^{\text {a }}$ |  |
| 119.455 |  | －1 | 0 | 2 | 4 | 2 | －0．00000003 | －6．0rvonooz | O．M（0） | 0．nomoniota |  | 0.089001 arar |
| 124.756 |  | $?$ | 1 | 2 | 0 | 2 | 0.000 onom | 0.00 cosom | －0．anonar 49 | －0．0nomri4． |  | －9）． 9 Mothne 44 |
| 125.745 | $1 \times 5.365$ | 2 | 0 | 2 | 0 | ， | －0．00000010 | －0．00000013 | 0.000033588 | 0.00003829 | 0.0000305 | （9．006036ス\％ |
| 125.755 | $1 \times 5.355$ | 2 | 0 | 2 | $\checkmark$ | 2 |  | －－omounataz | 0.001017 ari | O． 010 mbraza | 0.001017 mgl |  |
| 126．55 | － | 0 | 1 | 2 | 2 | 2 | O．bununnol | －．ombuntios | －0．00u00305 | －0．000013015 | －11． $0 \times 1040305$ |  |
| 126.655 | － | 1 | 0 | 2 | 1 | 2 | $0.0000000]$. | O．Onwonctul | －0．00060） 10 | －0．0nmonilo | －0．00000＜ $0^{10}$ |  |
| 126.754 | － | 2 | －1 | 2 | 0 |  | －- ． 0 00unnos | －0．006000001 | $0.00 y 002 \mathrm{R}$ | 0．octunozat |  |  |
| 127.545 | $1 \times 3.565$ | 1 | 0 | 2 | 2 | 1 | －0．00numal？ | －0．00cuonels | 0．0n0nzear | 0．nombicarat | 0.000112884 |  |
| 127.555 | $1 \times 3.555$ | 0 | 0 | 2 | 2 | 2 | －0．00000070 | －0．000000174 | 0.00071213 | 0.00072448 | $0.00421 / 13$ | O．060\％2 atre |
| 128.544 | － | 0 | －1 | 2 | 2 | ， | －0．00000nol | － 0 ．oriounciol | 0．anumatis | O．nomoobeas |  | O．10（10）3／（\％） |
| 128.550 | － | 0 | －5 | 6 | －2 | 6 | －－coununnos | －0．0r0enoue 5 | W．OnuO1498 | 0．000014498 | 1． 00060.498 |  |
| 129.355 | ． | －2． | 0 | 2 | 4 | 2 | －0．00nouno？ | －0．00060）cor | 0.0000016 h ？ | 0.0 nomornat | 0.00 Shliteraz | 13．6numbote |
| 133.855 |  | 3 | 0 | 2 | －2 |  | 0．nunuonol | O．0rnomol | －0．0000n424 | －0．00non424 | －0．00000424 | －0．uncmo4\％ 4 |
| 134.656 |  | 1. |  | 2 | 0 |  | 0.00000004 | 0.000000004 | －0．0nonl123 | －0．00011．73 | －0．00061123 | －0．0104以」ンタ |
| 135.435 |  | －1 | $1)$ | 4 | 0 | 2 | O．nonuono？ | 0．0rimatiou？ | －0．0nunn514 | －0．00nnot5l4 | －0．000005 54 | －6．06） 6 （1）${ }^{\text {a }}$ |
| 135.035 | 195.475 | 1 | 0 | 2 | 0 | 0 | O． 600000 n 5 | $0.00 n 00000$ | －0．00001290 | －0．00nocil 5 ？ | －1）．0n6itiz90 |  |
| 135.645 | 195.465 | 1 | 0 | 2 | 0 | 1 | －0．nomonnt？ | －0．0nvonory？ | 0.00022030 | 0.011027899 | （1．040202030 | （1．60）127418 |
| 135.455 | 195.455 | 1 | 0 | 2 | 0 |  | －0．00000420 | －0．06000450 | 0．00127850 | 0． 1101371924 | 1.00127851 |  |
| 135.855 | 195.255 | 3 | 0 | 0 | 0 |  | 0.10000002 | a．0nouomua | －0．0nutinfiza | － 0 ． n （mmonkr | －n． 010000029 | －0．0rmatuais |
| 136.456 | ， | $-1$ |  | 2 | ， | 2 | 0.00000001 | 0.00000001 | －0．nnunn） 38 | －0．00）noonzas | －0．00000238 |  |
| 136.555 |  |  | 0 | 2 | 1 | 2 | 0．oumunno？ | 0.00010 mog 2 | －0．00000714 | －0．00000714 | －0．00u007 714 |  |
| 136.644 | ． | 1 | －1 | 2 | 0 | ， | －0．0000n001 | －0．00000001 | 0.00000201 | $0.00000 \mathrm{Pa!}$ | 0．000000201 |  |
| 136.654 | － | 1 | －1 | 2 | 0 | 2 | －1）． 00000 nm 4 | －0．00000004 | 0.00001744 | 0.00001244 | 1． 1.00001244 | 0.0 0001844 |
| 137.445 | 193．665 | －1 | 0 | 2 | 2 | 1 | －0．000000 14 | －0．0n00tmi．7 | 0.00004148 | 0.000057 .77 | 0.010004145 | 0.0 unuese 73 |
| 137.455 | 193．655 | －1 | 0 | 2 | 2 | 2 | －6．nonuonto |  | $0.00 \cup 24153$ | 0．01025 5901 | （1．00024153 | 0.000025401 |
| 137.655 | 193.455 | ， | 0 | 0 | 2 | 0 | 0．0unonoer | 0.00000001 | －0．00007579 | －0．00000267 | －0．00002579 |  |
| 137.665 | ． | 1 | 0 | 0 | 2 | 1 | － 1.00000001 | －0．00000001 | 0.00000438 | 0.00000438 | 0.000006438 | 0.000000433 － |
| 138.444 | － | －1 | －1 | 2 | 2 | 1 | －0．00000n01 | －0．00000001 | 0.00000200 | 0.00000200 | 0.000002800 |  |
| 138.454 | － | －1 | －1 | 2 | 2 | 2 | －0．00000004 | －0．0n00000） | 0.00001 .165 | 0.0000114 .5 | 0.00601165 | 0.000001165 |
| 139.455 | 191．655 | －1 | 0 | 0 | 4 | 0 | 0.00000002 | 0.00000000 | －0．nnuo0477 | －1）．00060031 | －0．100000477 | －0．unorousi |
| 143.535 | － | 0 | 0 | 4 | －2 | 2 | 0.00000001 | 0.0 mounool | －0．0n000302 | －0．000003n？ | －0．001000302 | －0，uncoro302 |

Table 9－3－（continued）

> COEFFICIFMTS FOR EIILER AMGLE PERTURRATIONS IN SFGNNS OF ARC

C．IStNFS FOR melta thfta
INFS FOR DELTA PSI $\Rightarrow$ SIN（THETA）

| $\begin{array}{r} \text { TIOAL } \\ \text { CODE } \end{array}$ | ARGUMENT i．UMBBERS | COEFFICIENTS |  |  |  | －F | $\begin{gathered} \text { ROTATIOM } \\ \text { AXIS } \end{gathered}$ |  | AXIS AF <br> EIGIRF |  | TEKRESTRIAL $Z$ axis |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  | L． | LP | F | 1 | GM | $\cos$ | SIM | cos | $51^{\text {n }}$ | cuis | SIM |
| $\bigcirc 143.745$ |  | 2 | 0 | 2 | －2 | 1 | 0.00000001 | c． 000000001 | － 0.00000735 | －0．00000355 | －0．000001355 | －0．0n0u0355 |
| 143.755 |  | 2 | 0 | 2 | －2 | 2 | 0.00000007 | 0.061000007 | －0．00002004 | － 0.00002004 | －0．00002004 | －1． 0 00022004 |
| 144.546 |  | 0 | 1 | 2 | 0 |  | 6．nonoono 1. | 0.000000081. | －0．00unn ${ }^{\text {a }}$ 5 | －0．00000？ 65 | －0．001010 ${ }^{\text {a }}$－ 5 |  |
| 144.556 | － | 0 | 1 | 2 | 0 | 2 | 0.00000008 | 0.04000048 | －0．0n002300 | －0． 0000 man | －0．000023001 | －0．00002310 |
| 145.535 | 185.575 | 0 | 0 | 2 | 0 | 1 | 0．bununnza | 0.0000000 ？ | －0．0nun7169 | －0．000005 ${ }^{0} 4$ | －0．00007164 | －19．0000458． 4 |
| 145.545 | 185.565 | 0 | 0 | 2 | 0 | 1 | －0．0000036 | －0．00000464 | 0.00109508 | 0.00141185 | 0.00109509 | 0.00141127 |
| 145.555 | 185.555 | n | 0 | 2 | 0 | 2 | －0．00002102 | －0．000002264 | $0.0 n g<$ nobn | 0.0 ngras50 | 1．008440068 | O．uncrasis |
| 145.755 | 185.355 | 2 | 0 | 0 | 0 | 1 | ¢．00000026 | $0.0 n 000002$ | －0．00007945 | －0．0000042？ | －0．00007945 | －0．0n000nc2 |
| 145.765 |  | 2 | 0 | 0 | 1） | 1 | 0.00000002 | 6． －inguonmo 2 | －6．nnu00705 | －0．000000705 | －6．000000705 | －0．096кく\％705 |
| 146.544 |  | 0 | －1 | 2 | 0 | 1 | －0．00000001 | －0．00000001 | 0.000002711 | 0．0000n？${ }^{\text {a }}$ | U．000002211 | 1．Unomatil |
| 146.554 |  | 0 | －1 | ？ | 0 | 2 | －0．00000007 | －0．0noonmo 7 | 0．0000） 0 23 | 0.00002023 | 1．00002023 | 0．0000以2033 |
| 147.355 |  | －2． | ， | 2 | 2 | 2 | 0.00000001 | $0.0 n 00 \mathrm{coml}$ | － $0.0 n 000$ スe9 | －0．0ヶ00n3ka | －0．00000369 | －0．0000033 ${ }^{\text {a }}$ |
| 147.545 | 183.565 | 0 | 0 | 0 | 2 | －1 | 0.00000004 | －0．00000006 | －0．00001725 | 0.60001714 | －0．006m1225 | G．000m171t |
| 147.555 | 183.555 | 0 | 0 | 0 | 2 | 0 | 0.00000053 | 0.00000004 | －0．00016150 | －0．00001071 | －0．000） 0150 | －0．00001uld |
| 147.565 | 183.545 | 0 | 0 | 0 | 2 | 1 | －0． 0.0000005 | －0．0000000 | 0．0noolabi | 0.00002121 | 0.61001231 | 0.00642121 |
| 148.554 | 142.556 | 0 | －1 | 0 | 2 | 0 | 0．00000004 | 0.06000000 | －0．00001．069 | －0．000D00085 | － 1.00001 luta |  |
| 152.656 |  | 1 | 1 | 2 | －2 | 2 | O． 0 0000nol | a． 0 noonoti | －0．00enn2．40 | －0．0nonooz40 | －0．00000240 | －0．000101124 |
| 153.645 | － | 1 | 0 | 2 | －2 | 1 | 0.000000104 | 0.001000004 | －0．0not1075 | －0．00001075 | －0．00001075 | －0．000131475 |
| 153.655 | 177.455 | 1 | 0 | 2 | －2 | 2 | 0． 000000115 | $0.0 r 000016$ | －0．00un4555 | －0．00004932 | －0．00004555 | －0．60004493\％ |
| 154.656 | 176.454 | 1 | 1 | 0 | 0 | 0 | －0．00000002 | －0．00000000 | 0.000004921 | 0.00006079 | 0.00000491 | 0.00000019 |
| 155.435 | 175.675 | －1 | 0 | 2 | 0 | 0 | －0．0000n002 | －0．00000000 | $0.0 n \mathrm{O} 00557$ | 0.00000071 | 0.000100557 | 0．000000t21 |
| 155.445 | 175.665 | －1 | 0 | 2 | 0 | 1 | 0.00000009 | 0.00000012 | －0．00002888 | －0．010003803 | －0．00002888 | － 0.00003 ¢13 |
| 155.455 | 175.655 | －1 | 0 | 2 | 0 | 2 | 0.00000057 | 0.00000062 | －1）． 000173.57 | －0．0001 aras | －0．00017357 | －0．0001 8904 |
| 155.645 | 175.465 | 1 | 0 |  | 0 | －1 | 0.00000026 | －0．00000035 | －0．00007823 | 0.00010708 | －0．00007823 | 0.00016708 |
| 155.655 | 175.455 | 1 | $\bigcirc$ | ） | 0 | 0 | 1．00000319 | $0.0 n \mathrm{omon} 12$ | －0．0nu970a3 | －1）．00003507 | －0．00097084 | －0．0000 35ili |
| 155.665 | 175.445 | 1 | 0 | 0 | 0 | 1 | 0.00000029 | $0.0 斤 000038$ | －0．0nonr 704 | －0．00011451 | －0．00008704 | －0．00011451 |
| 155.675 |  | 1 | 0 | 0 | 0 | 2 | －0．00nu000］ | －0．0n000001 | $0.000 n 0288$ | C．Onomologr | 6.00000288 | $0.0 n \mathrm{mogez}$ |
| 156.555 | 174.555 | 0 | 0 | 0 | 1 | 0 | －0．00000002 | －0．00000000 | 0.00000524 | 0.00000018 | 0．00000524 | a．onourolas |
| 156.654 | 174.456 | 1 | －1 | 0 | 0 | 0 | 0．006000no？ | $0.0 n 000000$ | －0．0nuno5a9 | －0．000nom？ | －0．00000589 | －0．0nctoruz |
| 157.445 | 173.865 | －1 | 0 | 0 | 2 | －1 | 0.00000005 | －0．00000007 | －0．00001506 | 0.00002046 | －0．00001506 | ． 0.00002046 |
| 157．455 | 173.655 | －1 | 0 | 0 | 2 | 0 | O． 00000061. | 0.000000042 | －0．00019533 | －0．00000579 | －0．00018533 | －0．00000．574 |
| 157.465 | 173.645 | －1 | 0 | 0 | 2 | 1 | 0.00000008 | 0.00000008 | －0．00001208 | －0．00002379 | －0．00001808 | － 0.00002379 |
| 158.454 | 172.656 | －1 | －1 | 0 | 2 | 0 | 0.00000003 | $0.0000 n 000$ | －0．00u00786 | －0． 0 （noonola | －0．0000tipra | －0．00000m？2 |
| 161.557 | － |  | 2 | 2 | －2 | 2 | －0．00000002 | －0．00000002 | 0.00000894 | ． 0.00000804 | 0.00000894 | 9.00006694 |
| 162.55 h | 168.554 | 0 | ， | 2 | －2 | 2 | －0．00000053 | －0．000000．59 | 0.00016284 | U．nonl77？ | 0.00015284 | 0.00017721 |
| 163.535 | 167.575 | 0 | 0 | 2 | －2 | 0 | －0．00000002 | －0．00000．000 | 0.00000458 | 0.00000003 | 0.00000458 | 0.00000003 |
| 163.545 | 167.565 | 0 | 0 | 2 | －2 | 1 | 0.00000009 | 0.00000012 | －0．00002801 | －0．00003745 | －0．00002801 | －0．00003745 |
| 163.555 |  | 0 | 0 | 2 | －2 | 2 | －0．00000002 | －0．00000002 | $0.000 \cap 0493$ | 0.00000493 | 0.00000493 | 0.00000493 |
| 183.555 | 167．555 | 0 | 0 | 2 | －2 | 2 | －0．00000909 | －0．0n000990 | 0.00276942 | 0.00201587 | 0.00276945 | 0.00301590 |
| 163.557 | 167.553 | 0 | 2 | 0 | 0 | 0 | 0.00000001 | 0.00000000 | －0．00000362 | －0．0000non？ | －0．000003 A2 | －0．00000u02 |

## Table 9-3-(continued)

COEFFICIFNTS FOK FIHLER ANGLF PFRTIIPATIGOS IN SFCRNDS OF ARC.

CISSINFS FIHR DFLTA THFTA
SINFS FOR MFLTA HSI = SIMITHFTAI

| Tiost | ARGUMENT NUMHERS | COEFFICIENTS |  |  |  | ${ }^{0} \mathrm{~F}$ | ROTATIOM AXIs |  | AXIS OF FIGATRF |  | $\begin{gathered} \text { TEKRFSTHTAL } \\ Z \Delta x 15 \end{gathered}$ |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
| CODE |  | L | L* | F | $1)$ | un | cos | S] ${ }^{4}$ | cos | 510 | Cns | ¢1: |
| 163.755 | 167.355 | 2 | 0 | 0 | -2 | 0 | 19.0000nons | 1).060000008 | -0.0nu0n85 5 | -0.0nomonom |  |  |
| 164.554 | - | 0 | -1 | 2 | -2 | 2 | 0.00000008 | 0.081000004 | -0. (nnuoz2414 | -13.00002414 | -1. 00002414 | -r.aromez414 |
| 184.556 | 166.554 | 0 | 1 | 0 | 0 | U | 0.00000046 | 1).0006000\% | - 0 . Onularar | -0. กn000638 | -0.016013 3anc |  |
| 165.54 .5 | 165.565 | n | 0 | 0 | 0 | -1 | 4.0000n 329 | -0.0n0008442 | -0.00100274 | 0.001344 Cl + | - 1.0001004275 | 0.6ill 134648 |
| 165.555 | . | 0 | 0 | 0 | 0 | 0 | $0.0000 n 905$ | 0.0moungob | -0.010275A09 | -0. M1275409 | -0.00275813 | -1..unj $75 \cdots 3$ |
| 165.555 | . | $1)$ | 0 | 0 | 0 | 0 | 0.00041946 | 0.000101046 | -0.00592617 | -0.0059? 017 | -6.10154? ¢\% 4 | -19, 10195420\%4 |
| . | 185.575 | 0 | O | 0 | 0 | -2 | -0.0unonoor | 0. O (rumacher | 0.nnunzal8 | - 0.0 mon 5 c 18 |  |  |
| - | 173.445 | 1 | 0 | -2 | 2 | -1 | 0. 16000001 | - 6.00600001 | -0.00000270 | 0. manome76 |  |  |
| - | 175.475 | 1 | 0 | 1 | 0 | -2 | -0.0000 0000 | O.0.abiowil | n.mnonoras | -6.00000305 |  | -8.0.mbucab |
| - | 185.365 | 2 | 0 | 0 | 1 | -1 | 6.00000002 | -0.06000002 | -11. 0 nonot 72 | 0.000000732 | -6.00000732 | 0.061610732 |
| - | 185.585 | 0 | 0 | 2 | 0 | -1 | (1.0000000) | - 0.0 unumatil | -n.0numozel3 | ".00\%tur) 3 | -i. 6.0 ¢¢!0¢13 |  |
| , | 193.465 | 1 | 0 | 0 | 2 | -1 | 0.000 00001 | -0.00:000001 | -n. 0 0umoz22 | Q. momozz? | -(1.00040222 | G006mbrez |
| - | $1 E 3.455$ | 1 | 0 | 2 | 2 | 2 | 0.00000001 | -0.03000ticil | -0.0n000167 | (1.0nturs) 4.7 | -11.03以wolt? |  |

8

```
Explanation of symbols
    L, LP, F, D and OM are Brown'g fundamental
    arguments; l, 伯, F,D and \Omega
Constants
        k=0
        K}=-7!552430\times1\mp@subsup{0}{}{3}\mp@subsup{\textrm{Jullan}}{\mathrm{ centhry }}{
        K
    (C-A)/C = 3.272930 \times10
        \Omega=360:9856 day }\mp@subsup{}{}{-1
    (C,
        J}=1.082545\times1\mp@subsup{0}{}{-3
```



# Coefficients for Perturbations in the Euler Angles of the Rotation Axis, the Axis of Figure and the Angular Momentum Vector in a Deformable Earth 

C.IEFFICIFNTS FOR FILLFR ANGLF PFRTIIRATIUINS<br>IN SFCHNDS MF ARC<br>COSINFS FOR DFLTA THETA<br>SINFS FRR DFLTA PSI ; SIN(THFTA



# Table 9-4-(continued) 

COEFFICIFNTS FOR FILLER MMGLF PERTIRRATIONS
IN SFC.MMOS OF ARC
COSINFS FOR MELTA THFTA
SINFS FMR MFLTA YSI * SINGTHFTA:

| $\begin{array}{r} \text { TIUAL } \\ \text { CODE } \end{array}$ | ARGUMENT NUMBERS | CDEFFICIFNTS |  |  |  | UF | $\begin{aligned} & \text { ROTATION } \\ & \text { AXTS } \end{aligned}$ |  | AXIS nf Figilikf |  | $\begin{aligned} & \text { TFRKESTRIAL } \\ & Z \text { AXIS } \end{aligned}$ |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  | L | LP | F | L | U14 | cos | SIN | c 0 S | STM | cns | SIM |
| 149.745 | - | ? | 0 | 2 | -2 | 1 | 0.00000102 | 0.00000102 | -0.00031151 | -0.00031151. | -0.000000245 | -0.0nturicas |
| 143.755 | - | 2 | 0 | 2 | -2 | 2 | 0.00000578 | 0.000100578 | -0.00176001 | -0.00176001 | -0.00001384 | -0.0n001384 |
| 144.54 \% |  | 0 | 1 | 2 | 0 | 1 |  | 0.000000077 | -0.00023363 | -0.00n? 3383 | -0.0000ndra | - 6.0 (1)torris ${ }^{\text {a }}$ |
| 144.556 | - | 0 | 1 | 2 | 0 | 2 | 0.000006655 | 0.000016 ta | -0.00202476 | -0.0n20734 | -0.90001589 |  |
| 145.535 | 185.575 | 0 | 0 | 2 | 0 | U | 0.0m032329 | 0.00000001 | -0.00679414 | -1). $n$ nonoz 50 | -0.010004952 | -0.1006045tz |
| 145.545 | 185.565 | 0 | 0 | 2 | U | 1 | -0.00031027 | -0.00041647 | 0.00448886 | 0.12682962 | 0.00075648 | $0.000912 \%$ |
| 145.555 | 185.555 | 0 | 0 | 2 | 0 | 2 | -0.00184457 | -0.05201046 | 0.56174046 | (1. 121275932 | 0.00044214 A | 1). 10 (1446327 |
| 145.755 | 185.355 | 2 | 0 | 0 | 0 | 0 | 0.00002469 | 0.00000616 | -0.00751990 | -0.00004945 | -0.00005488 | -10.00000430 |
| 145.765 |  | 2 | 0 | 0 | 0 | 1 | 0.00000205 | . 0.041000205 | -n.0006zフ99 | -0.006) O 2.90 | -0.000000487 | -0.006mutari |
| 146.544 |  | 0 | -1 | 2 | 0 | 1 | -0.00000061 | -0.0n000061 | 0.00018690 | $0.0001860 n$ | 0.00000146 | 1. 00000414t |
| 146.554 |  | ก | -1 | 2 | 0 | 2 | -0.00000588 | -0.0n000588 | 0.00179108 | 0.00170108 | 9.000061397 | 0.00081297 |
| 14.7.355 | - | -2 | 0 | 2 | 2 | 2 | 0.00000107 | 0.00000107 | -0.00032706 | -0.00032706 | -0.00000255 | -0.00000<5s |
| 147.545 | 183.565 | 0 | 0 | 13 | 2 | -1 | 1. 000000419 | -0.0n000562 | -1).00127608 | n.0017121t. | -0. 016000846 | C.00mbllrt. |
| 147.555 | 183.555 | 0 | 0 | 0 | 2 | 0 | 0.00005025 | -0.00000003 | -0.01530437 | 0.00001038 | -0.00011155 | -0.000001744 |
| 147.565 | 183.545 | 0 | 0 | 0 | 2 | 1 | -0.00000465 | - 12.00000529 | 0.00141743 | 0.00191547 | 0.00001127 | 0.000014ヶ5 |
| 148.554 | 182.556 | 0 | -1 | 0 | 2 | 0 | 0.00000332 | 0.00000005 | -0.00101199 | -0.00001590 | -0.00000738 | -0.00000659 |
| 152.656 | - | 1 | 1 | 2 | -2 | 2 | 0.00000077 | 0.00000072 | -0.00021201 | -0.000731801 | -0.000008185 | - 6.0 aroogith |
| 153.645 |  | 1 | 0 | 2 | 2 | 1 | 0.00000322 | 0.00000322 | -0.00098104 | -0.00098104 | -0.00000743 | -0.60000743 |
| 153.655 | 177.455 | 1. | 0 | 2 | -2 | 2 | 13.00001360 | 0.01001483 | -0.0n414225 | -0.0045].582 | -0.0101103147 | -0.0010434.17 |
| 154.656 | 176.454 | 1 | 1 | 0 | 0 | 0 | -0.00000153 | -0.00000000 | 0.00046706 | 0.000000080 | 0.00000339 | 0.vomonula |
| 155.435 | 175.675 | -1 | 0 | 2 | 0 | 0 | -0.00000174 | -0.0nounaron | 0.00052934 | 0.00000010 | 0. 000100345 | 6. 0 000001614 |
| 155.445 | 175.665 | -1 | 0 | 2 | 0 | 1 | 0.00000859 | 0.00001156 | -0.0026ítz | -0.001351904 | -0.00001495 | -0.0040\%207 |
| 155.455 | 175.655 | -1 | 0 | 2 | U | 2 | 1).00005210 | 0.000056 kl | -0.01586778 | -0.01729985 | -0.00011989 |  |
| 255.645 | 175.465 | 1 | 0 | 0 | 0 | -1 | 0.00002566 | -0.00003435 | -0.00781363 | 0.01046081 | -0.00405403 | 0.00007345 |
| 155.655 | 175.455 | 1 | 0 | 0 | 0 | 0 | 0. 000330306 | 0.00000005 | -0.09220195 | -0.00001 6.67 | -0.00067058 | -0.000022421 |
| 155.665 | 175.445 | 1 | 0 | 0 | 0 | 1 | 0.00002593 | 0.00003482 | -0.00789530 | -0.01060378 | -0.00006012 | $-0.00007910$ |
| 155.6.75 | 17. | 1 | 0 | 0 | 0 | 2 | -0.00000087 | -0.00000087 | 0.00026472 | 0.000126472 | 0.010000199 | 11.00000139 |
| 156.555 | 174.555 | 0 | 0 | 0 | 1 | 0 | -0.00000164 | -0.00000000 | 0.00049920 | 0.00000008 | 0.00000362 | 0.00000012 |
| 156.654 | 174.456 | 1 | -1 | 0 | U | 0 | 0.00000184 | 0.00000000 | -0.00056049 | -0.00000000 | -0.00000407 | - 1.0000000144 |
| 157.445 | 173.665 | -1 | 0 | 0 | 2 | -1 | 0.00000491 | -0.00000654 | -0.00149429 | 0.00199257 | -0.00001040 | 0.00001412 |
| 157.455 | 173.655 | -1 | 0 | 0 | 2 | 0 | 0.000005787 | 0.00000001 | -0.01762385 | -0.00000275 | -0.00012801 | - 0.000000740 |
| 157.465 | 1.73.6.45 | -1 | 0 | 0 | , | 1 | 0.00000542 | 0.00000726 | -0.00165064 | -0.00221102 | -0.00001249 | -0.00001643 |
| 158.454 | 172.656 | -1 | -1 | 0 | 2 | 0 | 0.00000245 | 0.051000000 | -0.00074730 | -0.00000011 | - 0.0000054 .3 | -0.00000015 |
| 1.61 .557 |  | -0. | 2 | 2 | -2 | 2 | -0.00000215 | -0.00000215 | 0.00065362 | 0.00065362. | -.0.000004779 | 0.000004 .29 |
| 162.556 | 16R. 554 | , | 1 | 2 | -2 | 2 | -0.00005045 | -0.00005498 | 0.01536357 | 0.01674252 | $0.00011<48$ | 0.01012241 |
| 163.535 | 16.7 .575 | 0 | 0 | 2 | -2 | 0 | -0.00000143 | -0.00000000 | 0.00043592 | 0.00000001 | 0.00000316 | 0.00000002 |
| 163.545 | 167.565 | 0 | , | 2 | -2 | 1 | 0.00000869 | 0.000011 ht | -0.00264678 | -0.00354974 | -0.00001935 | -0.00002587 |
| 163.555 | - | 0 | 0 | 2 | -2 | 2 | -0.00000153 | -0.00000153 | 0.00046707 | 0.00046707 | 0.00000341 | 0.00000341 |
| 163.555 | 167.555 | 0 | 0 | 2 | -2 | 2 | -0.0008605? | -0.00093797 | 0.26206034 | 0.28564450 | 0.00191293 | 6.0n208316 |
| 163.557 | 167.553 | 0 | 2 | 0 | 0 | 0 | 0.00000113 | 0.00000000 | -0.00034475 | -0.00000001 | -0.00000250 | -0.00000001 |

Table 9-4-(continued)

COEFFICIFNTS FOR FILER ANGLF PFRTHRHATIONS
IN SFCOMOS OF aRC
COSIMES FUR DELTA THFTA
SINES FRR DFLTA PSI * SIN(THFTA)

| $\begin{gathered} \text { TIOAL } \\ \text { CODE } \end{gathered}$ | RGUMENT <br> NUMBERS | C.IEFFICIFNTS |  |  |  | $\mathrm{OF}_{\mathrm{OM}}$ | ROTATITAN AXIS |  | AXIS กF <br> FIGIRF |  | TFKRESTRIAL 2 AXIS |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
|  |  | L | LP | F | 0 |  | $\cos$ | SIN | cns | SIN | Cus | 5 JM |
| 183.755 | 167.355 | 2 | 0 | 0 | 2 | 0 | 0.00000266 | 0.0nomonoun | -0.000R0957 |  |  |  |
| 164.554 |  | 0 | -1 | 2 | -2 | 2 | 0.00000752 |  | -0.00229118 | -0.00000602 -0.00229118 | -0.00000587 -0.00001667 | $-0.0066 m 013$ -0.00001567 |
| 164.556 | 166.554 | 0 |  | 0 | 0 | 0 | 0.00004335 | $0.00100 n 00$ | -0.002320104 | -0.00229118 -0.0000018 | -0.00001667 | - 0.00001067 |
| 165.545 | 165.565 | 0 | 0 | 0 | 0 | -1 | 0.00031 .348 | -0.01042084 | -0.00546715 | 0.12816140 | -0.00009579 $-0.00064<62$ | -12.00600026 0.010692990 |
| 165.555 | - | 0 | 0 | 0 | 0 | 0 | 0,00086146 | 0.00086146 | -0.76234498 | -0. ${ }^{\text {a }}$ /34408 |  | 0.060192990 -0.00140371 |
| 165.555 |  | 0 | 0 | 0 | 0 | 0 | 0.00185233 | 0.00185233 | -0.56410116 | -0.5641G11A | -0.00190371 -0.00409338 | -0.00140371 -0.00409338 |
| - | 155.575 | 0 | 0 | 0 | 0 | -2 | -0.000007a7 | 0.001000787 | 0.00239757 | -0.002 29757 | -0.00409338 1.0001734 | -0.00409338 $-6.0 n 001739$ |
| - | 173.445 | 1 | 0 | -2 |  | -1 | 0.00000087 | -0.00000087 | -0.00026463 | $0.000264+3$ | \%,00017 | - $0 \cdot 01001734$ |
| - | 175.475 | 1 | 0 | 0 | 0 | -2 | -0.00000nth | 0.00000066 | 0.00020736 | -0.0002023a | -0.00000188 | 0.0000018 e |
| - | 195.365 | 2 | 0 | 0 | 0 | -1 | 0.00000245 | -0.0000 1245 | -0.00074704 | 0.00074704 | -.monol 42 | -0.0mmorit ${ }^{2}$ |
| - | 185.585 | 0 | 0 | 2 | 0 | -1 | 0.0000007 ? | -0.0000007? | -0.00021789 |  | -0.00000506. | O.00malibos |
| - | 193.465 | 1 | 0 | 0 | 2 | -1 | 0.00040077 | -0.000000 77 | -0.000 33242 | O.10n21784 | -0.00000147 | 6.00006147 |
| - | 1 E 3.455 | 1 | 0 | 2 | 2 | 2 | 0.100000061 | -0.0n000061 | -0.0n014668 | 0.00023342 | -6.00000 54 | 0.00000154 |
|  |  |  |  |  |  |  |  | - | -0.010 | O.0001Rtar | -0.00000115 | O. 0 monils |

Explanation of symbole
$\mathrm{L}, \mathrm{LP}, \mathrm{F}, \mathrm{D}$ and OM are Rrown's fundamental arguments; $\ell, \ell^{\prime}, F, D$ and $\cap$

## $\xrightarrow{\text { Constants }}$

$k=0$

$\mathrm{K}_{\mathrm{G}}=-3: 484150 \times 10^{3} \mathrm{Jultan}$ rentury ${ }^{-1}$
$(\mathrm{C}-\mathrm{A}) / \mathrm{C}=3.272930 \times 10^{-3}$
$\Omega=36059856$ day $^{-1}$
$\left(C_{6} / a_{E}\right)=2.343852 \times 10^{4}$
$\mathrm{J}_{2}=1.082645 \times 10^{-3}$
$\left(m_{E} / m_{E}\right)=3.334 .320 \times 10^{6}$


Figure 9.1. The Effect of Polar Motion on Latitude and Time.

Figure 9.2 shows the diurnal part of the latitude variation for Goddard Space Flight Center during the summer of 1970 . The polar motion components $\mathrm{m}_{1}$ and $\mathrm{m}_{2}$ in Equation (9.14) were computed using all of the coefficients from Table 9.2. Of the 135 tidal constituents, those with argument numbers
135.655
145.545
145.555
163.555
165.555
165.565


Figure 9.2. Diurnal Variation in the Latitude of Goddard Scpace Flight Center from June 22, 1970 to August 29, 1970, The Love Number $k=0.29$,
have amplitudes large enough to dominate the motion. In Figure 9.3 the diurnal latitude variation based upon all 135 tidal constituents is compared to that obtained by using only the terms with the argument numbers (9.16). The error due to the neglected terms is at most 0.0009 or 2.8 cm at the Earth's surface.


Figure 9.3. Effect of Neglecting all but the 6 Largest Tidal Components in Computing the Diurnal Variation of Latitude for Goddard Space Flight Center. Both curves are for a Deformable Earth with $\mathrm{k}=0.29$.

The effect of rotational and tidal deformation on the diurnal polar motion components is shown in Figure 9.4. The diurnal motion of the rotation axis within the Earth is affected only slightly by deformation, being decreased in amplitude by $0!00024$ which corresponds to 0.7 cm . In contrast to this, deformation has a substantial effect on the motion of the angular momentum vector within the Earth. As shown in Figure 9.5, the diurnal motion of the angular momentum vector is reduced in amplitude by $0: 00645$ or 20 cm . The shift in the angular momentum vector is due primarily to the factor

$$
\begin{equation*}
\left(1-\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right) \cong \frac{2}{3} \tag{9.17}
\end{equation*}
$$



Figure 9.4. The Effect of Rotational and Tidal Deformation on the Diurnal Motion of the Rotation Axis.


Figure 9.5. The Effect of Rotational and Tidal Deformation on the Diurnal Motion of the Angular Momentum Vector.


Figure 9.6. Diurnal Motion of the Axis of Figure Due to Rotational and Tidal Deformation.
which appears in the diurnal part of Equation (7.31) for $\mathrm{H} / \mathrm{C} \Omega$ but is absent from the diurnal part of Equation (7.18) for m. Rotational and tidal deformation cause the direction cosines of the Earth's axis of figure to oscillate with an amplitude of 1.9 or 58.7 m , as shown in Figure 9.6. The motion of the axis of figure is due mainly to the mass redistribution associated with the diurnal tidal bulge. The direction cosines of the angular momentum vector are given by Equation (2.24) with $\mathrm{h}=0$ as

$$
\begin{equation*}
\frac{\mathrm{H}}{\mathrm{C} \Omega}=\frac{\mathrm{A}}{\mathrm{C}} \mathrm{~m}+\frac{\mathrm{c}}{\mathrm{C}} \tag{9.18}
\end{equation*}
$$

and the direction cosines of the axis of figure are, from Equation (E.9),

$$
\begin{equation*}
\psi_{f}=\frac{\mathrm{C}}{\mathrm{C}-\mathrm{A}} \tag{9.19}
\end{equation*}
$$

Therefore

$$
\begin{equation*}
\frac{\mathrm{H}}{\mathrm{C} \Omega}=\frac{\mathrm{A}}{\mathrm{C}} m+\left(\frac{\mathrm{C}-\mathrm{A}}{\mathrm{C}}\right) \psi_{\mathrm{f}} \tag{9.20}
\end{equation*}
$$

so that the 58.7 m diurnal motion of $\psi_{\mathrm{f}}$ is reduced by the factor ( $\mathrm{C}-\mathrm{A}$ )/C to the 20 cm departure of $\mathrm{H} / \mathrm{C} \Omega$ from the position that it would occupy in a rigid Earth.

The theory developed in Section 7 is directly applicable to the problem of transforming from an inertial coordinate system to a terrestrial system that is fixed to a set of observatories in some prescribed manner. The geometry involved in transforming from inertial to terrestrial coordinates is shown in Figure 9.7. The precessional transformation is

$$
\begin{equation*}
\vec{x}_{\text {mean of date }}=R_{3}\left(-z_{p}\right) R_{2}\left(\theta_{p}\right) R_{3}\left(-\zeta_{0}\right) \vec{x}_{\text {mean of epoch }} \tag{9.21}
\end{equation*}
$$

where $z_{p}, \theta_{p}$, and $\zeta_{p}$ are the precessional elements written as $z, \theta$, and $\zeta_{0}$ in the Explanatory Supplement to the AENA [1961, p. 29]. The transformation from the mean sidereal system of date to the true sidereal system of date is

$$
\begin{equation*}
\vec{x}_{\text {true of date }}=R_{1}\left(-\epsilon_{T D}\right) R_{3}\left(-\Delta \psi_{\mathrm{rD}}\right) \mathrm{R}_{1}\left(\epsilon_{\mathrm{M}}\right) \overrightarrow{\mathrm{x}}_{\text {mean of date }} \tag{9.22}
\end{equation*}
$$

where the true obliquity of date is given by

$$
\begin{equation*}
\epsilon_{\mathrm{TD}}=\epsilon_{\mathrm{M}}+\Delta \epsilon_{\mathrm{rD}} \dot{D} \tag{9.23}
\end{equation*}
$$



Figure 9.7. Transformation From Inertial Coordinates to Terrestrial Coordinates.

The nutations in longitude and obliquity referred to the mean ecliptic of date are denoted by $\Delta \psi_{\mathrm{rD}}$ and $\Delta \epsilon_{\mathrm{rD}}$ respectively. The nutation in longitude is reckoned positive westward. The transformation from the true sidereal system of date to the terrestrial system is

$$
\begin{equation*}
\vec{x}_{\text {terrestrial }}=R_{2}\left(-m_{1}\right) R_{1}\left(m_{2}\right) R_{3}(\text { GASTI }) \vec{x}_{\text {true of date }} \tag{9.24}
\end{equation*}
$$

The transformations (9.21), (9.22) and (9.24) are combined to form the complete transformation from inertial to terrestrial coordinates.

In Woolard's development of the theory of precession and nutation, the nutations $\Delta \psi_{H}$ and $\Delta \epsilon_{\mathrm{H}}$ for the angular momentum vector referred to the fixed mean ecliptic of epoch are obtained as a result of integrating Poisson's equations and are given in Woolard's Table 24 [1953, p. 138]. The Euler angle perturbations $\Delta \psi_{\mathrm{r}}$ and $\Delta \theta_{\mathrm{r}}$ computed from Woolard's Equations (55) are used to convert $\Delta \psi_{\mathrm{H}}$ and $\Delta \epsilon_{\mathrm{H}}$ into the nutations $\Delta \psi_{\mathrm{r}}$ and $\Delta \epsilon_{\mathrm{r}}$ corresponding to the rotation axis.

$$
\begin{align*}
\Delta \psi_{\mathbf{r}} & =\Delta \psi_{\mathbf{H}}-\delta \psi_{\mathbf{r}}  \tag{9.25}\\
\Delta \epsilon_{\mathbf{r}} & =\Delta \epsilon_{\mathbf{H}}+\delta \theta_{\mathbf{r}} \tag{9.26}
\end{align*}
$$

In (9.25), $\delta \psi_{\mathrm{r}}$ is subtracted because it is reckoned positive eastward whereas $\Delta \psi_{\mathrm{r}}$ and $\Delta \psi_{\mathrm{H}}$ are reckoned positive westward. The values of $\Delta \psi_{\mathrm{r}}$ and $\Delta \epsilon_{\mathrm{r}}$ are listed in Woolard's Table 24 as well as his values for the Euler angle perturbations $\delta \psi_{\mathrm{r}}$ and $\delta \theta_{\mathrm{r}}$. The precession and nutation results represented in Table 24 undergo a reduction to the mean ecliptic of date. This reduction consists of adding the corrections in Woolard's Table 25 [1953, p. 152] to the entries in Table 24 so as to produce Table 26 [1953, p. 153].

In order to apply the deformable-Earth theory of Section 7 in connection with Woolard's solution for the angular momentum vector, his rigid-Earth values for $\delta \psi_{r}$ and $\delta \theta_{r}$ must first be removed in Table 24 so as to give the precession and nutation of the angular momentum vector referred to the mean ecliptic of epoch. Reduction to the mean ecliptic of date is accomplished by adding the terms listed in Woolard's Table 25 to the modified entries from Table 24.

The rigid-Earth values of $\delta \psi_{\mathrm{r}}$ and $\delta \theta_{\mathrm{r}}$ from the diurnal terms in Woolard's Equations (55) are of order 0.000050 r 0.15 cm . They are therefore small enough to be lost in the numerical round-off error involved in truncating the 5th decimal place in Table 24 prior to presenting the nutation series in seconds of arc to 4 decimal places as Table 26 and as Table 2.5 in the Explanatory Supplement to the AENA [1961, p. 44]. The Eulerian terms in Woolard's (55) are of order
$0!0005$ or 1.5 cm and are therefore large enough to affect the 4th decimal place in the nutation series. The Eulerian terms are not included in Woolard's nutation tables, however, because they involve the arbitrary constant of integration $\mathrm{m}_{0}$. The nutation series as presented in Woolard's Table 26 and in Table 25 of the Explanatory Supplement to the AENA therefore represent the direction in space of the Earth's angular momentum vector to within the 4 decimal places given in the tables, and subject to the understanding that the Eulerian terms in Equations (7.39) and (7.40) are to be added to the tabular values of $\delta \theta_{r D}$ and $\delta \psi_{\mathrm{rD}}$.

The Euler angle perturbations $\delta \theta_{r}$ and $\delta \psi_{r}$ of Equations (7.39) and (7.40) must be reduced to the mean ecliptic of date before they can be applied to the nutation series in Table 2.5 of the Explanatory Supplement to the AENA. The Euler angle perturbations $\delta \theta_{\mathrm{rD}}$ and $\delta \psi_{\mathrm{rD}}$ referred to the mean ecliptic of date are given by

$$
\begin{equation*}
\delta \theta_{r D}+i \delta \psi_{r D} \sin \epsilon_{T D}=-i e^{i(\operatorname{GASTI})}\left(m-\frac{H}{C \Omega}\right) \tag{9.27}
\end{equation*}
$$

which is analogous to Equation (6.10). It follows from (9.27) and (6.10) that

$$
\begin{equation*}
\delta \theta_{r D}+\mathbf{i} \delta \psi_{r D} \sin \epsilon_{T D}=\left(\delta \theta_{r}+\mathbf{i} \delta \psi_{r} \sin \theta\right) \mathrm{e}^{\mathrm{i}(\mathrm{GASTI}-\phi)} \tag{9.28}
\end{equation*}
$$

The Greenwich apparent sidereal time is related to the Greenwich mean sidereal time $\phi_{M}$ by

$$
\begin{equation*}
\mathrm{GAST} 1=\phi_{\mathrm{M}}+\Delta \psi_{\mathrm{rD}} \cos \epsilon_{\mathrm{TD}} \tag{9.29}
\end{equation*}
$$

and $\phi_{\mathrm{M}}$ is related to the Euler angle $\phi$ by Equation (C.4). Equations (9.29) and (C.4) are combined to give

$$
\begin{equation*}
\operatorname{GAST} 1-\phi=-\mathrm{a}+\gamma \cos \left(\Gamma+\phi_{\mathrm{M}}\right) \cot \epsilon_{\mathrm{M}} \tag{9.30}
\end{equation*}
$$

The planetary precession a is, from Woolard's Equation (67),

$$
\begin{align*}
\mathrm{a}= & 12.473 \mathrm{~T}-2.43804 \mathrm{~T}^{2} \\
& -0.00133 \mathrm{~T}^{3} \tag{9.31}
\end{align*}
$$

where T is measured in Julian centuries since 1900 Jan 0.5 ET , and the polar motion amplitude $\gamma$ is of order $0!15$. The reduction (9.28) of the Euler angle perturbations to the mean ecliptic of date therefore involves adding terms of order $0!0000004$ and can be neglected in numerical computations.

The nutations $\Delta \psi_{\mathrm{rD}}$ and $\Delta \epsilon_{\mathrm{rD}}$ in longitude and obliquity referred to the mean ecliptic of date are given by

$$
\begin{align*}
& \Delta \psi_{\mathrm{rD}}=\Delta \psi_{\mathrm{HD}}-\delta \psi_{\mathrm{rD}}  \tag{9.32}\\
& \Delta \epsilon_{\mathrm{rD}}=\Delta \epsilon_{\mathrm{HD}}+\delta \theta_{\mathrm{rD}} \tag{9.33}
\end{align*}
$$

Neglecting second order terms,

$$
\begin{align*}
& \delta \psi_{\mathrm{rD}}=\delta \psi_{\mathrm{r}}  \tag{9.34}\\
& \delta \theta_{\mathrm{rD}}=\delta \theta_{\mathbf{r}} \tag{9.35}
\end{align*}
$$

The nutations $\Delta \psi_{\mathrm{HD}}$ and $\Delta \epsilon_{\mathrm{HD}}$ are taken from Table 2.5 of the Explanatory Supplement to the AENA.

It is now possible to write down a transformation from inertial to terrestrial coordinates which takes into account the diurnal motion of the rotation axis and the angular momentum vector within a deformable Earth. The form of the transformation is

$$
\begin{align*}
\overrightarrow{\mathrm{x}}_{\text {terrestrial }}= & \mathrm{R}_{2}\left(-\mathrm{m}_{1}\right) \mathrm{R}_{1}\left(\mathrm{~m}_{2}\right) \mathrm{R}_{2}\left(-\mathrm{H}_{1} / \mathrm{C} \Omega+\mathrm{m}_{1}\right) \mathrm{R}_{1}\left(\mathrm{H}_{2} / \mathrm{C} \Omega-\mathrm{m}_{2}\right) \\
& \mathrm{R}_{3}\left(\mathrm{GAST} 1+\delta \psi_{\mathrm{rD}} \cos \epsilon_{\mathrm{HD}}\right) \\
& \mathrm{R}_{1}\left(-\epsilon_{\mathrm{HD}}\right) \mathrm{R}_{3}\left(-\Delta \psi_{\mathrm{HD}}\right) \mathrm{R}_{1}\left(\epsilon_{\mathrm{M}}\right) \\
& \mathrm{R}_{3}\left(-z_{\mathrm{p}}\right) \mathrm{R}_{2}\left(\theta_{\mathrm{p}}\right) \mathrm{R}_{3}\left(-\zeta_{0}\right) \overrightarrow{\mathrm{x}}_{\text {mean of epoch }} \tag{9.36}
\end{align*}
$$

and the associated geometry is shown in Figure 9.7. The transformation from the mean system of epoch to the mean system of date is the same as in Equation (9.21). Instead of a transformation from the mean equinox of date $P_{M}$ directly to the true equinox of date $\Upsilon_{T}$ as in Equation (9.22), (9.36) involves a transformation from $\Upsilon_{M}$ to the ascending node of the mean ecliptic of date on the equator normal to the angular momentum vector. The obliquity $\epsilon_{\mathrm{HD}}$ is given by

$$
\begin{equation*}
\epsilon_{\mathrm{HD}}=\epsilon_{\mathrm{M}}+\Delta \epsilon_{\mathrm{HD}} \tag{9.37}
\end{equation*}
$$

The Greenwich apparent sidereal time is modified so as to account for the displacement $\delta \psi_{\mathrm{rD}} \cos \epsilon_{\mathrm{HD}}$. The transformation from the angular momentum vector to the rotation axis is made using the components of $m-H / C \Omega$. The diurnal terms in $\mathrm{m}-\mathrm{H} / \mathrm{C} \Omega$ are computed by differencing the respective components given in Table 9.2. The Eulerian part of $m-H / C \Omega$ is from (9.6) and (9.8),

$$
\begin{equation*}
(\mathrm{m}-\mathrm{H} / \mathrm{C} \Omega)_{\text {Eulerian }}=\left(\frac{\mathrm{C}-\mathrm{A}}{\mathrm{C}}\right)\left(1-\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right) \mathrm{m}_{0} \mathrm{e}^{\mathrm{i} \sigma_{0} \mathrm{t}} \tag{9.38}
\end{equation*}
$$

The vector $m_{0} e^{i \sigma_{0} t}$ must be determined from observations. As shown in Figure 9.8 , it represents the displacement of the rotation pole, with diurnal motion filtered out, relative to the mean pole position. Because of the small factor

$$
\begin{equation*}
\left(\frac{C-A}{C}\right)\left(1-\frac{\mathrm{k}}{\mathrm{k}_{\mathrm{s}}}\right)=0.0023 \tag{9.39}
\end{equation*}
$$

in (9.38), an uncertainty of, say 1 m , in $\mathrm{m}_{0} \mathrm{e}^{\mathrm{i} \sigma_{0}{ }^{t}}$ will produce an uncertainty of only 0.23 cm in ( $\mathrm{m}-\mathrm{H} / \mathrm{C} \Omega)_{\text {Eulerian }}$.


Figure 9.8. Determination of $m_{o} e^{i \sigma_{0} t}$ from Observations.

The transformation from the rotation axis to the $z$ axis of the terrestrial system is based upon the components of the polar motion $m$. There is as yet no reliable model for the "non-diurnal" or "Eulerian" part of the polar motion and it is certainly true that the idealized circular motion given by the terms,

$$
\begin{equation*}
m_{0} e^{i \sigma_{0} t}+\frac{\left(\frac{c_{0}}{C-A}\right)}{\left(1-\frac{k}{k_{s}}\right)} \tag{9.40}
\end{equation*}
$$

in (9.6), does not adequately represent non-diurnal polar motion in the real Earth. In order to represent the polar motion a "semi-analytical" pole path is determined by adding the diurnal terms of (9.6) to an observed pole position with diurnal motion filtered out.

$$
\begin{equation*}
m=m_{\text {diurnal }}+m_{\text {non-diurnal }} \tag{9.41}
\end{equation*}
$$

The polar motion transformation is then broken into an empirical part and an analytical part as follows:

$$
\begin{align*}
R_{2}\left(-m_{1}\right) R_{1}\left(m_{2}\right)= & R_{2}\left(-m_{1, \text { non-diurnal }}\right) R_{1}\left(m_{2, \text { non-diurnal }}\right) \\
& R_{2}\left(-m_{1, \text { diurnal }}\right) R_{1}\left(m_{2, \text { diurnal }}\right) \tag{9.42}
\end{align*}
$$

An alternative form of the transformation (9.36) is

$$
\begin{align*}
\vec{x}_{\text {terrestrial }}= & R_{2}\left(-m_{1}\right) R_{1}\left(m_{2}\right) R_{3}(\text { GAST1 }) \\
& R_{1}\left(-\epsilon_{T D}\right) R_{3}\left(-\Delta \psi_{r D}\right) R_{1}\left(\epsilon_{M}\right) \\
& R_{3}\left(-z_{p}\right) R_{2}\left(\theta_{p}\right) R_{3}\left(-\zeta_{0}\right) \vec{x}_{\text {mean epoch }} \tag{9.43}
\end{align*}
$$

where the true obliquity of date $\epsilon_{\mathrm{TD}}$ is given in terms of the nutation $\Delta \epsilon_{\mathrm{rD}}$ in obliquity by Equation (9.23). The nutations $\Delta \psi_{r D}$ and $\Delta \epsilon_{r D}$ are given by Equations (9.32) and (9.33). The Euler angle perturbations $\delta \theta_{r D}$ and $\delta \psi_{r D}$ are computed from Equations (7.37) and (7.38) with the Eulerian terms determined empirically. The diurnal terms in $\delta \theta_{\mathrm{rd}}$ and $\delta \psi_{\mathrm{rd}}$ are given in Table 9.4. The polar motion transformation can be broken into diurnal and non-diurnal parts as in Equation (9.42).

A second alternative form of Equation (9.36) is

$$
\begin{align*}
\overrightarrow{\mathrm{x}}_{\text {terrestrial }}= & \mathrm{R}_{3}\left[\text { GAST } 1-\left(\delta \psi_{\mathrm{zD}}-\delta \psi_{\mathrm{rD}}\right) \cos \epsilon_{\mathrm{TD}}\right] \\
& R_{1}\left(-\theta_{M}\right) \mathrm{R}_{3}\left(-\Delta \psi_{\mathrm{zD}}\right) \mathrm{R}_{1}\left(\epsilon_{\mathrm{M}}\right) \\
& R_{3}\left(-z_{p}\right) \mathrm{R}_{2}\left(\theta_{\mathrm{p}}\right) \mathrm{R}_{3}\left(-\zeta_{0}\right) \overrightarrow{\mathrm{x}}_{\text {mean of epoch }} \tag{9.44}
\end{align*}
$$

where the obliquity $\theta_{M}$ of the terrestrial equator is given by

$$
\begin{equation*}
\theta_{M}=\epsilon_{M}+\Delta \epsilon_{z D} \tag{9.45}
\end{equation*}
$$

The nutations $\Delta \psi_{2 \mathrm{D}}$ and $\Delta \epsilon_{2 \mathrm{D}}$ are given by

$$
\begin{align*}
& \Delta \psi_{\mathrm{zD}}=\Delta \psi_{\mathrm{HD}}-\delta \psi_{\mathrm{zD}}  \tag{9.46}\\
& \Delta \epsilon_{2 \mathrm{D}}=\Delta \epsilon_{\mathrm{HD}}+\delta \theta_{\mathrm{zD}} \tag{9.47}
\end{align*}
$$

Except for second order terms, the Euler angle perturbations $\delta \theta_{\mathrm{zD}}$ and $\delta \psi_{\mathrm{zD}}$ are the same as the corresponding perturbations referred to the fixed mean ecliptic of epoch.

$$
\begin{align*}
& \delta \psi_{\mathrm{zD}}=\delta \psi_{\mathrm{z}}  \tag{9.48}\\
& \delta \theta_{\mathrm{zD}}=\delta \theta_{\mathbf{z}} \tag{9.49}
\end{align*}
$$

Equations (7.43) and (7.44) are used to compute $\delta \theta_{2 \mathrm{D}}$ and $\delta \psi_{\mathrm{zD}}$. The Eulerian terms are determined empirically and the diurnal terms are given in Table 9.4. No polar motion rotations are needed because the transformation goes directly from the mean sidereal system of date to the terrestrial system, bypassing the rotation axis altogether.

The choice of the best alternative form of the transformation from inertial to terrestrial coordinates depends upon the particular application at hand. Equation (9.36) has the advantage that the nutational part is made using $\Delta \psi_{H D}$ and $\Delta \epsilon_{\mathrm{HD}}$ which come directly from Table 2.5 of the Explanatory Supplement to the AENA. Modification of an existing transformation that does not include diurnal polar motion and deformable Earth effects is accomplished by inserting the transformation

$$
\begin{align*}
& \mathrm{R}_{2}\left(-\mathrm{m}_{1}, \text { diurna1 }\right) \mathrm{R}_{1}\left(\mathrm{~m}_{2, \text { diurnal }}\right) \\
& \mathrm{R}_{2}\left(-\mathrm{H}_{1} / \mathrm{C} \Omega+\mathrm{m}_{1}\right) \mathrm{R}_{1}\left(\mathrm{H}_{1} / \mathrm{C} \Omega-m_{2}\right) \\
& \mathrm{R}_{3}\left(\text { GAST } 1+\delta \psi_{\mathrm{rD}} \cos \epsilon_{\mathrm{HD}}\right) \tag{9.50}
\end{align*}
$$

in place of

$$
\begin{equation*}
\mathrm{R}_{3}(\mathrm{GAST1}) \tag{9.51}
\end{equation*}
$$

Equation (9.43) involves fewer individual rotations than (9.36) but the nutations $\Delta \psi_{\mathrm{HD}}$ and $\Delta \epsilon_{\mathrm{HD}}$ must be modified using (9.32) and (9.33) to form $\Delta \psi_{\mathrm{rD}}$ and $\Delta \epsilon_{\mathrm{rD}}$. Equation (9.44) involves still fewer individual rotations but the absence of a polar motion transformation makes it impossible to incorporate the semianalytical pole path defined by Equation (9.41).

## APPENDIX A <br> THE FUNCTIONAL FORM OF DOODSON'S EXPANSION OF THE TIDAL POTENTIAL

Doodson [1922] developed a formal expansion of the tide generating potential based upon Brown's [1905] lunar theory and Newcomb's [1898] theory of the sun. Doodson's tabulated results are represented here in functional form as

$$
\begin{align*}
& U=\sum_{j}\left[G_{20,} \bar{A}_{20 j 9} \cos \left(\omega_{20 j \rho} t+\beta_{20 j \otimes}\right)\right. \\
& \left.+G_{20 \lambda} \bar{A}_{20 j \odot} \cos \left(\omega_{20 j \odot} t+\beta_{20 j \odot}\right)\right] \\
& +\sum_{j}\left[G_{21 \nu} \bar{A}_{21 \mathrm{j},} \sin \left(\omega_{21 \mathrm{j} \vartheta} \mathrm{t}+\beta_{21 \mathrm{j} 9}+\lambda\right)\right. \\
& \left.+G_{21}, \bar{A}_{21 j \odot} \sin \left(\omega_{21 \mathrm{j} \odot}{ }^{t}+\beta_{21 j \odot}+\lambda\right)\right] \\
& +\sum_{j}\left[G_{22,} \bar{A}_{22 j 2} \cos \left(\omega_{22 j \partial} t+\beta_{22 j y}+2 \lambda\right)\right. \\
& \left.+G_{22 \lambda} \bar{A}_{22 j \odot} \cos \left(\omega_{22 j \odot} t+\beta_{22 j \odot}+2 \lambda\right)\right] \\
& +\sum_{j} G_{308} \bar{A}_{30 j \nu} \sin \left(\omega_{30 j_{2}} t+\beta_{30 j \otimes}\right) \\
& +\sum_{j} G_{31 \geqslant} \bar{A}_{31 j \otimes} \cos \left(\omega_{31 j \lambda} \mathbf{t}+\beta_{31 j \geqslant}+\lambda\right) \\
& +\sum_{j} G_{32 \Omega} \bar{A}_{32 j 9} \sin \left(\omega_{32 j \rho} t+\beta_{32 j \partial}+2 \lambda\right) \\
& +\sum_{j} G_{332} \bar{A}_{33 j 2} \cos \left(\omega_{33 j 2} t+\beta_{33 j 2}+3 \lambda\right) \tag{A.1}
\end{align*}
$$

Solar terms of degree 3 and all terms of degree 4 and higher are neglected because they are small. The coefficients denoted by $\overline{\mathrm{A}}_{\mathrm{nmjd}}$ are the numbers that appear in Doodson's tables [1922, pp. 322-325]. The spherical harmonic degree and order are denoted respectively by $n$ and $m$. Individual tabular entries are denoted by the index $j$. The subscript $d$ stands for "disturbing body" which is either the moon, ${ }^{3}$, or the sun, $\odot$.

The geodetic coefficients $G_{n m d}$ are functions of the latitude $\phi$ and Doodson's tidal parameter $G_{D d}$.

$$
\begin{align*}
& G_{20 d}=\frac{1}{2} G_{D d}\left(1-3 \sin ^{2} \phi\right)  \tag{A.2}\\
& G_{2 d d}=G_{D d} \sin 2 \phi  \tag{A.3}\\
& G_{22 d}=G_{D d} \cos ^{2} \lambda  \tag{A.4}\\
& G_{30 d}=1.11803 G_{D d} \sin \phi\left(3-5 \sin ^{2} \phi\right)  \tag{A.5}\\
& G_{31 d}=0.72618 G_{D d} \cos \phi\left(1-5 \sin ^{2} \phi\right)  \tag{A.6}\\
& G_{32 d}=2.59808 G_{D d} \sin \phi \cos ^{2} \phi  \tag{A.7}\\
& G_{33 d}=G_{D d} \cos ^{3} \phi \tag{A.8}
\end{align*}
$$

where

$$
\begin{align*}
& \qquad G_{D d}=\frac{3}{4} \frac{\mathrm{~m}_{\mathrm{d}} \mathrm{Gr}^{2}}{\mathrm{c}_{\mathrm{d}}^{3}}  \tag{A.9}\\
& \mathrm{~m}_{\mathrm{d}}
\end{aligned}=\text { mass of disturbing body } \quad \begin{aligned}
\mathrm{c}_{\mathrm{d}} & =\text { mean distance of disturbing body } \\
\mathrm{G} & =\text { universal gravitational constant } \\
\mathrm{r} & =\text { geocentric radius }
\end{align*}
$$

In order to consolidate his results, Doodson adjusts each solar coefficient $\bar{A}_{n m j \odot}$ so that it gives the correct result when multiplied by the corresponding lunar geodetic coefficient $\mathrm{G}_{\mathrm{nm}}^{\mathrm{nmj}}$. This adjustment involves the choice of specific values for the ratio of lunar to solar mass and the ratio of lunar to solar mean distance. If $\overline{\bar{A}}_{\mathrm{nnj}}$ 。 denotes the solar coefficient before adjustment, then

$$
\begin{equation*}
G_{n m j} \bar{A}_{n m j \odot}=G_{n m \odot} \overline{\bar{A}}_{n m j \odot} \tag{A.10}
\end{equation*}
$$

By first dividing out Doodson's numerical factor so as to recover the $\overline{\bar{A}}_{n m j \odot}$, it is possible to incorporate revised values of the lunar and solar masses and mean distances into subsequent calculations. The only solar terms in Doodson's table are those of degree 2, and for these

$$
\begin{equation*}
\frac{G_{2 m \odot}}{G_{2 m \rho}}=\frac{G_{D \odot}}{G_{D s}}=\left(\frac{m_{\odot}}{m_{3}}\right)\left(\frac{c_{9}}{c_{\odot}}\right)^{3} \tag{A.11}
\end{equation*}
$$

Doodson [1922, p. 318]-used

$$
\begin{equation*}
\frac{G_{2 \mathrm{~m} \odot}}{G_{2 \mathrm{~m}}}=0.46040 \tag{A.12}
\end{equation*}
$$

The solar coefficients $\overline{\bar{A}}_{2 \mathrm{mj} \odot}$ for use with solar geodetic coefficients are therefore given in terms of Doodson's tabulated coefficients, $\overline{\mathrm{A}}_{2 \mathrm{mj} \odot}$, by

$$
\begin{equation*}
\overline{\bar{A}}_{2 \mathrm{mj} \odot}=\frac{\overline{\mathrm{A}}_{2 \mathrm{mj} \odot}}{0.46040} \tag{A.13}
\end{equation*}
$$

The terms of degree 2 and order 1 in the tidal potential are the only ones that enter into the torque components (4.19) and (4.20). The corresponding geodetic coefficient (A.3) is written in terms the Legendre function $P_{2}^{1}(\sin \phi)$ as

$$
\begin{equation*}
G_{21 d}=\frac{1}{2} m_{d} \frac{G^{2}}{c_{d}^{3}} P_{2}^{1}(\sin \phi) \tag{A.14}
\end{equation*}
$$

The arguments of the trigonometric functions appearing in (A.1) are given by linear combinations of the following six standard variables chosen by Doodson [1922, p. 310]

$$
\begin{aligned}
\tau & =\text { local mean lunar hour angle referred to lower transit } \\
\mathrm{s} & =\text { lunar mean longitude } \\
\mathrm{h} & =\text { solar mean longitude } \\
\mathrm{p} & =\text { longitude of lunar perigee } \\
-\mathrm{N}^{\prime}=\mathrm{N} & =\text { longitude of the lunar ascending node } \\
\mathrm{p}_{\mathrm{s}} & =\text { longitude of solar perigee }
\end{aligned}
$$

The standard variables are discussed in greater detail in Appendix B. Doodson's arguments are of the form

$$
\begin{align*}
d_{1} \tau & +\left(d_{2}-5\right) s+\left(d_{3}-5\right) h \\
& +\left(d_{4}-5\right) p+\left(d_{5}-5\right) N^{\prime}+\left(d_{6}-5\right) p_{s} \tag{A.15}
\end{align*}
$$

where $d_{1}$ through $d_{6}$ are the integers in a code number written as

$$
\mathrm{d}_{1} \mathrm{~d}_{2} \mathrm{~d}_{3} \cdot \mathrm{~d}_{4} \mathrm{~d}_{5} \mathrm{~d}_{6}
$$

The integer $d_{1}$ is always equal to the order $m$ of the harmonic in which the argument arises, or equivalently, to Doodson's "schedule number" $0,1,2$, or 3 . The standard variables are given in terms of Brown's fundamental arguments by Equations (B.1) through (B.5). When terms involving second and higher powers of the time are neglected in Brown's arguments, the standard variables are expressible as

$$
\begin{equation*}
W_{i} t+P_{i} \tag{A.16}
\end{equation*}
$$

where $t$ is the Greenwich mean solar time, $W_{i}$ is a frequency and $P_{i}$ is a phase angle. In the expression for the local mean lunar time $\tau$, the longitude $\lambda$ is separated from the rest of the phase angle so that

$$
\begin{equation*}
\tau=W_{\tau} t+P_{\tau}+\lambda \tag{A.17}
\end{equation*}
$$

Doodson's arguments may be written as

$$
\begin{equation*}
m\left(W_{\tau} t+P_{\tau}\right)+m \lambda+\sum_{i=2}^{6}\left(d_{i}-5\right)\left(W_{i} t+P_{i}\right) \tag{A.18}
\end{equation*}
$$

The frequency and phase of an argument are defined by

$$
\begin{align*}
& \omega_{n m j d}=m W_{\tau}+\sum_{i=2}^{6}\left(d_{i}-5\right) W_{i}  \tag{A.19}\\
& \beta_{n m j d}=m P_{\tau}+\sum_{i=2}^{6}\left(d_{i}-5\right) P_{i} \tag{A.20}
\end{align*}
$$

and the arguments are written in the form

$$
\begin{equation*}
\omega_{\mathrm{nmj} \mathrm{~d}} \mathrm{t}+\beta_{\mathrm{nm} j \mathrm{~d}}+\mathrm{m} \lambda \tag{A.21}
\end{equation*}
$$

which appears in (A.1).
For the analysis of polar motion and tidal deformation, the tide generating potential (A.1) is rewritten as

$$
\begin{align*}
\mathrm{U}= & \sum_{\mathrm{d}} \frac{G m_{d}}{c_{d}} \sum_{n=2}^{\infty} \sum_{m=0}^{n}\left(\frac{r}{c_{d}}\right)^{n} P_{n}^{m}(\sin \phi) \\
& \cdot \sum_{j} A_{n m j d} \cos \left[\omega_{n m j d} t+\beta_{n m j d}+m \lambda+(n-m) \frac{\pi}{2}\right] \tag{A.22}
\end{align*}
$$

The tesseral diurnal coefficients $A_{21 j d}$ in (A.21) are related to those in (A.1) by

$$
\begin{equation*}
A_{21 \mathrm{jd}}=-\frac{1}{2} \bar{A}_{21 \mathrm{jd}} \tag{A.23}
\end{equation*}
$$

## APPENDIX B

DEFINITIONS OF THE STANDARD TIDAL VARIABLES

Explicit relationships are presented here between Doodson's standard variables, the similar variables defined by Melchior [1966, p. 26] and Brown's fundamental arguments as given in the Explanatory Supplement to the AENA [1961, p. 44$]$. The astronomical variables $s, h, p, N^{\top}$, and $p_{s}$ are defined in Table B. 1 and the time variables T, $\phi_{\mathrm{M}}$ and t are defined in Table B.2.

Table B. 1
Astronomical Variables

| Definition | Symbols Used by <br> Doodson [1922, p. 310] and Melchior [1966, p. 26] | Symbols Used in the Expl. Supp. [1961, p. 107] |
| :---: | :---: | :---: |
| Iunar mean longitude, measured in the ecliptic from the mean equinox of date to the mean ascending node of the lunar orbit, and then along the orbit. <br> solar mean longitude, measured in the ecliptic from the mean equinox of date <br> mean longitude of lunar perigee, measured in the ecliptic from the mean equinox of date to the mean ascending node of the lunar orbit, and then along the orbit <br> longitude of the mean ascending node of the lunar orbit on the ecliptic, measured from the mean equinox of date <br> mean longitude of solar perigee, measured in the ecliptic from the mean equinox of date | h <br> p $\mathrm{N}=-\mathrm{N}^{\dagger}$ $p_{s}$ | L <br> $\Gamma^{\prime}$ <br> $\Omega$ <br> $\Gamma$ |

Table B2
Time Variables

| Definition | This <br> Paper | Doodson <br> [1922, p. 310] | Melchior <br> $[1966, \mathrm{p} .26]$ |
| :--- | :---: | :---: | :---: |
| Greenwich mean sidereal <br> hour angle <br> local mean lunar hour <br> angle, measured from <br> lower transit of the mean <br> moon past the local <br> meridian | $\phi_{\text {M }}$ | $\tau$ | $\theta$ |
| Greenwich mean lunar <br> hour angle, measured <br> from upper transit of <br> the mean moon past the <br> Greenwich meridian |  | $\tau$ |  |
| Greenwich mean solar <br> hour angle, measured <br> from lower transit of the <br> mean sun past the Green- <br> wich meridian |  |  |  |
| Greenwich mean solar <br> hour angle, measured <br> from upper transit of <br> the mean sun past the <br> Greenwich meridian <br> Greenwich mean solar <br> time |  | t |  |
| Greenwich mean solar <br> hour angle |  |  |  |

Doodson's standard variables are given in terms of Brown's fundamental arguments, $l, l^{\prime}, F, D$, and $\Omega$ by

$$
\begin{align*}
& \mathrm{s}=\mathrm{F}+\Omega  \tag{B.1}\\
& \mathrm{h}=\mathrm{F}-\mathrm{D}+\Omega  \tag{B.2}\\
& \mathrm{p}=-\ell+\mathrm{F}+\Omega  \tag{B.3}\\
& \mathrm{N}^{\prime}=-\Omega  \tag{B.4}\\
& \mathrm{p}_{\mathrm{s}}=-l^{\prime}+\mathrm{F}-\mathrm{D}+\Omega \tag{B.5}
\end{align*}
$$

and the fundamental arguments are given as polynomials in time in the Explanatory Supplement to the AENA [1961, p. 44].

Doodson's mean lunar time is reckoned from lower transit of the moon just as conventional mean solar time is reckoned from lower transit of the mean sun. The relationship between $\phi_{M}, \tau$, and $t_{s}$ is shown in Figure B.1.

$$
\begin{align*}
& \phi_{M}=\tau+s-\pi-\lambda  \tag{B.6}\\
& \phi_{M}=t_{s}+\mathrm{h}-\pi \tag{B.7}
\end{align*}
$$

From equations (B.6) and (B.7),

$$
\begin{equation*}
\tau=\mathrm{t}_{\mathrm{s}}+\mathrm{h}-\mathrm{s}+\lambda \tag{B.8}
\end{equation*}
$$

Melchior's mean lunar and solar hour angles $\tau$ and t are each referred to the Greenwich meridian and measured from upper transit. Thus Melchior [1966, p. 26] writes

$$
\begin{align*}
& \theta=\tau+\mathrm{s}  \tag{B.9}\\
& \theta=\mathrm{t}+\mathrm{h} \tag{B.10}
\end{align*}
$$

which have the same meaning as (B.6) and (B.7).
Doodson's standard variables are each expressible in terms of the Greenwich mean sidereal hour angle $\phi_{M}$. Equating the forms (A.15) and (A.21) gives

$$
\begin{align*}
\omega_{n m j d} t & +\beta_{n m j d}+m \lambda=d_{1} \tau+\left(d_{2}-5\right) s+\left(d_{3}-5\right) h \\
& +\left(d_{4}-5\right) p+\left(d_{5}-5\right) N^{\prime}+\left(d_{6}-5\right) p_{s} \tag{B.11}
\end{align*}
$$



Figure B.1. Time Variables

Substituting for $\tau$ from (B.6) into (B.11), with $n=2$ and $m=d_{1}=1$ gives

$$
\begin{equation*}
\omega_{\mathrm{j}} \mathrm{t}+\beta_{\mathrm{j}}=\left(\phi_{\mathrm{M}}+\pi\right)+\alpha_{\mathrm{j}} \tag{B.12}
\end{equation*}
$$

where the subscripts 2,1 , and $d$ are omitted and

$$
\begin{align*}
a_{j}= & \left(d_{2}-6\right) s+\left(d_{3}-5\right) h+\left(d_{4}-5\right) p \\
& +\left(d_{5}-5\right) N^{\prime}+\left(d_{6}-5\right) p_{s} \tag{B.13}
\end{align*}
$$

The frequency of a particular term is written in terms of the Earth's constant nominal rotation rate $\Omega$ [not the $\Omega$ in Table B. 1 and Equations (B.1) through (B.5)] as

$$
\omega_{\mathrm{j}}=\Omega-\mathrm{n}_{\mathrm{j}}
$$

where

$$
\begin{equation*}
n_{j}=\Omega-\dot{\phi}_{M}-\dot{\alpha}_{j} \tag{B.14}
\end{equation*}
$$

## APPENDIX C

## EULER ANGLES RELATED TO THE MEAN SIDEREAL TIME

In order to apply Doodson's expansion of the tidal potential to problems in polar motion dynamics it is necessary to relate the Greenwich mean sidereal hour angle $\phi_{M}$ to the Euler angle $\phi$ defined in Figure 3.1. The difference ( $\phi_{M}-\phi$ ) is small and is due to polar motion, lunisolar precession and nutation, and planetary precession.

As shown in Figure C.1, the Euler angle $\phi$ is measured in the terrestrial equator from its descending node $\Upsilon_{1 E}$ on the fixed ecliptic. Polar motion causes $\phi$ to depart from the angle $\phi_{\mathrm{r}}$ measured in the true equator of date. The difference $\left(\phi_{\mathrm{r}}-\phi\right)$ is expressed in terms of the phase $\Gamma$ and amplitude $\gamma$ of the polar motion, shown in Figure C.2. Solution of the spherical triangle shown in Figure C. 3 gives

$$
\begin{equation*}
\phi_{\mathrm{r}}-\phi=\gamma \cos (\Gamma+\phi) \cot \epsilon_{1 \mathrm{~T}} \tag{C.1}
\end{equation*}
$$

in which second order terms in $\gamma$ are neglected.
The true equator of date moves relative to the mean equator of date because of the lunisolar nutation. The nutation $\Delta \psi_{\mathrm{rD}}$ in longitude is responsible for the difference between mean and apparent sidereal time. From Figure C.1,

$$
\begin{equation*}
\text { GAST } 1=\phi_{\mathrm{M}}+\Delta \psi_{\mathrm{TD}} \cos \epsilon_{\mathrm{TD}} \tag{C.2}
\end{equation*}
$$

where second order terms in $\Delta \psi_{r D}$ are neglected.
Planetary precession of the mean ecliptic causes the true equinox $\gamma_{T}$ to move relative to the equinox $\boldsymbol{r}_{1 \mathrm{~T}}$ on the fixed ecliptic. From Figure C.1, the angle $\phi_{\mathrm{r}}$ is related to the Greenwich apparent sidereal time by

$$
\begin{equation*}
\phi_{\mathrm{r}}=\mathrm{GAST} 1+\mathrm{a} \tag{C.3}
\end{equation*}
$$

where second order terms in $\Delta \psi_{r D}$ are neglected. The second order terms in $\Delta \psi_{\mathrm{rD}}$ appear because a is measured in the mean equator of date rather than in the true equator of date.


Figure C.1. Relationship Between the Greenwich Mean Sidereal Time $\phi_{M}$ and the Euler Angle $\phi$


Figure C.2. Amplitude and Phase of the Polar Motion

Equations (C.1), (C.2) and (C.3) are combined to give the difference ( $\phi_{M}-\phi$ ) as

$$
\begin{align*}
\phi_{\mathrm{M}}-\phi= & -\Delta \psi_{\mathrm{rD}} \cos \epsilon_{\mathrm{TD}}-\mathrm{a} \\
& +\gamma \cos (\Gamma+\phi) \cot \epsilon_{1 T} \tag{C.4}
\end{align*}
$$

On the right hand side of (C.3) the replacement of $\epsilon_{\mathrm{TD}}$ and $\epsilon_{1 \mathrm{~T}}$ by $\epsilon_{\mathrm{M}}$ and the replacement of $\phi$ by $\phi_{\mathrm{M}}$ introduces only additional second order terms. Therefore, with second order terms neglected,

$$
\begin{align*}
\phi_{M}-\phi= & -\Delta \psi_{r D} \cos \epsilon_{M}-a \\
& +\gamma \cos \left(\Gamma+\phi_{M}\right) \cot \epsilon_{M} \tag{C.5}
\end{align*}
$$



Figure C.3. Departure of $\phi$ From $\phi_{r}$ Due to Polar Motion

## APPENDIX D

## GEOPOTENTIAL COEFFICIENTS IN TERMS OF THE EARTH'S INERTIA TENSOR

The Earth's gravitational potential is expanded in the form

$$
\begin{align*}
V= & \frac{G m_{E}}{r}\left\{1-\sum_{n=2}^{\infty} J_{n}\left(\frac{a_{E}}{r}\right)^{n} P_{n}(\sin \phi)\right. \\
& +\sum_{n=2}^{\infty} \sum_{m=1}^{n}\left(\frac{a_{E}}{r}\right)^{n} P_{n}^{m}(\sin \phi)\left(C_{n m} \cos m \lambda+S_{n m} \sin m \lambda\right) \tag{D.1}
\end{align*}
$$

The spherical harmonic coefficients are given by

$$
\begin{align*}
& J_{n}=-\frac{1}{m_{E} a_{E}^{n}} \int r^{n} P_{n}(\sin \phi) d m  \tag{D.2}\\
& \left\{\begin{array}{l}
C_{n m} \\
S_{n m}
\end{array}\right\}=\frac{W_{n m}}{m_{E} a_{E}^{n}} \int r^{n} P_{n}^{m}(\sin \phi)\left\{\begin{array}{l}
\cos m \lambda \\
\sin m \lambda
\end{array}\right\} \tag{D.3}
\end{align*}
$$

where

$$
\begin{equation*}
W_{\mathrm{nm}}=\frac{2(\mathrm{n}-\mathrm{m})!}{(\mathrm{n}+\mathrm{m})!} \quad(\mathrm{m}=1,2, \ldots, n) \tag{D.4}
\end{equation*}
$$

The origin of the $x, y, z$ system is at the Earth's center of mass so that

$$
\begin{equation*}
J_{1}=C_{11}=S_{11}=0 \tag{D.5}
\end{equation*}
$$

$J_{2}$ is given by

$$
\begin{equation*}
J_{2}=-\frac{1}{m_{E} a_{E}^{2}} \int r^{2}\left(\frac{3}{2} \sin ^{2} \phi-\frac{1}{2}\right) d m \tag{D.6}
\end{equation*}
$$

Substituting for $x$ and $\phi$ in terms of $x, y$ and $z$ gives

$$
\begin{equation*}
J_{2}=-\frac{1}{2 m_{E} a_{E}^{2}} \int\left(2 z^{2}-x^{2}-y^{2}\right) d m \tag{D.7}
\end{equation*}
$$

Equation (D.7) is written in terms of the principal moments of inertia defined in Equation (2.3) as

$$
\begin{equation*}
J_{2}=-\frac{1}{2 m_{E} \mathrm{a}_{E}^{2}}\left(\mathrm{I}_{11}+\mathrm{I}_{22}-2 \mathrm{I}_{33}\right) \tag{D.8}
\end{equation*}
$$

Similar developments give the other second degree geopotential coefficients in terms of inertial integrals as

$$
\begin{gather*}
C_{21}=-\frac{1}{m_{E} a_{E}^{2}} I_{13}  \tag{D.9}\\
S_{21}=-\frac{1}{m_{E} a_{E}^{2}} I_{23}  \tag{D.10}\\
C_{22}=\frac{1}{4 m_{E} a_{E}^{2}}\left(I_{22}-I_{11}\right)  \tag{D.11}\\
S_{22}=-\frac{1}{2 m_{E} a_{E}^{2}} I_{12} \tag{D.12}
\end{gather*}
$$

Equations (D.8) through (D.12) are rewritten in terms of the inertia tensor perturbations defined in Equation (2.5) as

$$
\begin{gather*}
J_{2}=\frac{\mathrm{C}-\mathrm{A}}{\mathrm{~m}_{\mathrm{E}} a_{E}^{2}}+\frac{2 \mathrm{c}_{33}-\mathrm{c}_{11}-\mathrm{c}_{22}}{2 \mathrm{~m}_{E} a_{E}^{2}}  \tag{D.13}\\
\mathrm{C}_{21}=-\frac{c_{13}}{m_{E} a_{E}^{2}} \tag{D.14}
\end{gather*}
$$

$$
\begin{align*}
& S_{21}=-\frac{c_{23}}{m_{E} a_{E}^{2}}  \tag{D.15}\\
& C_{22}=\frac{c_{22}-c_{11}}{4 m_{E} a_{E}^{2}}  \tag{D.16}\\
& S_{22}=-\frac{c_{12}}{2 m_{E} a_{E}^{2}} \tag{D.17}
\end{align*}
$$

The trace of the inertia tensor (2.5) is invariant under a small deformation. In order to show this, the trace is written as

$$
\begin{equation*}
\sum_{j=1}^{3} I_{j j}=2 \int r^{2} d m \tag{D.18}
\end{equation*}
$$

The mass redistribution resulting from a small deformation is treated as a surface layer as shown in Figure D.1. The mass element dm is written in terms of a surface density $\rho(\theta, \lambda)$ as

$$
\begin{equation*}
\mathrm{dm}=\rho(\phi, \lambda) \mathrm{ds} \tag{D.19}
\end{equation*}
$$

where ds denotes a surface element. The density is expanded in spherical surface harmonics

$$
\begin{equation*}
\rho(\phi, \lambda)=\sum_{n=0}^{\infty} \mathbf{S}_{\mathbf{n}} \tag{D.20}
\end{equation*}
$$

where

$$
\begin{equation*}
S_{0}=0 \tag{D.21}
\end{equation*}
$$

in order that mass be conserved. If $V_{0}$ and $S$ denote the volume and surface corresponding to the sphere of radius $a_{E}$, then

$$
\begin{equation*}
\int r^{2} d m=\int_{V_{0}} r^{2} d m+\int_{S} a_{E}^{2} \rho(\phi, \lambda) d s \tag{D.22}
\end{equation*}
$$

Substituting for $\rho$ from (D.20) gives

$$
\begin{equation*}
\int_{S} a_{E}^{2} \rho(\phi, \lambda) d s=4 \pi a_{E}^{2} S_{0}=0 \tag{D.23}
\end{equation*}
$$

so that

$$
\begin{equation*}
\int \mathrm{r}^{2} \mathrm{dm}=\int_{\mathrm{v}_{0}} \mathrm{r}^{2} \mathrm{dm} \tag{D.24}
\end{equation*}
$$

Therefore

$$
\begin{equation*}
\left(A+c_{11}\right)+\left(A+C_{22}\right)+\left(C+c_{33}\right)=2 A+C \tag{D.25}
\end{equation*}
$$

from which

$$
\begin{equation*}
c_{11}+c_{22}+c_{33}=0 \tag{D.26}
\end{equation*}
$$



Figure D.1. Small Deformation Represented as a Surface Layer

## APPENDIX E

DIRECTION COSINES OF THE AXIS OF FIGURE RELATED TO PERTURBATIONS IN THE INERTIA TENSOR

The Earth's inertia tensor is written in the form

$$
I=\left[\begin{array}{ccc}
A+c_{11} & c_{12} & c_{13}  \tag{E.1}\\
c_{12} & A+c_{22} & c_{23} \\
c_{13} & c_{23} & C+c_{33}
\end{array}\right]
$$

given by Equation (2.5). The perturbations $c_{i j}$ are small in relation to the principal moments of inertia so that the xyz system of Figure 2.1 is almost a principal axis system. The xyz system must be rotated through the small angles $\ell$ and $p$ shown in Figure E. 1 in order to make the $z^{\prime}$ axis coincide with the principal axis of inertia.


Figure E.1. Rotation of the terrestrial System to Make the $\mathbf{z}$ axis Coincide With the Principa! Axis of Inertia

Neglecting second order terms in $l$ and $p$, the coordinate transformation from xyz to $x^{\prime} y^{\prime} z^{\prime}$ is

$$
R=\left[\begin{array}{rrr}
1 & 0 & -\ell  \tag{E.2}\\
0 & 1 & -p \\
\ell & p & \underline{0}
\end{array}\right]
$$

The inertia tensor transforms as

$$
\begin{equation*}
I^{\prime}=R I R^{T} \tag{E.3}
\end{equation*}
$$

The transformed inertia tensor is

$$
I^{\prime}=\left[\begin{array}{ccc}
A+c_{11} & c_{12} & {\left[\ell(A-C)+c_{13}\right]}  \tag{E.4}\\
c_{12} & A+c_{22} & {\left[p(A-C)+c_{23}\right]} \\
{\left[\ell(A-C)+c_{13}\right]} & {\left[p(A-C)+c_{23}\right]} & C+c_{33}
\end{array}\right]
$$

The direction cosines

$$
\begin{align*}
& l=\frac{c_{13}}{C-A}  \tag{E.5}\\
& p=\frac{c_{23}}{C-A} \tag{E.6}
\end{align*}
$$

will diagonalize $I^{\prime}$ except for the $c_{12}$ terms. If $c_{12} \neq 0$ an additional rotation about $z^{\prime}$ is necessary to produce a principal axis system.

The direction cosines of the principal axis of ineria are combined into a complex number $\psi_{f}$ called the axis of figure

$$
\begin{equation*}
\psi_{f}=\ell+i p \tag{E.7}
\end{equation*}
$$

In terms of the complex representation,

$$
\begin{equation*}
c=c_{13}+i c_{23} \tag{E.8}
\end{equation*}
$$

of the products of inertia, the axis of figure is given by

$$
\begin{equation*}
\psi_{f}=\frac{c}{C-A} \tag{E.9}
\end{equation*}
$$

## APPENDIX $\mathbf{F}$ <br> SYMMETRIC TIDAL ARGUMENTS

Summations of the form

$$
\begin{equation*}
\sum_{j} \widetilde{A}_{j} e^{-i\left(\omega_{j} t+\beta_{j}-\phi\right)} \tag{F.1}
\end{equation*}
$$

arise in deriving the formulas for the Euler angle perturbations in Sections 6 and 7. The $\widetilde{A}_{j}$ denote general coefficients, $\omega_{j} t+\beta_{j}$. is a tidal argument and $\phi$ is the Euler angle defined in Figure 3.1 which describes the Earth's diurnal rotation. Two tidal arguments with distinct indices $j_{+}$and $j_{-}$are called symmetric when

$$
\begin{align*}
& \omega_{j+} t+\beta_{j+}=\left(\phi_{M}+\pi\right)+\alpha_{j}  \tag{F.2}\\
& \omega_{j_{-}} t+\beta_{j^{-}}=\left(\phi_{M}+\pi\right)-\alpha_{j} \tag{F.3}
\end{align*}
$$

where $\alpha_{j}$ denotes a linear combination of Doodson's arguments and $\phi_{M}$ is the Greenwich mean sidereal hour angle.

The arguments of symmetric terms in (F.1) are

$$
\begin{equation*}
\omega_{j \pm} \mathrm{t}+\beta_{\mathrm{j} \pm}-\phi=\left(\phi_{\mathrm{M}}-\phi\right)+\pi \pm \alpha_{\mathrm{j}} \tag{F.4}
\end{equation*}
$$

Two symmetric terms are combined to form

$$
\begin{align*}
& \widetilde{A}_{j+} e^{-i\left(\omega_{j+} t+\beta_{j \pm}-\phi\right)}+\widetilde{A}_{j-} e^{-i\left(\omega_{j-}-t+\beta_{j-}-\phi\right)} \\
&= {\left[-\cos \left(\phi_{M}-\phi\right) \cos \alpha_{j}\left(\widetilde{A}_{j+}+\widetilde{A}_{j-}\right)\right.} \\
&\left.-\sin \left(\phi_{M}-\phi\right) \sin \alpha_{j}\left(-\widetilde{A}_{j+}+\widetilde{A}_{j-}\right)\right] \\
&++i\left[-\cos \left(\phi_{M}-\phi\right) \sin \alpha_{j}\left(-\widetilde{A}_{j+}+\widetilde{A}_{j-}\right)\right. \\
&\left.+\sin \left(\phi_{M}-\phi\right) \cos \alpha_{j}\left(\widetilde{A}_{j+}+\widetilde{A}_{j-}\right)\right] \tag{F.5}
\end{align*}
$$

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